



The Most Massive galaxy Clusters (M2C) across cosmic time: link between radial total mass distribution and dynamical state

I. Bartalucci, M. Arnaud, G. W. Pratt, J. Démoclès, L. Lovisari

► To cite this version:

I. Bartalucci, M. Arnaud, G. W. Pratt, J. Démoclès, L. Lovisari. The Most Massive galaxy Clusters (M2C) across cosmic time: link between radial total mass distribution and dynamical state. *Astronomy and Astrophysics - A&A*, 2019, 628, pp.A86. 10.1051/0004-6361/201935984 . cea-02266323

HAL Id: cea-02266323

<https://cea.hal.science/cea-02266323>

Submitted on 13 Aug 2019

HAL is a multi-disciplinary open access archive for the deposit and dissemination of scientific research documents, whether they are published or not. The documents may come from teaching and research institutions in France or abroad, or from public or private research centers.

L'archive ouverte pluridisciplinaire **HAL**, est destinée au dépôt et à la diffusion de documents scientifiques de niveau recherche, publiés ou non, émanant des établissements d'enseignement et de recherche français ou étrangers, des laboratoires publics ou privés.

The Most Massive galaxy Clusters (M2C) across cosmic time: link between radial total mass distribution and dynamical state

I. Bartalucci¹, M. Arnaud¹, G. W. Pratt¹, J. Démoclès¹, and L. Lovisari²

¹ AIM, CEA, CNRS, Université Paris-Saclay, Université Paris Diderot, Sorbonne Paris Cité, 91191 Gif-sur-Yvette, France
 e-mail: iacopo.bartalucci@cea.fr

² Center for Astrophysics | Harvard & Smithsonian, 60 Garden Street, Cambridge, MA 02138, USA

Received 29 May 2019 / Accepted 3 July 2019

ABSTRACT

We study the dynamical state and the integrated total mass profiles of 75 massive ($M_{500} > 5 \times 10^{14} M_{\odot}$) Sunyaev–Zeldovich (SZ)-selected clusters at $0.08 < z < 1.1$. The sample is built from the *Planck* catalogue, with the addition of four SPT clusters at $z > 0.9$. Using *XMM-Newton* imaging observations, we characterise the dynamical state with the centroid shift $\langle w \rangle$, the concentration C_{SB} , and their combination, M , which simultaneously probes the core and the large-scale gas morphology. Using spatially resolved spectroscopy and assuming hydrostatic equilibrium, we derive the total integrated mass profiles. The mass profile shape is quantified by the sparsity, that is the ratio of M_{500} to M_{2500} , the masses at density contrasts of 500 and 2500, respectively. We study the correlations between the various parameters and their dependence on redshift. We confirm that SZ-selected samples, thought to most accurately reflect the underlying cluster population, are dominated by disturbed and non-cool core objects at all redshifts. There is no significant evolution or mass dependence of either the cool core fraction or the centroid shift parameter. The M parameter evolves slightly with z , having a correlation coefficient of $\rho = -0.2 \pm 0.1$ and a null hypothesis p -value of 0.01. In the high-mass regime considered here, the sparsity evolves minimally with redshift, increasing by 10% between $z < 0.2$ and $z > 0.55$, an effect that is significant at less than 2σ . In contrast, the dependence of the sparsity on dynamical state is much stronger, increasing by a factor of $\sim 60\%$ from the one third most relaxed to the one third most disturbed objects, an effect that is significant at more than 3σ . This is the first observational evidence that the shape of the integrated total mass profile in massive clusters is principally governed by the dynamical state and is only mildly dependent on redshift. We discuss the consequences for the comparison between observations and theoretical predictions.

Key words. dark matter – galaxies: clusters: intracluster medium – X-rays: galaxies: clusters

1. Introduction

The shape of the dark matter profile in galaxy clusters is a sensitive test of the nature of dark matter and of the theoretical scenario of structure formation. In the standard framework, cosmological structures form hierarchically from initial density fluctuations that grow under the influence of gravity. In a Λ cold dark matter (Λ CDM) Universe, the dark matter (DM) collapse is scale-free and one expects objects to form with similar internal structure. The DM shape is expected to depend on the halo assembly history, which is a function of the redshift, the total mass, and the underlying cosmology (e.g. Dolag et al. 2004; Kravtsov & Borgani 2012). A certain scatter of the shape and break of strict self-similarity are expected, reflecting the detailed formation history of each halo.

Clusters of galaxies are ideal targets to test the above scenario: the dark matter is the dominant component by far except in the very centre. Furthermore, complementary techniques can measure the total mass density profile, for example, galaxy velocities; strong gravitational lensing in the centre and weak lensing at large scale; X-ray estimates using the intra-cluster medium (ICM) density and temperature profiles and the hydrostatic equilibrium (HE) equation (see Pratt et al. 2019, for a review).

Numerical simulations indeed predict that the cold dark matter density profiles of virialised objects follows a ‘universal’ form. Well-known parametric models for the DM density profiles include those proposed by Navarro et al. (1997; hereafter

NFW) and the Einasto profile (Einasto 1965), which is currently considered to be a more accurate description of the profiles in state-of-the-art simulations (Navarro et al. 2004; Wu et al. 2013).

A fundamental parameter of these parametric models is their concentration, which in general terms describes the relative distribution in the core as compared to the outer region (Klypin et al. 2016). The concentration is characterised in the NFW and Einasto¹ models by the ratio $c \equiv r_{-2}/R_{\Delta}^2$, where r_{-2} is the radius at which the logarithmic density slope is equal to -2 . The relation between the concentration and the total mass (hereafter $c-M$) has been very widely used as an indicator of the dark matter profile shape in cosmological simulations (e.g. Diemer & Kravtsov 2015, and references therein). However, it has been shown that the $c-M$ relation of the most relaxed haloes is different to that derived for the full population (e.g. Neto et al. 2007; Bhattacharya et al. 2013). In fact, the capability of these parametric models to reproduce the DM profile (i.e. the goodness of the fit) depends on the dynamical state of the halo and on the formation time (Jing 2000; Wu et al. 2013, for the NFW case). Results on the dependence of the DM profile shape on mass and redshift based on $c-M$ relations may thus be ambiguous, and depend on the sample selection (Klypin et al. 2016; Balmès et al. 2014). For this reason, recent works based on simulations have

¹ In this model, the actual shape also depends on a second parameter.

² R_{Δ} is defined as the radius enclosing Δ times the critical matter density at the cluster redshift. M_{Δ} is the corresponding mass.

often focused on relaxed haloes (e.g. Dutton & Macciò 2014; Ludlow et al. 2014; Correa et al. 2015).

However, a rigorous comparison with cluster observations cannot be made, as there is a continuous distribution of dynamical states, and the definition of what is a relaxed object is therefore somewhat arbitrary. Moreover, in the absence of fully realistic simulations including the complex baryonic physics, it is nearly impossible to define common criteria that quantify the dynamical state in a consistent manner, both in simulations and observations. Ideally we would compare the full cluster population from numerical simulations to observed samples chosen to reflect as closely as possible the true underlying population. The advent of Sunyaev–Zeldovich (SZ)-selected cluster catalogues offers a unique opportunity to build such observational samples. Surveys such as those from the Atacama Cosmology Telescope (ACT; Marriage et al. 2011), the South Pole Telescope (SPT; Reichardt et al. 2013; Bleem et al. 2015), and the *Planck* Surveyor (Planck Collaboration VIII 2011; Planck Collaboration XXIX 2014; Planck Collaboration XXVII 2016) have provided SZ-selected cluster samples up to $z \sim 1.5$. The magnitude of the SZ effect being closely linked to the underlying mass with small scatter (da Silva et al. 2004), these are thought to be as near as possible to being mass selected, and as such unbiased.

The observational study of the DM profile shape using such samples requires investigation of the dependence on fundamental cluster quantities such as mass, dynamical state, and redshift. X-ray observations are an excellent tool with which to undertake such studies – the ICM morphology can be used to infer the dynamical state, while the total mass profile can be derived by applying the hydrostatic equilibrium (HE) equation to spatially resolved density and temperature profiles. While this method yields the highest statistical precision on individual profiles over a wide radial range and up to high z (Amodeo et al. 2016; Bartalucci et al. 2018), it has the drawback of a systematic uncertainty due to any departure of the gas from HE, which must be taken into account.

Here we apply the sparsity parameter introduced by Balmès et al. (2014) to quantify the shape of the total mass profiles derived from the X-ray observations. The sparsity is defined as the ratio of the integrated mass at two overdensities. This non-parametric quantity is capable of efficiently characterising the profile shape, as long as the two overdensities are separated enough to probe the shape of the mass profile (Balmès et al. 2014; Corasaniti et al. 2018). Formally, the DM profiles can be derived by subtracting the gas and galaxy distribution from the total. However in the following we focus on the total distribution, in view of the negligible impact of the baryonic component outside the very central region on the total profile shape for the halo masses and density contrasts under consideration (e.g. Velliscig et al. 2014, Fig. 2).

In this work, we present the dynamical properties and the individual spatially resolved radial total mass profiles of a sample of 75 SZ-selected massive clusters in the $[0.08–1]$ redshift range with $M_{500} = [5–20] \times 10^{14} M_{\odot}$. We discuss the dispersion and evolution of the total mass profile shape, and demonstrate the link between the diversity in profile shape and the underlying dynamical state. In Sect. 2 we present the sample; in Sects. 3 and 4 we describe the methodology used to derive the HE mass profiles and the morphological parameters of each cluster, respectively; in Sect. 5 we discuss the morphological properties of the sample and its evolution; in Sect. 6 we investigate the dependence of the profile shape on mass, redshift, and dynamical state using the sparsity; and finally, in Sects. 7 and 8 we discuss our results and present our conclusions.

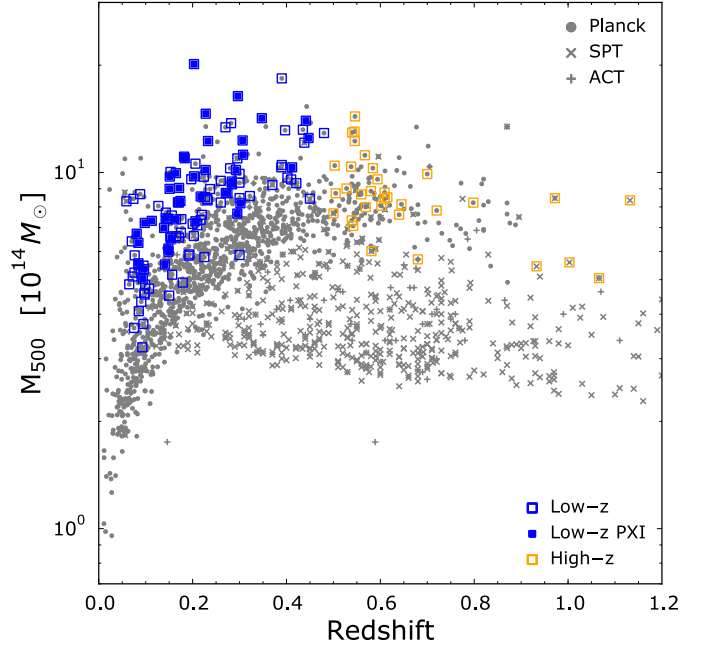


Fig. 1. Distribution in the mass–redshift plane of all the clusters published in the *Planck*, SPT, and ACT catalogues. Filled circles: *Planck* clusters with available redshifts (Planck Collaboration VIII 2011; Planck Collaboration XXIX 2014; Planck Collaboration XXVII 2016); crosses: SPT (Bleem et al. 2015); plus symbols: ACT (Hasselfield et al. 2013). Masses in the *Planck* catalogue are derived iteratively from the M_{500} – Y_{SZ} relation calibrated using hydrostatic masses from *XMM-Newton*; they are not corrected for the hydrostatic equilibrium (HE) bias. In the figure the *Planck* masses are multiplied by a factor of 1.2 (i.e. assuming 20% bias). The blue and orange open squares identify the Low- z and High- z samples considered in this study (see Sect. 2). The clusters that are part of the Low- z PXI sample are shown as filled blue squares.

We adopt a flat Λ -cold dark matter cosmology with $\Omega_m = 0.3$, $\Omega_\Lambda = 0.7$, $H_0 = 70 \text{ km Mpc s}^{-1}$, and $h(z) = (\Omega_m(1+z)^3 + \Omega_\Lambda)^{1/2}$, where $h(z) = H(z)/H_0$ throughout. Uncertainties are given at the 68% confidence level (1σ). All fits were performed via χ^2 minimisation.

2. The sample

2.1. The high- z SZ-selected sample

Our initial sample is built from the 28 clusters with spectroscopic $0.5 < z < 0.9$ in the sample of the *XMM-Newton* Large Programme (LP) ID069366 (with re-observation of flared targets in ID72378), consisting of clusters detected at high S/N with *Planck*, and confirmed by autumn 2011 to be at $z > 0.5$. The exposure times of these observations were optimised to allow the determination of spatially resolved radial total mass profiles at least up to R_{500} . To extend the redshift coverage to $z \sim 1$ we included the five most massive clusters detected by SPT and *Planck* in the redshift range $0.9 < z < 1.1$. Using deep *XMM-Newton* and *Chandra* observations, Bartalucci et al. (2017, 2018) examined the X-ray properties of these objects and determined ICM and HE total mass profiles up to R_{500} . The resulting observation properties are detailed in Table 1 of Bartalucci et al. (2017).

Combining the above, we obtain the full ‘High- z ’ sample of 33 objects, with $M_{500}^{SZ} > 5 \times 10^{14} M_{\odot}$, shown with orange open squares in Fig. 1. The observation details of this sample are detailed in Table B.1.

2.2. The Low- z SZ selected sample

Any evolution study requires a local reference sample with similar selection and quality criteria to act as an ‘anchor’ to compare to high redshifts. The ongoing (AO17) *XMM-Newton* heritage program ‘Witnessing the culmination of structure formation in the Universe’ (PIs M. Arnaud and S. Ettori), based on the final *Planck* PSZ2 catalogue, will serve this function in the future (observations are to be completed by 2021). In the interim, for the present study we use published *XMM-Newton* follow-up of *Planck*-selected local clusters taken from the ESZ sample.

The early SZ (ESZ; [Planck Collaboration VIII 2011](#)) catalogue represents the first release derived from the *Planck* all-sky SZ survey, containing 188 clusters mostly at $z < 0.3$. There are a number of studies in the literature describing the X-ray properties of this sample. In particular, [Lovisari et al. \(2017\)](#) characterised the global properties and the morphological state of the ESZ clusters covered by *XMM-Newton* observations. We use their results on the morphological properties of the 118 ESZ clusters at $0.05 < z < 0.5$ for which R_{500} is within the field of view, and the relative error on the morphological parameters is less than 50%. This is about 80% of the corresponding parent ESZ subsample defined with the same z and size criteria, so we do not expect any major bias due to the incomplete *XMM-Newton* coverage. These objects cover one decade in mass, $M_{500} \sim [3-20] \times 10^{14} M_{\odot}$, and are shown with blue open squares in Fig. 1. Henceforth we refer to this sample as the ‘Low- z ’ sample.

The spatially resolved thermodynamic properties of an ESZ subsample were analysed by the *Planck* Collaboration, who combined X-ray and SZ data to calibrate the local scaling relations ([Planck Collaboration XI 2011](#)) and measure the pressure profiles ([Planck Collaboration Int. V 2013](#)). We use the published thermodynamic profiles of 42 clusters for which R_{500} is within the field of view (the subsample ‘A’ defined in Sect. 3.1 of [Planck Collaboration XI 2011](#)) to derive the total mass profiles. These clusters are shown as filled blue squares in Fig. 1. We henceforth refer to this sample as the ‘Low- z PXI’ sample. Its representativeness with respect to the full Low- z sample is excellent, as discussed in Appendix B.

2.3. Data preparation

The observations used in this work were taken using the European Photon Imaging Camera (EPIC; [Turner et al. 2001](#) and [Strüder et al. 2001](#)) instrument on board the *XMM-Newton* satellite. This instrument is composed of three CCD arrays, namely MOS1, MOS2, and PN, which simultaneously observe the target. Datasets were reprocessed using the Science Analysis System³ (SAS) pipeline version 15.0 and calibration files as available in December 2016. Event files with this calibration applied were produced using the *emchain* and *epchain* tools.

The reduced datasets were filtered in the standard fashion. Events for which the keyword PATTERN is < 4 and < 13 for MOS1, 2, and PN cameras, respectively, were filtered out from the analysis. Flares were removed by extracting a light curve, and removing from the analysis the time intervals where the count rate exceeded 3σ times the mean value. We created the exposure map for each camera using the SAS tool *eexmap*. We merged multiple observation datasets of the same object, if available. We report in Table B.1 the effective exposure times after all these procedures. We corrected vignetting following the weighting scheme detailed in [Arnaud et al. \(2001\)](#). The weight

for each event was computed by running the SAS *evigweight* tool on the filtered observation datasets.

Point sources were identified using the Multi-resolution wavelet software ([Starck et al. 1998](#)) on the exposure-corrected [0.3–2] keV and [2–5] keV images. We inspected each resulting list by eye to check for false detections and missed sources. We defined a circular region around each detected point source and excised these from the subsequent analysis.

We also defined regions encircling obvious sub-structures. Identified by eye, these regions were considered in the morphological analysis because the parameters we used to characterise the morphological state of a cluster (Sect. 4) are sensitive to the presence of any such sub-structure. However, these regions were excluded in the radial profile analysis detailed in Sect. 3.

X-ray observations are affected by instrumental and sky backgrounds. The former is due to the interaction of the instrument with energetic particles, while the latter is caused by Galactic thermal emission and the superimposed emission of all the unresolved point sources, namely the cosmic X-ray background ([Lumb et al. 2002](#); [Kuntz & Snowden 2000](#)). These components were estimated differently for the radial profile 1D analysis and for the morphological 2D analysis, as described in Sects. 3 and 4, respectively.

3. Radial profile analysis

3.1. Instrumental background estimation

We evaluated the instrumental background for the radial profile analysis following the procedures described in [Pratt et al. \(2010\)](#). Briefly, observations taken with the filter wheel in CLOSED position were renormalised to the source observation count rate in the [10–12] and [12–14] keV bands for the EMOS and PN cameras, respectively. We then projected these event lists in sky coordinates to match our observations. We applied the same point source masking and vignetting correction to the CLOSED event lists as for the source data. We also produced event lists to estimate the out-of-time (OOT) events using the SAS-*epchain* tool.

3.2. Density and 3D temperature profiles

To determine the radial profiles of density and temperature of the ICM we followed the same procedures and settings detailed in Sects. 3.2 and 3.3 of [Bartalucci et al. \(2017\)](#). Briefly, we firstly determined the X-ray peak by identifying the peak of the emission measured in count-rate images in the [0.3–2.5] keV band smoothed using a Gaussian kernel with a width of between 3 and 5 pixels. We extracted the vignettted-corrected and background-subtracted surface brightness profiles, S_X , from concentric annuli of width $2''$, centred on the X-ray peak, from both source and background event lists. These profiles were used to derive the radial density profiles, $n_e(r)$, employing the deprojection and PSF correction with regularisation technique described in [Croston et al. \(2006\)](#).

We obtained the deprojected temperature profiles by performing the spectral analysis described in detail in [Pratt et al. \(2010\)](#) and Sect. 3.4 of [Bartalucci et al. \(2017\)](#). The background-subtracted spectrum of a region free of cluster emission was fitted with two unabsorbed MEKAL thermal models plus an absorbed power law with fixed slope of $\Gamma = 1.4$. The resulting best-fitting model, renormalised by the ratio of the extraction areas, was then added as an extra component in each annular fit. The cluster emission was modelled by an absorbed MEKAL

³ cosmos.esa.int/web/xmm-newton

model with N_H given in Table B.1 using the absorption cross sections from Morrison & McCammon (1983). Spectral fitting was performed using XSPEC⁴ version 12.8.2. The deprojected 3D temperature profile, T_{3D} , was derived from the projected profile using the ‘Non parametric-like’ technique described in Sect. 2.3.2 of Bartalucci et al. (2018).

For the Low- z PXI sample, we used the temperature and density profiles published in Planck Collaboration Int. V (2013) and Planck Collaboration XI (2011), which were derived employing identical methods to those used in this work.

3.3. Global properties

We determined the mass at density contrast $\Delta = 500$, $M_{500}^{Y_X}$, and corresponding $R_{500}^{Y_X}$ radius, from the mass proxy Y_X . This was computed iteratively from the $M_{500}-Y_X$ relation, calibrated by Arnaud et al. (2010) using HE mass estimates of local relaxed clusters. We assumed that the $M_{500}-Y_X$ relation obeys self-similar evolution. The starting mass value was obtained from the $M-T$ relation of Arnaud et al. (2005), and the computation converges typically within 5–10 iterations. The quantity Y_X is defined as the product of the temperature measured in the $[0.15-0.75]R_{500}$ region and the gas mass within R_{500} (Kravtsov et al. 2006), the gas mass profiles being computed from the density profiles. The mass for each cluster and associated errors are reported in Table B.1.

We used the $M_{500}^{Y_X}$ published by Lovisari et al. (2017) for the Low- z sample, and the $M_{500}^{Y_X}$ computed by us for the Low- z PXI and High- z samples. We investigated the coherence between these measurements by comparing the masses for the clusters in common between the Low- z plus High- z samples and the Low- z PXI sample. The excellent agreement between the two is discussed in Appendix A, and the comparison is shown the right panel of Fig. A.1.

3.4. Derivation of the total mass profiles

Under the assumption that the ICM is in hydrostatic equilibrium in the gravitational potential well the relation between the total halo mass within radius R and the ICM thermodynamic properties is:

$$M(\leq R) = -\frac{kT(R)R}{G\mu m_H} \left[\frac{d \ln n_e(R)}{d \ln R} + \frac{d \ln T(R)}{d \ln R} \right], \quad (1)$$

where G is the gravitational constant, m_H is the hydrogen atom mass, and $\mu = 0.6$ is the mean molecular weight in atomic mass unit. We used the deprojected density and temperature profiles and the relation in Eq. (1) to derive the mass profiles, applying the ‘forward non-parametric-like technique’ detailed in Sect. 2.4.1 of Bartalucci et al. (2018). The ‘forward’ in the technique name is due to the fact that we started our analysis from the surface brightness and projected temperature observables to derive the mass profiles at the end. The ‘non-parametric-like’ refers to the fact that we used a deprojection technique to derive the density and temperature profiles, that is, without using parametric models.

The mass profiles of the Low- z PXI and High- z samples and their associated uncertainties are shown in Fig. 2 (those of the five highest-redshift clusters are reproduced as published in Bartalucci et al. 2018). The profiles of the High- z clusters are

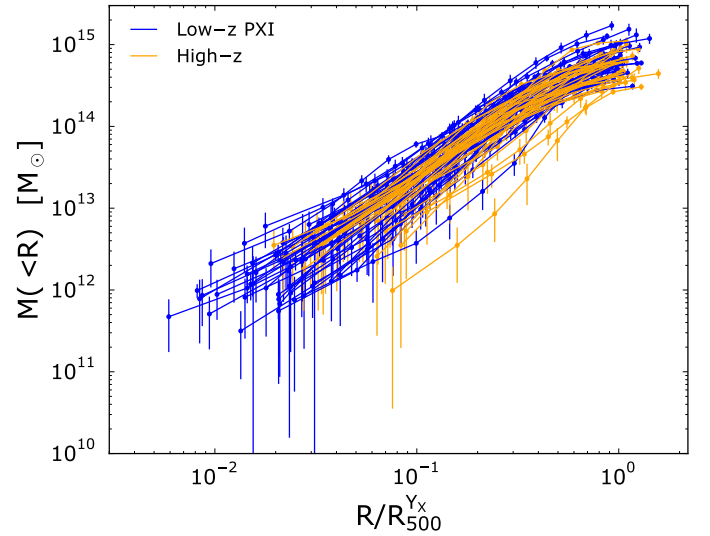


Fig. 2. Integrated mass profiles as a function of scaled radius, estimated from the hydrostatic equilibrium equation, for all clusters considered in this work. The Low- z PXI and High- z samples are plotted in blue and red, respectively.

mapped at least up to $0.8R_{500}$, with 20 out of 33 objects measured up to R_{500} , thus not requiring extrapolation to compute the HE mass at a density contrast of $\Delta = 500$, M_{500}^{HE} . The median statistical error is about 20%. The quality of the Low- z PXI mass profiles, based on archival data, is lower and less homogeneous. Only 18 out of 42 clusters have temperature profiles extending to R_{500} , with a maximum radius between $0.6R_{500}$ and $1.4R_{500}$. Bartalucci et al. (2018) showed that while the HE mass is very robust when the temperature profiles extend up to R_{500} , the M_{500}^{HE} mass is very sensitive to the mass estimation method when extrapolation is required. This is particularly the case for irregular clusters. To minimise systematic errors, we used $M_{500}^{Y_X}$ rather than M_{500}^{HE} in the following for all scaling with mass, and for the computation of the sparsity (see below). The possible impact of this choice on our results is discussed in Sect. 7.

3.5. Sparsity

The sparsity, S , was introduced by Balmès et al. (2014) to quantify the shape of the dark matter profile. It is defined as the ratio of the integrated mass at two over-densities, and has the advantage of being non-parametric, as there is no a priori assumption on the form of the profile. The sparsity therefore represents a useful measure when dealing with a population of objects with a wide variety of dynamical states. Another non-parametric approach, advocated by Klypin et al. (2016), considers the maximum circular velocity, which is linked to the traditional NFW or Einasto concentration. As the velocity is basically the square root of $M(<R)/R$, it can also be derived from observations. In practice however, measuring such a maximum is much more difficult than measuring a ratio of integrated masses.

In the following, we concentrate on the sparsity to investigate the shape of the mass profiles, which is defined as:

$$S_{\Delta_1, \Delta_2} \equiv \frac{M_{\Delta_1}}{M_{\Delta_2}}, \quad (2)$$

where M_{Δ} is the mass corresponding to the density contrast Δ and with $\Delta_1 < \Delta_2$. We recall that $M_{\Delta} = M(<R_{\Delta})$, which is the

⁴ <https://heasarc.gsfc.nasa.gov/xanadu/xspec/>

mass enclosed within R_Δ , such that

$$\frac{M(< R_\Delta)}{(4\pi/3) R_\Delta^3} = \Delta \rho_{\text{crit}}. \quad (3)$$

Balmès et al. (2014) argue that the general properties of the sparsity do not depend on the choice of $\Delta_{1,2}$ as long as the halo is well defined (i.e. Δ_1 is not too small), and that the interaction between dark and baryonic matter in the central region can be neglected (i.e. Δ_2 is not too large). We use $\Delta_1 = 500$ and $\Delta_2 = 2500$; the choice of the latter is further discussed in Sect. 6.2.

4. Morphological analysis

4.1. Centroid shift

We produced count images for each camera in the soft band, [0.3–2] keV, binned using 2'' pixel size, on which we excised and refilled the masked regions where point sources were detected using the *Chandra* interactive analysis of observation (CIAO; Fruscione et al. 2006) *dmfilth* tool. Sub-structures were not masked for this analysis. We estimated the background following a similar approach to that of Böhringer et al. (2010). We computed the background map for each camera by fitting the refilled count images using a linear combination of the vignettted and unvignettted exposure maps to account for the instrumental and sky background, respectively. We removed the cluster emission by masking a circular region within R_{500}^{X} and centred on the X-ray peak. Exposure maps and background and count images of MOS1, MOS2, and PN were combined, weighting by the ratio of integrated surface brightness profile of each individual camera to that of the combined profile. The combined count images were then background subtracted and exposure corrected. We produced 100 realisations of the count-rate maps by applying the same procedure to 100 Poisson realisations of the count maps.

The centroid shift parameter, $\langle w \rangle$, was introduced by Mohr et al. (1993) as a proxy to characterise the dynamical state of a cluster. The centroid (x_c, y_c) within an aperture is defined as

$$(x_c, y_c) \equiv \frac{1}{N_i} \sum_k n_k(x_k, y_k), \quad (4)$$

where N_i is the total number of counts per second within the i th aperture and n_k is the count rate in the pixel k of coordinates (x_k, y_k) . We computed the mean deviation of the centroid from the X-ray peak by measuring the displacement within $N = 10$ apertures using the definition of Böhringer et al. (2010):

$$\langle w \rangle = \left(\frac{1}{N-1} \sum_{i=0}^{10} (\Delta_i - \langle \Delta \rangle)^2 \right)^{1/2} \frac{1}{R_{500}^{\text{X}}}, \quad (5)$$

where Δ_i is the projected distance between the X-ray peak and the centroid computed within the i th concentric annulus, each one being $i \times 0.1 R_{500}^{\text{X}}$ in width. Uncertainties on the centroid shift were estimated by measuring $\langle w \rangle$ on the 100 Poisson count map realisations, and taking the values within 68% of the median.

Maughan et al. (2008) measured $\langle w \rangle$ on a sample of clusters observed by *Chandra*, excluding the inner 30 kpc to make the parameter less sensitive to very bright cores. The PSF of *XMM-Newton* is larger for all the clusters considered in this work and, for this reason, we did not excise the core from the analysis. The good agreement between the $\langle w \rangle$ values derived at $z > 0.9$ by *Chandra* and *XMM-Newton* shown by Bartalucci et al. (2017) indicates that the *XMM-Newton* PSF is not an issue. Nurgaliev et al. (2013) showed that the centroid shift can be biased high in

the case of observations with a low number (< 2000) of counts. In our sample the minimum number of counts in the [0.3–2] keV band we used to measure $\langle w \rangle$ is 3000. Furthermore, all clusters are in the high-SN regime, the lowest SN in our sample being 40.

Following Pratt et al. (2009), we initially classify an object as ‘morphologically disturbed’ if $\langle w \rangle > 0.01$ and ‘morphologically regular’ if $\langle w \rangle < 0.01$. The results of the centroid shift characterisation for the Low- z and High- z samples is shown in the top- and bottom-left panels of Fig. 3, and the corresponding values are given in Table B.1.

4.2. Surface brightness concentration C_{SB}

The ratio of the surface brightness profile within two concentric apertures, hereafter the C_{SB} , was introduced by Santos et al. (2008) to quantify the concentration of cluster X-ray emission. We computed the C_{SB} using the following definition:

$$C_{\text{SB}} = \frac{\int_0^{0.1 \times R_{500}^{\text{X}}} S_X(r) dr}{\int_0^{0.5 \times R_{500}^{\text{X}}} S_X(r) dr}, \quad (6)$$

where the error was computed using a Monte-Carlo procedure on 100 Gaussian realisations of the surface brightness profiles and taking the 68% value around the median. The C_{SB} parameter is a robust X-ray measurement, as it relies on the extraction of surface brightness profiles only and is not model-dependent. We nonetheless corrected for the *XMM-Newton* PSF in view of the high z and small angular size of the high z sample.

Santos et al. (2010) demonstrated that for objects at high redshift the emission within two apertures requires a different k -correction due to the presence of a cool core, that is cool cores will have typically a softer spectrum than the surrounding regions. This correction is potentially important for this study, as we are comparing C_{SB} in a wide redshift range. Santos et al. (2010) proposed a correction for this effect which requires spatially resolved temperature profiles. At the median redshift of the Low- z sample, the k correction is negligible ($< 1\%$), but it can be up to $\sim 5\%$ at the highest redshifts. We therefore did not make this correction for the Low- z sample. For the High- z sample, we applied the k -correction to the C_{SB} of all the clusters as if they were observed at the median redshift of the Low- z sample, $z = 0.19$. Henceforth, all the C_{SB} values shown and used in this work are k -corrected in this manner. The results of the C_{SB} analysis are reported in the top and bottom central panels of Fig. 3 and in Table B.1.

The C_{SB} allows the identification of cool-core (hereafter CC) clusters, the parameter being tightly correlated with the cooling time (e.g. Croston et al. 2008; Santos et al. 2010; Pascut & Ponman 2015). From the correlation between central density and cooling time in the REXCESS sample, Pratt et al. (2009) defined a central density of $n_{e,0} = 0.04 \text{ cm}^{-3} h(z)$ as a threshold which segregates CC and non-CC clusters (their Fig. 2). We computed the central density for the Low- z PXI and High- z objects and used the correlation between this quantity and the C_{SB} to translate this density threshold in terms of the C_{SB} , finding that CC clusters have $C_{\text{SB}} > 0.35$ for this classification scheme.

The C_{SB} value can also be used as an indicator of the relaxation state of the cluster. A high concentration is an indication that the core has not been disturbed by recent merger events. The corresponding threshold defined to distinguish relaxed clusters, for example from the anti-correlation observed between $\langle w \rangle$ and C_{SB} and/or visual inspection (Cassano et al. 2010; Lovisari et al. 2017) differs from that used to define the CC/NCC segregation.

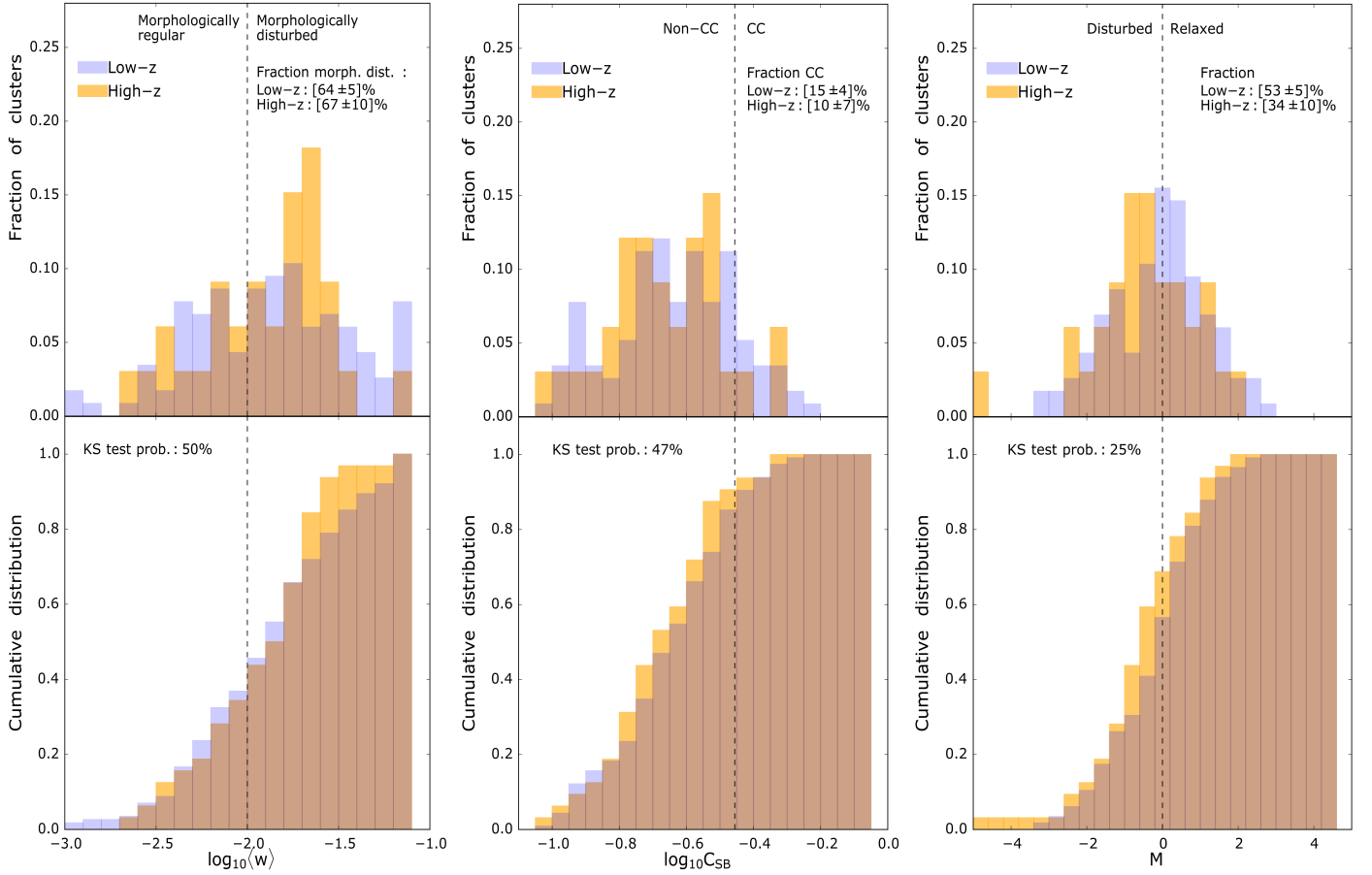


Fig. 3. From left to right, normalised histogram (top panels) and cumulative distributions (bottom panels) of the centroid shift, $\langle w \rangle$, the concentration C_{SB} , and the M parameter (Eq. (7)). The C_{SB} of the High- z sample has been k -corrected as described in Sect. 4.2. The Low- z and High- z distributions are shown in blue and orange, respectively. The vertical dotted line represents the threshold value for each parameter. *Left panel:* $\langle w \rangle = 0.01$ threshold, separating morphologically regular and disturbed clusters. *Central panels:* $C_{SB} = 0.35$ threshold, between cool core (CC) and non-CC objects. *Right panels:* $M = 0$ threshold, between disturbed and relaxed objects. The corresponding fraction above each threshold is given in the top right of each figure. In the *bottom panels*, we report the p -value of the null hypothesis (i.e. that the two distributions are drawn from the same distribution) from the Kolmogorov–Smirnov (KS) test. The distribution of morphologically disturbed and CC objects, based on the parameter $\langle w \rangle$ and C_{SB} , are not statistically different in the High- z and Low- z samples. However, the combination of the two parameters, M , indicates that the fraction of disturbed objects is significantly higher in the High- z sample.

4.3. Combined dynamical indicator, M

The combination of certain morphological parameters has been shown to identify the most disturbed and relaxed clusters (Cassano et al. 2010; Rasia et al. 2013; Lovisari et al. 2017; Cialone et al. 2018). We take advantage of the observed anti-correlation between the centroid shift and the C_{SB} to compute the M parameter introduced by Rasia et al. (2013) and use it as an additional dynamical indicator. M is defined as follows.

$$M \equiv \frac{1}{2} \left(\frac{C_{SB} - C_{SB,med}}{|C_{SB,quar} - C_{SB,med}|} - \frac{\langle w \rangle - \langle w \rangle_{med}}{|\langle w \rangle_{quar} - \langle w \rangle_{med}|} \right), \quad (7)$$

where $C_{SB,med}$ and $\langle w \rangle_{med}$ are the median values of the C_{SB} and centroid shift, respectively, and $C_{SB,quar}$ and $\langle w \rangle_{quar}$ are the first or third quartile depending on whether the parameter value is larger or smaller than the median, respectively. The M parameter is therefore an indicator that combines the large-scale (i.e. centroid shift) and the core (i.e. concentration) properties. It is interesting to note that the two morphological parameters appear in Eq. (7) with the same weight to distinguish relaxed and disturbed objects. This is consistent with what has been derived by Cialone et al. 2018. According to this definition, clusters which are characterised by the presence of a cool core and

are morphologically regular will have $M > 0$ and very disturbed objects with a very diffuse core will have $M < 0$.

Henceforth, we refer to the former and latter objects as ‘relaxed’ and ‘disturbed’, respectively. The choice of this dual classification is arbitrary and does not correspond to a strict segregation between two types of objects. The distribution of M both for Low- z and High- z samples is continuous, as shown in the top right panel of Fig. 3. The results of the M characterisation are reported in the top- and bottom-right panels of Fig. 3 and in Table B.1. We note that the numerical value of $C_{SB,med}$ that we use is smaller than the threshold used in Sect. 4.2 to define CC clusters, and is closer to the value chosen by Lovisari et al. (2017) after visual classification of clusters when defining a similar M parameter⁵.

4.4. Consistency of the morphological characterisation

We used the results of the morphological analysis published in Lovisari et al. (2017) to characterise the Low- z cluster

⁵ $C_{SB,med} = 0.23$ which is close to $C_{SB} = 0.15$, used by Lovisari et al. (2017), after correction for the different aperture definition (their Fig. C1).

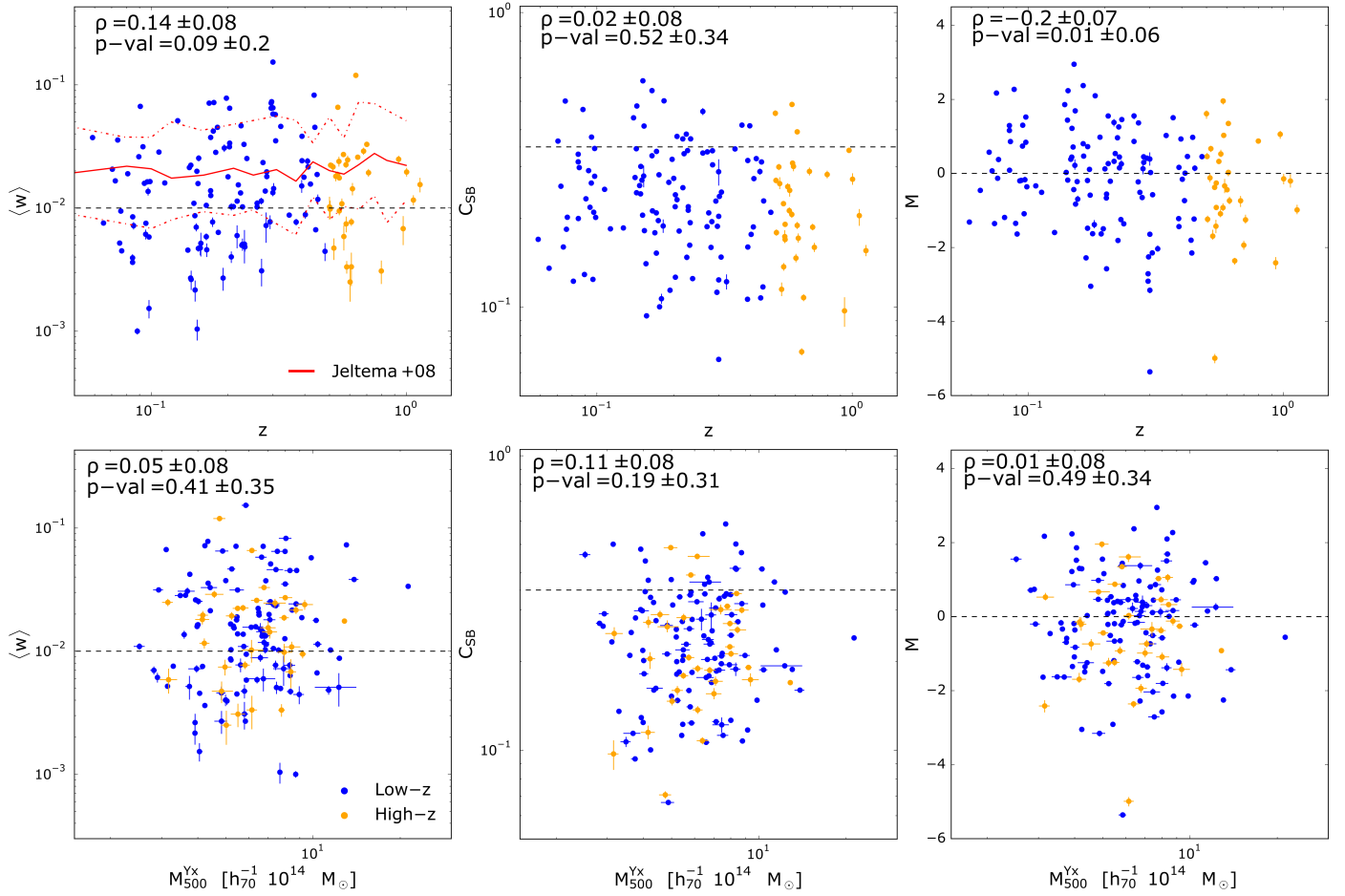


Fig. 4. Morphological parameters vs. redshift (*top panel*) and mass (*bottom panel*) of all the clusters used in this work. High- z and Low- z sample clusters are colour-coded in blue and orange, respectively, following the sample colour code of Fig. 3. *Left panels*: centroid shift, $\langle w \rangle$. The solid red line is the mean $\langle w \rangle$ derived by Jeltema et al. (2008) from numerical simulations. The dotted lines correspond to the $\pm 68\%$ dispersion, *Middle panel*: concentration, C_{SB} . *Right panel*: M parameter. In each panel, the horizontal dotted line identifies the corresponding threshold, as defined in Fig. 3. The symbols ρ and p -value correspond to the Spearman's rank correlation factor and the corresponding null hypothesis p -value, respectively. Errors on these quantities are computed through 1000 bootstrap resampling. The figure corroborates the lack of evolution with redshift of $\langle w \rangle$ and C_{SB} shown in Fig. 3. However, there is a mild but significant evolution of the combined parameter M . There is no dependence on mass of the morphological parameters.

morphological properties, using their $\langle w \rangle$ and C_{SB} values to derive the M parameter. We obtained the morphological parameters of the High- z sample using a different pipeline and different analysis settings. To avoid potentially biased conclusions on morphological evolution or the dependence of the mass profiles on morphology for the full sample, it is necessary to check the consistency between our morphological analysis and that of Lovisari et al. (2017). We thus compared the morphological parameters derived from each pipeline independently for the common clusters of the Low- z PXI sample. The agreement is excellent, as shown in the left and central panel of Fig. A.1. Full details of the comparison are discussed in Appendix A.

5. Morphology and dynamical state

5.1. Sample characterisation and comparison

The results of the Low- z and High- z morphological characterisation are shown in Fig. 3, where the top and bottom of each panel show the normalised and cumulative distribution for each parameter, respectively. We also show the fraction of objects above the fiducial thresholds discussed in Sect. 4.1 in each panel. Errors are computed from 1000 Monte-Carlo bootstrap resamples. They

are dominated by the number of clusters, the individual uncertainties being much smaller than the intrinsic dispersion.

The top left panel shows that both samples contain a majority of disturbed ($\langle w \rangle > 0.01$) objects. The shape of the distributions differs, the High- z sample having a prominent peak of objects around $\langle w \rangle \sim 0.02$. This is reflected in the cumulative distribution, where there is an excess of High- z objects at $\log_{10} \langle w \rangle \sim -1.4$. However, the fraction of disturbed objects in the Low- z sample ($64 \pm 5\%$) is nearly identical to that in the High- z sample ($67 \pm 10\%$), and is consistent within the uncertainties. We investigated if the two samples are representative of the same population by performing the Kolmogorov–Smirnov (KS) test. We determine the p -value of the null hypothesis, that is the two samples are drawn from the same distribution. The p -value is 50%, indicating no significant difference between the Low- z and High- z samples.

We obtained similar results studying the distribution of CC clusters using the C_{SB} parameter, as shown in the central panels of Fig. 3. The fraction of CC clusters is low. The High- z sample has a slightly lower fraction of CC objects ($10 \pm 5\%$) as compared to the Low- z sample ($15 \pm 4\%$), but the difference is not significant. Consistently, the KS test yields a high p -value of 47%.

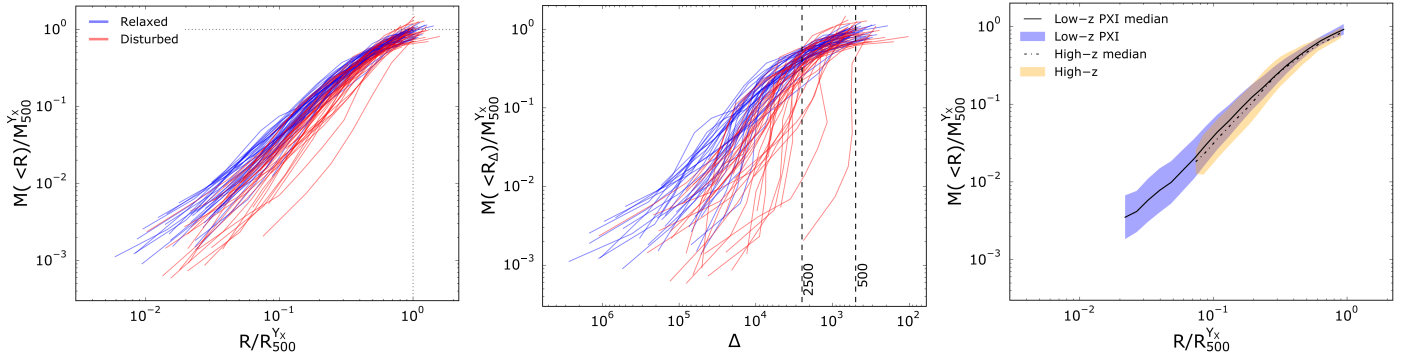


Fig. 5. Scaled radial mass profiles extracted assuming hydrostatic equilibrium for the samples considered in this work. *Left panel:* radius and mass are scaled by R_{500}^{Yx} and M_{500}^{Yx} , respectively. Blue and red profiles represent relaxed and disturbed objects, respectively, according to the M parameter. *Central panel:* same as the left panel, except for the fact that we show the mass profile as a function of the density contrast, Δ . The vertical line $\Delta = 2500$ is the overdensity used in this work to compute the sparsity described in Sect. 6. *Right panel:* comparison of the scaled mass profiles for the Low- z PXI and High- z samples. The solid line and dotted lines represent the median computed for the Low- z PXI and High- z samples, respectively. The gold and green shaded regions represent the 1σ dispersion of the Low- z PXI and High- z samples, respectively. Disturbed clusters have a shallower mass distribution, and present a larger dispersion than that of the most relaxed objects. On the other hand, the median HE mass profile depends mildly on redshift.

Interestingly, the M distributions shown in the right panels of Fig. 3 suggest some evolution. While the distributions have qualitatively similar shapes, the High- z sample has a peak which is clearly shifted towards disturbed objects. Furthermore, the fraction of relaxed clusters in the Low- z sample, $53 \pm 5\%$, is 50% higher than that of the High- z sample, $34 \pm 8\%$, a 2σ effect. This evolution can be seen in the cumulative distribution as a systematic over-abundance of disturbed objects in the High- z sample as compared to the Low- z PXI sample. The KS test yields a smaller p -value of 25%, but this is not small enough to reject the null-hypothesis.

5.2. Mass dependence and redshift evolution

For each morphological parameter, we further quantified the relation with mass and redshift by computing the Spearman's rank (SR) coefficient ρ and the null hypothesis p -value, the probability that the observed coefficient is obtained by chance if the two parameters are completely independent. We also considered the sum square difference of ranks, D , and the number of standard deviations by which D deviates from its null-hypothesis expected value, σ_D . As in the previous section, we performed 1000 bootstrap resamples to estimate these values and their 68% errors. The top and bottom panels of Fig. 4 show each parameter as a function of redshift and mass, respectively. The centroid shift, C_{SB} , and M parameters are shown in the left, middle, and right panels, respectively, with the SR coefficient and corresponding p -value indicated in the top left of each plot.

The only parameter for which there is a correlation with z is the combined M parameter, for which $\rho = -0.2 \pm 0.1$. The correlation is not very significant, with a null-hypothesis p -value of 0.01 ± 0.07 , and a standard deviation on the null hypothesis of $\sigma_D = 2.5 \pm 1$. The Kendall test gives consistent results. This weak correlation of M with z comes from the amplification of the positive (but not significant) trend in $\langle w \rangle$ versus z , while there is no correlation between C_{SB} and redshift.

In summary, consistent with the trend observed in Sect. 5.1 above, there is weak evidence that clusters at higher redshift are slightly more disturbed. On the other hand, there is no evidence for any trend with mass, the SR coefficient for all parameters being consistent with zero and the corresponding p -value in the range 20%–50%. This is in agreement with the mass indepen-

dence of the dynamical state found by Böhringer et al. (2010) and Lovisari et al. (2017) for local clusters. We must note however that the mass range is limited in the present sample.

6. Total mass profile shape

6.1. Radial mass profiles

The individual scaled HE integrated total mass profiles are shown in Fig. 5. In the left and central panels, the profiles are colour coded according to their morphological state according to the M parameter: in blue for the relaxed ($M > 0$) and in red for the disturbed ($M < 0$) clusters.

There is a clear difference between the two populations in the $[0.01-0.5]R_{500}^{Yx}$ range. As compared to the relaxed clusters, the disturbed objects have a shallower mass distribution on average (i.e. there is less mass within the central region), and these profiles show a larger dispersion. This effect is even more evident in the central panel of Fig. 5, where the mass profiles are plotted as a function of the total density contrast, Δ . As Δ is proportional to the mean total density within a sphere of radius R_Δ (Eq. (3)), it decreases with radius, more or less rapidly depending on the steepness of the density profile. For density profiles that are very peaked towards the centre, Δ decreases rapidly, or equivalently $M(<R_\Delta)$ slowly increases with decreasing Δ . On the other hand, a flat profile within the core would correspond to a constant mean density, thus quasi-constant Δ , and a very steep variation of $M(<R_\Delta)$ with Δ in the central region, with a maximum value of Δ corresponding to that of the core. This is what is observed for the relaxed and disturbed clusters, respectively. In summary, the dynamical state of the cluster clearly has a very strong impact on the shape of the total mass profile.

The right panel of Fig. 5 shows the 68% dispersion envelopes of the Low- z PXI and High- z samples with light and dark green, respectively. The black solid and dotted lines represent the median profiles for the Low- z PXI and High- z samples, respectively. The envelopes of the two samples are consistent. However, the median HE mass profile of the High- z sample is slightly lower and shallower than that of the Low- z PXI sample, being lower by 6.5% and 16% at $0.9R_{500}^{Yx}$ and $0.3R_{500}^{Yx}$, respectively. The HE radial mass profiles thus show a hint of evolution, but this is not statistically significant considering the dispersion of

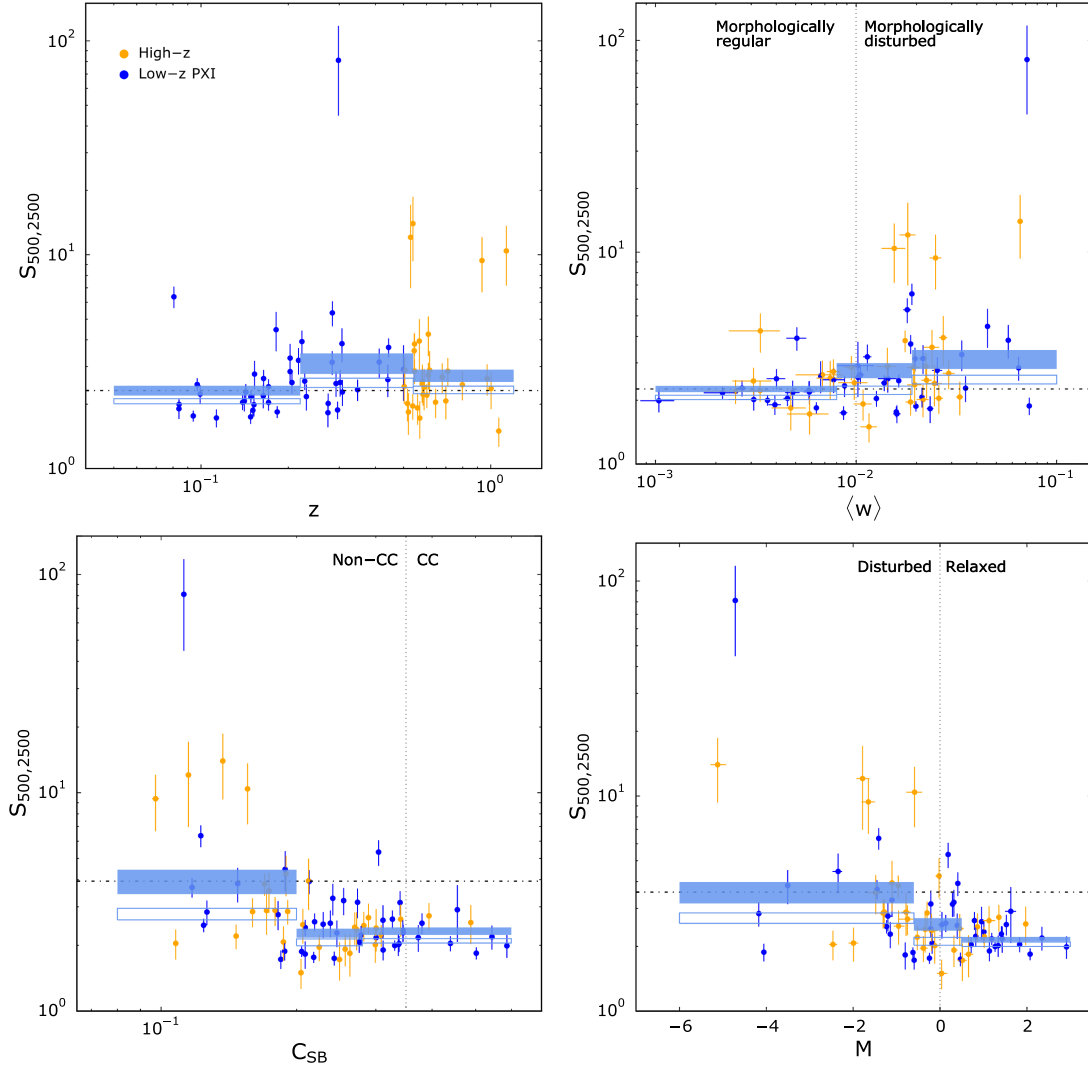


Fig. 6. *Top left panel:* sparsity of the Low- z PXI and High- z samples as a function of redshift. The filled blue rectangles represent the mean sparsity weighted by the statistical errors and the intrinsic dispersion, estimated iteratively (see text), and its 1σ uncertainty. This is computed in three bins, defined to have approximately the same number of objects (see Table 1). The open blue rectangles represent the same quantity computed removing the outliers. We show as reference the median value of the first bin with the black dotted line. *Top right panel:* same but with the sparsity as a function of $\langle w \rangle$. *Bottom left panel:* same but with sparsity as a function of the C_{SB} . *Bottom right panel:* same but with the sparsity as a function of M . The sparsity, i.e. the shape of the total mass profile, varies significantly with the dynamical state indicators. More disturbed clusters have higher sparsity i.e. they are less concentrated. The dependence on redshift is smaller than the dependence on the dynamical state.

the profiles. Interestingly, the dispersions of the two samples are similar. We recall that the samples contain a similar number of objects, the Low- z PXI and High- z having 42 and 33 objects, respectively, and the data quality ensures that HE mass profiles are computed for each object. In view of our finding above that the morphology evolution is negligible, the absence of evolution of the HE mass profile dispersion is a natural consequence of their shape being driven by the dynamical state.

6.2. The shape of the mass profiles

We further quantified the evolution and the impact of dynamical state on the shape of the profile using the sparsity, $S_{500/2500}$. We chose $\Delta_2 = 2500$, which is large enough to encompass the dependence of the mass profile shape as a function of the dynamical state and is reached by all haloes, as shown in the middle panel of Fig. 5.

Figure 6 shows the sparsity as a function of the redshift in the top-left panel, and the three morphological parameters, $\langle w \rangle$,

C_{SB} , and M in the top-right, bottom-left, and bottom-right panels, respectively. The statistical errors on the sparsity are not negligible as compared to the intrinsic scatter, and the simple correlation tests used in the previous section cannot be applied. We therefore computed the mean sparsity in bins. We defined three bins in redshift and for each parameter, the width of each bin being defined so as to have roughly the same number of objects in each bin. The logarithmic mean of the sparsity was computed in each bin, weighting each value by the quadratic sum of the statistical error and the intrinsic scatter. This scatter and the weighted mean were estimated simultaneously by iteration. The mean and intrinsic scatter, together with their 68% errors, were computed using 1000 bootstrap resamples. The results are reported in Table 1 and as filled blue rectangles in each panel of Fig. 6. As there is a large scatter with the presence of strong outliers, we also computed the mean within the same bins excluding the $>3\sigma$ outliers. The results are shown with the open blue rectangles.

There seems to be a slight although not very significant evolution of the sparsity with z . Higher-redshift ($z > 0.54$) clusters

Table 1. Mean values of the sparsity and their uncertainties computed in the bins shown in Fig. 6 as a function of z , $\langle w \rangle$, C_{SB} , and M .

z	Mean	$\sigma_{\text{log,int}}$	$\langle w \rangle [10^{-2}]$	Mean	$\sigma_{\text{log,int}}$	C_{SB}	Mean	$\sigma_{\text{log,int}}$	M	Mean	$\sigma_{\text{log,int}}$
[0.05, 0.22]	2.32 ± 0.12 (2.06 ± 0.06)	0.11 ± 0.04	[0.10, 0.80]	2.26 ± 0.08 (2.06 ± 0.05)	0.05 ± 0.03	[0.08, 0.20]	3.94 ± 0.51 (2.78 ± 0.17)	0.28 ± 0.09	[−6.00, −0.60]	3.53 ± 0.39 (2.71 ± 0.15)	0.25 ± 0.10
[0.22, 0.54]	3.11 ± 0.33 (2.52 ± 0.13)	$[0.23 \pm 0.12]$	[0.80, 1.90]	2.77 ± 0.23 (2.24 ± 0.10)	0.17 ± 0.04	[0.20, 0.28]	2.28 ± 0.11 (2.06 ± 0.07)	0.08 ± 0.02	[−0.60, 0.50]	2.55 ± 0.17 (2.11 ± 0.09)	0.13 ± 0.03
[0.54, 1.20]	2.73 ± 0.19 (2.34 ± 0.09)	0.13 ± 0.05	[1.90, 10.00]	3.12 ± 0.32 (2.51 ± 0.12)	0.23 ± 0.11	[0.28, 0.60]	2.33 ± 0.09 (2.10 ± 0.05)	0.07 ± 0.03	[0.50, 3.00]	2.16 ± 0.05 (2.04 ± 0.05)	0.02 ± 0.02

Notes. The means are computed in logarithmic space and take into account both statistical errors and the intrinsic dispersion, estimated iteratively. The intrinsic dispersion $\sigma_{\text{log,int}}$ in dex are given in the table. The values between parentheses are the sparsities computed within the same bins excluding the outliers.

have a slightly larger value of $S_{500/2500}$ by 18%, that is these profiles are less concentrated, which is not consistent with the low-redshift ($z < 0.2$) clusters at $\sim 1.8\sigma$. We found the same behaviour when excluding the outliers.

There is a much stronger and significant variation of the sparsity with dynamical state: morphologically disturbed (high $\langle w \rangle$), non-CC (low C_{SB}), and disturbed (low M) clusters have larger sparsity. For all parameters considered, the sparsity of the first and third bins is not consistent at more than 3σ , and the difference between the sparsities of these bins is of the order of $\sim 50\%$. At the same time, the intrinsic scatter increases significantly, reaching ~ 0.2 dex for the most disturbed or least concentrated objects. For example, the sparsity increases by 64% at a significance level of 3.4σ between $M < -0.6$ and $M > 0.5$, that is, between the one third most relaxed and the one third most disturbed as defined by this parameter. Only an upper limit on the intrinsic scatter can be estimated from the former (< 0.04 dex), while the intrinsic scatter for the latter reaches 0.25 ± 0.09 dex. There are strong outliers at high sparsity. Excluding the outliers yields the same qualitative results, although the variation between bins is weaker. This suggests a sparsity distribution skewed towards high values, with a skewness increasing with departure from dynamical relaxation.

7. Discussion

7.1. Dynamical state

The first result of our study is that SZ-selected samples are dominated at all redshifts by disturbed and non-CC objects. Recent observational work on local clusters has converged to similar results (e.g. Lovisari et al. 2017; Rossetti et al. 2017; Andrade-Santos et al. 2017; Lopes et al. 2018). In particular, these works highlight the higher fraction of disturbed objects or the lower fraction of CC objects in SZ-selected samples as compared to X-ray-selected samples, a fact interpreted to be due to preferential detection of relaxed or more concentrated clusters in X-ray surveys.

In contrast, there is little consensus in the literature concerning the evolution of the dynamical state, as determined from various morphological parameters. Studies of the evolution up to $z \sim 1$ of the centroid shift and/or power ratios of X-ray-selected clusters indicate a larger fraction of disturbed clusters at high z (Maughan et al. 2008; Jeltema et al. 2005; Weißmann et al. 2013). However, the latter study is also consistent with no evolution, and Nurgaliev et al. (2017) did not find any significant evolution in the redshift range $0.3 < z < 1$ for the 400 d X-ray-selected clusters.

These somewhat contradictory results may simply be due to selection effects: X-ray detectability is clearly not independent of

cluster morphology. More peaked clusters (usually relaxed) are not only more luminous at a given mass, but are also easier to detect at a given X-ray luminosity. Such effects are particularly important in flux-limited surveys, as shown by Chon & Böhringer (2017) in the context of a volume-limited X-ray survey. It is therefore difficult to disentangle selection effects and/or z and mass dependence (see also the discussion in Mantz et al. 2015).

Sunyaev–Zeldovich detection does not suffer from these limitations, and an SZ-selected sample is expected to be close to mass-limited. The morphological evolution of SPT clusters was studied recently by Nurgaliev et al. (2017) using their newly introduced a_{phot} parameter. They did not find a significant difference between the redshift ranges [0.3–0.6] and [0.6–1.2]. This absence of significant evolution was also observed by McDonald et al. (2017) up to $z = 1.6$, also for an SPT sample. The comprehensive study of classification criteria for the most relaxed clusters by Mantz et al. (2015) indicates that the fraction of relaxed clusters in the SPT and *Planck* samples is consistent with being constant with redshift.

In the present study, we extend the morphological analysis of the full SZ-selected population of high-mass clusters, from very local systems $z = 0.05$ up to $z = 1$, applying a consistent sample construction and analysis strategy over the full z range. The high quality of our data allows us to investigate core (i.e. C_{SB}) and bulk (i.e. $\langle w \rangle$) properties at high precision. There are no significant trends either with z or mass in these parameters individually. However, we find a significant evolution with z of the M dynamical indicator, which combines these large-scale and core parameters, with a null-hypothesis p -value of 1%.

It has been suggested that the fraction of disturbed clusters should increase with z and mass in a hierarchical formation model (e.g. Böhringer et al. 2010; McDonald et al. 2017). To explain the observed absence of evolution in the SPT sample, McDonald et al. (2017) proposed a simple model, combining the merger rate from the simulations of Fakhouri et al. (2010) and a fiducial relaxation time of the hot gas equal to the crossing time. In fact, the link between cluster formation history and morphological state as observed in X-rays, as a function of z and mass, is very complex. This first depends on relating the individual mass assembly history to dynamical state (e.g. Power et al. 2012), and then the dynamical state to morphological indicators (e.g. Cui et al. 2017). Individual cluster history is never observed directly and has to be translated into the ensemble properties of cluster samples at different z (e.g. see Mostoghiu et al. 2019). A further complication is the relation between the gas dynamical history and that of the underlying dark matter, and how X-ray morphological observables relate to the gas dynamical state. To our knowledge the only theoretical prediction of the evolution of observed ICM morphological parameters is that of Jeltema et al. (2008). These latter authors claim a significant evolution of the mean $\langle w \rangle$ with

z , although a comparison of their results with our data in Fig. 4 shows that any such evolution is mild, is much smaller than the dispersion, and is fully in agreement with our results.

7.2. Total mass profile

Extending the pilot study of Bartalucci et al. (2018) with a fully SZ-selected sample, the overall picture emerging from the present work is that the shape of the dark matter profiles is affected both by evolution and by dynamical state (Table 1). The evolution effect is mild, increasing the sparsity of objects by $\sim 15\%$ from $z \sim 0.1$ to $z \sim 0.8$. In contrast, the M dependence is much stronger, with a sparsity increasing by $\sim 60\%$ with decreasing M , that is from the most relaxed to the most disturbed objects. This variation of profile shape with dynamical state is likely a fundamental property, rather than a secondary consequence of the mutual variation of the $S_{500,2500}$ and M parameters with z , which are both less significant. A multi-component analysis requiring a larger sample is needed to firmly assess this point.

An obvious question is whether the observed dependence of the sparsity on dynamical state is an artefact of systematic error in the X-ray mass estimate. As discussed in detail by Corasaniti et al. (2018) the sparsity derived from HE mass estimates is essentially bias-free (less than 5%). As it is a mass ratio, the sparsity is only sensitive to the radial dependence of any bias, which is usually small between the density contrasts under consideration. The mean $S_{500,2500}$ bias from the HE mass estimate for example is of $\sim 3.2\%$ from the simulations of Biffi et al. (2016). Generally, although the exact radial dependence of the HE bias will differ from object to object, we expect it to increase towards larger radii (smaller Δ) meaning that sparsity $S_{500/2500}$ measured from HE profiles would be biased low as compared to the true value, especially for the most disturbed objects. This effect is the opposite to the observed increase of sparsity for increasingly perturbed systems. As detailed in Sect. 3.5 we are not directly using the HE mass at $\Delta = 500$ but its proxy, M_{500}^{YX} . For the clusters for which no extrapolation is needed, the differences between these measures are the order of 10%, with no systematic trend with dynamical state. Thus, the measured sparsity may be slightly lower than if we had used the HE mass, but the clear trend of the sparsity with dynamical state would not be changed.

The trend of sparsity with dynamical state indicates that morphologically disturbed objects are less concentrated than relaxed objects. There is also evidence for increased scatter. This is qualitatively in agreement with the difference in the c - M relations for relaxed versus disturbed objects seen in numerical simulations (e.g. Bhattacharya et al. 2013, Figs. 1 and 4). In Le Brun et al. (2018) we performed a preliminary investigation using numerical simulations tailored to cover the high-mass high- z range considered here. In these simulations, the sparsity of the 25 most massive clusters at all redshifts shows a correlation with the DM dynamical indicator Δr^6 with a p -value of $[0.5-2]\%$ (their Fig. 3), indicating that sparser clusters are less regular. We will revisit the link between dynamical state and DM sparsity for the full simulated sample in a forthcoming paper.

8. Conclusions

We present new *XMM-Newton* observations of a *Planck* SZ-selected sample of 28 massive clusters in the redshift range $z = [0.5-0.9]$. These were combined with the sample of

Bartalucci et al. (2018) at $0.9 < z < 1.1$ for a total High- z sample of 33 objects at masses $M_{500}^{YX} > 5 \times 10^{14} M_{\odot}$. We characterised the dynamical state with the centroid shift $\langle w \rangle$, the concentration C_{SB} , and the combination of the two parameters, M , which simultaneously probes the large-scale and core morphology. The shape of the total mass profile, derived from the hydrostatic equilibrium equation, was quantified using the sparsity, the ratio of M_{500} to M_{2500} , that is the masses at density contrast 500 and 2500, respectively. This parameter, introduced by Balmès et al. (2014) offers a non-parametric measurement of the shape which is thought to be relatively insensitive to HE bias (Corasaniti et al. 2018).

We first combined the High- z morphology measurements with those of the ESZ clusters at $z < 0.5$ in Lovisari et al. (2017), for a total sample of 151 objects. In this study:

- We confirmed that SZ-selected samples, thought to best reflect the underlying cluster population, are dominated by disturbed ($\sim 65\%$) and non-CC ($\sim 80\%$) objects, at all redshifts.
- There is no significant evolution or mass dependence of the fraction of cool core or of the centroid shift parameter. The only parameter for which there is a significant correlation with z is the combined M parameter, for which $\rho = -0.2 \pm 0.1$ and a null-hypothesis p -value of 0.01.

We then combined the mass measurements obtained for our new data with those from a subsample of 42 ESZ objects with spatially resolved ICM profiles presented in Planck Collaboration Int. V (2013), and which we confirmed is representative of the full Low- z sample. The total sample of 75 objects covers the redshift range $0.08 < z < 1.1$ and mass range $[5-20] \times 10^{14} M_{\odot}$. We made the following findings.

- The median scaled mass profile differs by less than 6.5% and 16% at $0.9R_{500}^{YX}$ and $0.3R_{500}^{YX}$, respectively, between the Low- z PXI and High- z samples, with no difference in the dispersion. The evolution of the sparsity with z is mild: it increases by only 18% between $z < 0.2$ and $z > 0.55$, an effect significant at less than 2σ .
- When expressed in terms of a scaled mass profile, there is a clear difference between relaxed and disturbed objects. The latter have a less concentrated mass distribution on average, and their scaled profiles show a much larger dispersion.
- Consequently there is a clear dependence of the sparsity on the dynamical state. When expressed as a function of the M parameter, the sparsity increases by $\sim 60\%$ from the one third most relaxed to the one third most disturbed objects, an effect significant at more than 3σ level. We discussed the fact that the HE bias will not significantly change this result.

The main result of this work is that the radial mass distribution is chiefly governed by the dynamical state of the cluster and only mildly dependent on redshift. This has important consequences:

- A coherent sample selection at all z is key. For instance, one cannot compare the c - M relation calibrated at low z on the most relaxed X-ray selected clusters to that of a SZ-selected sample at high z . Ideally, one should consider complete samples, representative of the full true underlying cluster population; for example, a mass-selected sample. This is even more critical when comparing theory and observation, in view of the difficulty of defining coherent dynamical indicators between the two.
- To test theoretical predictions, it is insufficient to simply compare median or stacked properties at each z . The dispersion is a critical quantity, as is the profile distribution, in view of likely departure from log-normality. This requires the measurement of individual profiles. X-ray observations

⁶ Δr is defined as the distance between the centre of mass and the centre of the shrinking sphere (Le Brun et al. 2018).

currently provide the best way to obtain such profiles at high statistical precision.

- In view of this, observational and theoretical efforts to understand the HE bias and its radial dependence are all the more important.

In a forthcoming paper, we will extend our study of the dependence of the sparsity on the dynamical state using an extension of the dark matter simulations presented in [Le Brun et al. \(2018\)](#) to a larger sample of clusters at $M_{500} > 5 \times 10^{14} M_{\odot}$. On a longer timescale, the link between the true dark matter distribution and its dynamical state and the X-ray observables will need to be better understood. This includes the link between ICM morphological proxies and the true underlying dynamical state, and the potential critical issue of the radial variation of the HE bias. We will address these issues with dedicated simulations.

On the observational side, we will investigate the HE bias by comparing our results to weak lensing mass measurements for a subsample of the present data set. The observed lack of significant evolution needs to be tested with a larger sample, particularly at $z > 0.7$ where the present sample is limited, with data of the same or better quality. The fundamental link uncovered by the present paper between the mass profile and dynamical state will also need to be consolidated with better Low- z data. Our current Low- z sample relies on archival data of uneven quality and is not a complete sample. The building of a new local reference SZ-selected sample, with high-quality ICM thermodynamic and HE mass profiles, will be one of the main outcomes of the AO17 *XMM-Newton* heritage program ‘Witnessing the culmination of structure formation in the Universe’, and will provide the necessary inputs.

Acknowledgements. The authors would like to thank Barbara Sartoris for helpful comments and suggestions. The results reported in this article are based on data obtained with *XMM-Newton*, an ESA science mission with instruments and contributions directly funded by ESA Member States and NASA. This work was supported by CNES. The research leading to these results has received funding from the European Research Council under the European Union’s Seventh Framework Programme (FP7/2007-2013) ERC grant agreement no 340519. LL acknowledges support from NASA through contracts 80NSSCK0582 and 80NSSC19K0116.

References

- Amodeo, S., Ettori, S., Capasso, R., & Sereno, M. 2016, [A&A](#), **590**, A126
- Andrade-Santos, F., Jones, C., Forman, W. R., et al. 2017, [ApJ](#), **843**, 76
- Arnaud, M., Neumann, D. M., Aghanim, N., et al. 2001, [A&A](#), **365**, L80
- Arnaud, M., Pointecouteau, E., & Pratt, G. W. 2005, [A&A](#), **441**, 893
- Arnaud, M., Pratt, G. W., Piffaretti, R., et al. 2010, [A&A](#), **517**, A92
- Balmès, I., Rasera, Y., Corasaniti, P.-S., & Alimi, J.-M. 2014, [MNRAS](#), **437**, 2328
- Bartalucci, I., Arnaud, M., Pratt, G. W., et al. 2017, [A&A](#), **598**, A61
- Bartalucci, I., Arnaud, M., Pratt, G. W., & Le Brun, A. M. C. 2018, [A&A](#), **617**, A64
- Bhattacharya, S., Habib, S., Heitmann, K., & Vikhlinin, A. 2013, [ApJ](#), **766**, 32
- Biffi, V., Borgani, S., Murante, G., et al. 2016, [ApJ](#), **827**, 112
- Bleem, L. E., Stalder, B., de Haan, T., et al. 2015, [ApJS](#), **216**, 27
- Böhringer, H., Pratt, G. W., Arnaud, M., et al. 2010, [A&A](#), **514**, A32
- Cassano, R., Ettori, S., Giacintucci, S., et al. 2010, [ApJ](#), **721**, L82
- Chon, G., & Böhringer, H. 2017, [A&A](#), **606**, L4
- Cialone, G., De Petris, M., Sembolini, F., et al. 2018, [MNRAS](#), **477**, 139
- Corasaniti, P. S., Ettori, S., Rasera, Y., et al. 2018, [ApJ](#), **862**, 40
- Correa, C. A., Wyithe, J. S. B., Schaye, J., & Duffy, A. R. 2015, [MNRAS](#), **452**, 1217
- Croston, J. H., Arnaud, M., Pointecouteau, E., & Pratt, G. W. 2006, [A&A](#), **459**, 1007
- Croston, J. H., Pratt, G. W., Böhringer, H., et al. 2008, [A&A](#), **487**, 431
- Cui, W., Power, C., Borgani, S., et al. 2017, [MNRAS](#), **464**, 2502
- da Silva, A. C., Kay, S. T., Liddle, A. R., & Thomas, P. A. 2004, [MNRAS](#), **348**, 1401
- Diemer, B., & Kravtsov, A. V. 2015, [ApJ](#), **799**, 108
- Dolag, K., Bartelmann, M., Perrotta, F., et al. 2004, [A&A](#), **416**, 853
- Dutton, A. A., & Macciò, A. V. 2014, [MNRAS](#), **441**, 3359
- Einasto, J. 1965, *Trudy Astrofizicheskogo Instituta Alma-Ata*, **5**, 87
- Fakhouri, O., Ma, C.-P., & Boylan-Kolchin, M. 2010, [MNRAS](#), **406**, 2267
- Fruscione, A., McDowell, J. C., Allen, G. E., et al. 2006, [Proc. SPIE](#), **6270**, 62701V
- Hasselfield, M., Hilton, M., Marriage, T. A., et al. 2013, [JCAP](#), **7**, 008
- Jeltema, T. E., Canizares, C. R., Bautz, M. W., & Buote, D. A. 2005, [ApJ](#), **624**, 606
- Jeltema, T. E., Hallman, E. J., Burns, J. O., & Motl, P. M. 2008, [ApJ](#), **681**, 167
- Jing, Y. P. 2000, [ApJ](#), **535**, 30
- Kalberla, P. M. W., Burton, W. B., Hartmann, D., et al. 2005, [A&A](#), **440**, 775
- Klypin, A., Yepes, G., Gottlöber, S., Prada, F., & Heß, S. 2016, [MNRAS](#), **457**, 4340
- Kravtsov, A. V., & Borgani, S. 2012, [ARA&A](#), **50**, 353
- Kravtsov, A. V., Vikhlinin, A., & Nagai, D. 2006, [ApJ](#), **650**, 128
- Kuntz, K. D., & Snowden, S. L. 2000, [ApJ](#), **543**, 195
- Le Brun, A. M. C., Arnaud, M., Pratt, G. W., & Teyssier, R. 2018, [MNRAS](#), **473**, L69
- Lopes, P. A. A., Trevisan, M., Laganá, T. F., et al. 2018, [MNRAS](#), **478**, 5473
- Lovisari, L., Forman, W. R., Jones, C., et al. 2017, [ApJ](#), **846**, 51
- Ludlow, A. D., Navarro, J. F., Angulo, R. E., et al. 2014, [MNRAS](#), **441**, 378
- Lumb, D. H., Warwick, R. S., Page, M., & De Luca, A. 2002, [A&A](#), **389**, 93
- Mantz, A. B., Allen, S. W., Morris, R. G., et al. 2015, [MNRAS](#), **449**, 199
- Marriage, T. A., Acquaviva, V., Ade, P. A. R., et al. 2011, [ApJ](#), **737**, 61
- Maughan, B. J., Jones, C., Forman, W., & Van Speybroeck, L. 2008, [ApJS](#), **174**, 117
- McDonald, M., Allen, S. W., Bayliss, M., et al. 2017, [ApJ](#), **843**, 28
- Mohr, J. J., Fabricant, D. G., & Geller, M. J. 1993, [ApJ](#), **413**, 492
- Morrison, R., & McCammon, D. 1983, [ApJ](#), **270**, 119
- Mostoghiu, R., Knebe, A., Cui, W., et al. 2019, [MNRAS](#), **483**, 3390
- Navarro, J. F., Frenk, C. S., & White, S. D. M. 1997, [ApJ](#), **490**, 493
- Navarro, J. F., Hayashi, E., Power, C., et al. 2004, [MNRAS](#), **349**, 1039
- Neto, A. F., Gao, L., Bett, P., et al. 2007, [MNRAS](#), **381**, 1450
- Nurgaliev, D., McDonald, M., Benson, B. A., et al. 2013, [ApJ](#), **779**, 112
- Nurgaliev, D., McDonald, M., Benson, B. A., et al. 2017, [ApJ](#), **841**, 5
- Pascut, A., & Ponman, T. J. 2015, [MNRAS](#), **447**, 3723
- Planck Collaboration VIII. 2011, [A&A](#), **536**, A8
- Planck Collaboration XI. 2011, [A&A](#), **536**, A11
- Planck Collaboration XXIX. 2014, [A&A](#), **571**, A29
- Planck Collaboration XXVII. 2016, [A&A](#), **594**, A27
- Planck Collaboration Int. V. 2013, [A&A](#), **550**, A131
- Power, C., Knebe, A., & Knollmann, S. R. 2012, [MNRAS](#), **419**, 1576
- Pratt, G. W., Croston, J. H., Arnaud, M., & Böhringer, H. 2009, [A&A](#), **498**, 361
- Pratt, G. W., Arnaud, M., Piffaretti, R., et al. 2010, [A&A](#), **511**, A85
- Pratt, G. W., Arnaud, M., Biviano, A., et al. 2019, [Space Sci. Rev.](#), **215**, 25
- Rasia, E., Meneghetti, M., & Ettori, S. 2013, [Astron. Rev.](#), **8**, 40
- Reichardt, C. L., Stalder, B., Bleem, L. E., et al. 2013, [ApJ](#), **763**, 127
- Rossetti, M., Gastaldello, F., Eckert, D., et al. 2017, [MNRAS](#), **468**, 1917
- Santos, J. S., Rosati, P., Tozzi, P., et al. 2008, [A&A](#), **483**, 35
- Santos, J. S., Tozzi, P., Rosati, P., & Böhringer, H. 2010, [A&A](#), **521**, A64
- Starck, J. L., Murtagh, F., & Bijaoui, A. 1998, *Image Processing and Data Analysis: The Multiscale Approach* (New York, USA: Cambridge University Press)
- Strüder, L., Briel, U., Dennerl, K., et al. 2001, [A&A](#), **365**, L18
- Turner, M. J. L., Abbey, A., Arnaud, M., et al. 2001, [A&A](#), **365**, L27
- Velliscig, M., van Daalen, M. P., Schaye, J., et al. 2014, [MNRAS](#), **442**, 2641
- Weißmann, A., Böhringer, H., & Chon, G. 2013, [A&A](#), **555**, A147
- Wu, H.-Y., Hahn, O., Wechsler, R. H., Mao, Y.-Y., & Behroozi, P. S. 2013, [ApJ](#), **763**, 70

Appendix A: Low- z versus Low- z PXI characterisation

In this work we used the results of [Lovisari et al. \(2017\)](#) to characterise the morphological properties of the Low- z sample, using it as anchor for the local universe properties. For this reason, it is mandatory that the morphological parameters we derived for the High- z sample are coherent with the values computed from [Lovisari et al. \(2017\)](#). We derived the centroid shift, C_{SB} , and the M_{500}^{YX} for the Low- z PXI objects we have in common with [Lovisari et al. \(2017\)](#). The comparison between these values are shown in Fig. A.1, denoting our and [Lovisari et al. \(2017\)](#) values with ‘This work’ and ‘L+17’ labels, respectively. The measurements of the centroid shifts shown in the left panel are in good

agreement, with a ratio of 1.07 and a standard deviation of 0.45. The strongest outlier is MACS J2243.3–0935, shown with a red point. The difference is probably caused by the different choice of the X-ray peak (distant by $\sim 7''$) amplified by the high ellipticity of this object. The concentration parameter C_{SB} and the masses, shown in the central and right panels, respectively, are in excellent agreement. The median ratio for both quantities is excellent and the dispersion is remarkably small, the two analysis being performed with different pipelines. This comparison shows that the measurements that require larger samples for statistical reasons, such as the centroid shift, are more sensitive to analysis parameters such as the choice of the centre or the exclusion of point sources; integrated quantities such as the C_{SB} are more robust.

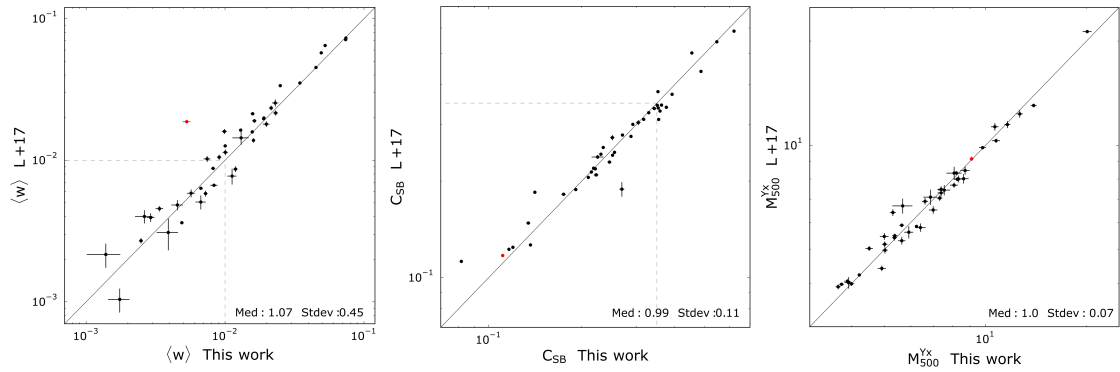


Fig. A.1. *Left panel:* comparison between the centroid shifts computed in this work and published in [Lovisari et al. \(2017\)](#) on the x and y axis, respectively. The red points highlight the outlier. The dotted lines indicate the threshold used to discriminate between disturbed and relaxed clusters. The solid line is the identity relation. The median and standard deviation were computed weighting by the errors and excluding the two outliers. *Central panel:* same as the left panel except that the comparison is done for the C_{SB} . The dotted lines indicate the threshold used to discriminate between CC and non-CC clusters. *Right panel:* same as the left panel except that the comparison is done for the M_{500}^{YX} .

Appendix B: Representativeness of the Low- z PXI

We derived the spatially resolved HE mass profiles for a subsample of 42 clusters of the Low- z sample, namely the Low- z PXI. This subsample covers the same Low- z redshift range and comprises clusters with $M_{500}^{YX} \geq 4 \times 10^{14} M_{\odot}$. We investigated if this subsample is representative of the Low- z population in terms of morphological status.

The Low- z and the Low- z PXI $\langle w \rangle$ distributions are shown in the left panels of Fig. B.1 with green and blue polygons, respectively. In particular, the normalised and cumulative distributions

are shown in the left and bottom panels, respectively. The two samples present qualitatively the same distribution and have a similar fraction of disturbed clusters, the Low- z and the Low- z PXI yielding a fraction of disturbed clusters equal to $[65 \pm 2]\%$ and to $[60 \pm 2]\%$, respectively. This result is confirmed by the high value of the KS test probability: 60%. We found similar results using the C_{SB} . The cluster distributions as a function of this parameter are shown in the right top and bottom panels of Fig. B.1. The two samples have the same fraction of CC objects (15%), and the KS test probability value of 40% confirms that they are representative of the same population.

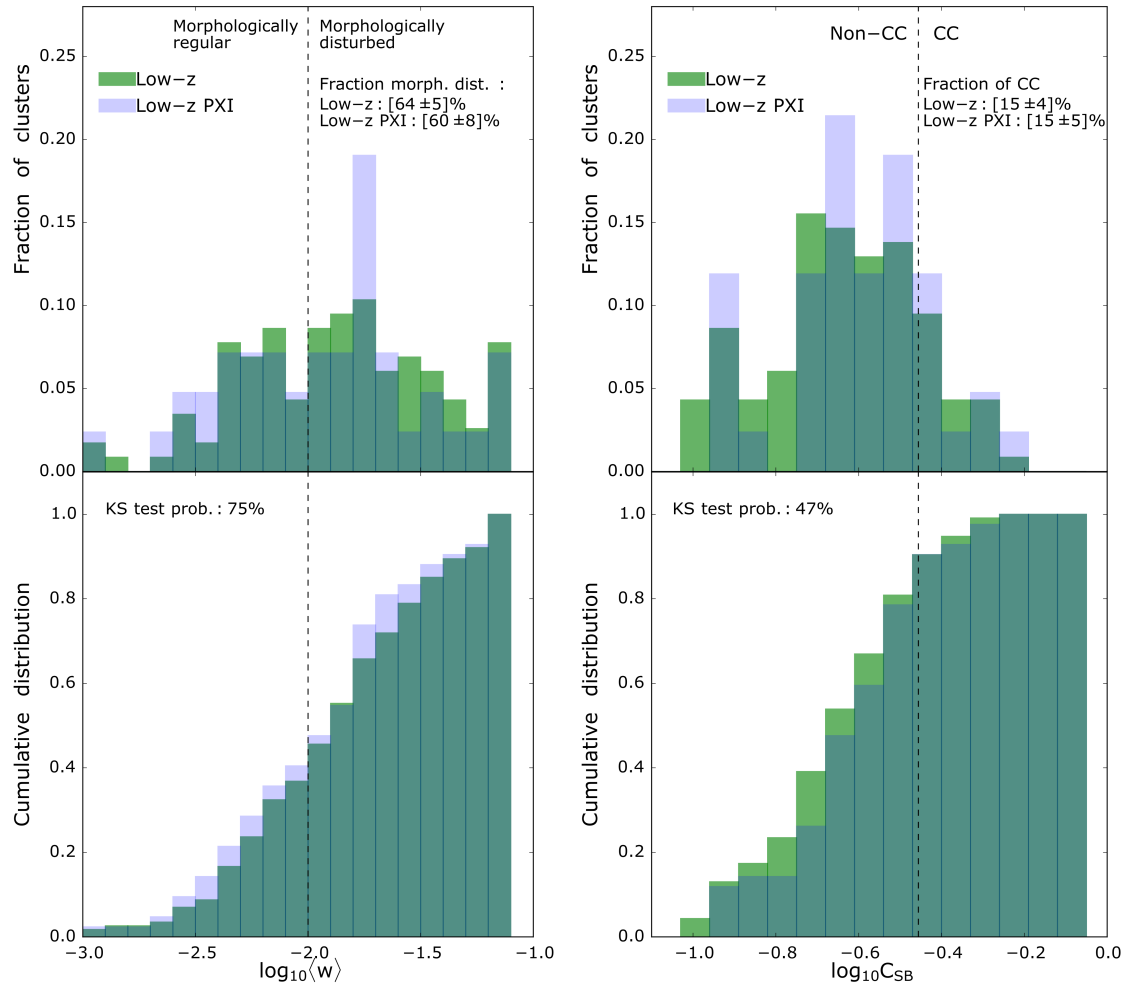


Fig. B.1. *Left top and bottom panels:* normalised histogram and cumulative distribution of the centroid shift distribution of the Low- z and Low- z PXI samples in green and blue, respectively. We report the probability that the two samples are correlated using the KS test. *Left top and bottom panels:* same as the left top and bottom panels, except that we show the normalised and cumulative distribution of the C_{SB} in the top and bottom panels, respectively.

Table B.1. Observational and derived dynamical and morphological properties of the Low- z PXI and High- z samples used in this work.

Planck name	SPT name	ACT name	Alt. name	z	RA [J2000]	Dec [J2000]	N_H [10^{20} cm^{-2}]	Exp. MOS1+2, PN [ks]	Obs. Id	R_{500}^X [kpc]	M_{500}^X [$10^{14} M_\odot$]	M_{2500}^{HE} [$10^{14} M_\odot$]	S	$\langle u \rangle$ [10^{-2}]	C_{58} [10^{-2}]	kC_{58} [10^{-2}]	M
PSZ2 G093.92+34.92			A2255	0.081	258.2029	64.0636	2.50	6.4	0112260801	1176 $^{+16}_{-12}$	5.00 $^{+0.21}_{-0.15}$	0.79 $^{+0.37}_{-0.70}$	6.36 $^{+0.74}_{-0.73}$	1.90 $^{+0.95}_{-0.86}$	12.26 $^{+0.90}_{-0.90}$		-1.41 $^{+0.85}_{-0.85}$
PSZ2 G306.77+58.61			A1651	0.084	194.8438	-4.1983	1.81	7.4	0203020101	1135 $^{+10}_{-12}$	4.51 $^{+0.12}_{-0.15}$	2.37 $^{+2.61}_{-2.61}$	1.90 $^{+0.20}_{-0.20}$	0.39 $^{+0.03}_{-0.03}$	31.14 $^{+0.12}_{-0.12}$		1.14 $^{+0.02}_{-0.02}$
PSZ2 G306.66+61.06			A1650	0.084	194.6734	-1.7622	0.72	34.28	0093200101	1110 $^{+4}_{-4}$	4.21 $^{+0.05}_{-0.05}$	2.11 $^{+2.22}_{-2.22}$	2.00 $^{+0.11}_{-0.11}$	0.36 $^{+0.01}_{-0.01}$	33.05 $^{+0.03}_{-0.03}$		1.27 $^{+0.06}_{-0.06}$
PSZ2 G321.98+47.96	SPT-CL J2249-6426		A3921	0.084	342.4917	-64.4294	1.61	30.23	0112240101	1062 $^{+6}_{-6}$	3.73 $^{+0.06}_{-0.06}$	2.12 $^{+2.00}_{-2.00}$	1.76 $^{+0.10}_{-0.10}$	1.59 $^{+0.04}_{-0.04}$	21.84 $^{+0.01}_{-0.01}$		-0.24 $^{+0.02}_{-0.02}$
PSZ2 G336.60-55.43			A3911	0.097	341.5711	-52.7261	1.50	23.10	0149670301	1085 $^{+7}_{-7}$	3.99 $^{+0.08}_{-0.08}$	1.61 $^{+1.50}_{-1.50}$	2.47 $^{+0.18}_{-0.18}$	1.63 $^{+0.04}_{-0.04}$	12.43 $^{+0.01}_{-0.01}$		-1.23 $^{+0.02}_{-0.02}$
PSZ2 G332.23-46.37			A3827	0.099	330.4720	-59.9454	2.09	21.11	0149670101	1217 $^{+7}_{-7}$	5.64 $^{+0.09}_{-0.09}$	2.53 $^{+2.78}_{-2.78}$	2.23 $^{+0.22}_{-0.22}$	0.58 $^{+0.03}_{-0.03}$	27.83 $^{+0.05}_{-0.05}$		0.82 $^{+0.02}_{-0.02}$
PSZ2 G353.53+59.52	SPT-CL J2201-5956		A2034	0.113	227.5528	33.5104	1.54	10.6	0303930101	1211 $^{+10}_{-10}$	5.64 $^{+0.14}_{-0.14}$	3.27 $^{+3.57}_{-3.57}$	1.73 $^{+0.17}_{-0.17}$	1.60 $^{+0.07}_{-0.07}$	18.45 $^{+0.01}_{-0.01}$		-0.60 $^{+0.03}_{-0.03}$
PSZ2 G241.79-24.01			A3378	0.139	91.4752	-35.3023	4.02	14.11	0201901001	1061 $^{+8}_{-8}$	3.88 $^{+0.09}_{-0.09}$	1.90 $^{+1.76}_{-1.76}$	2.04 $^{+0.16}_{-0.16}$	0.46 $^{+0.02}_{-0.02}$	43.95 $^{+0.10}_{-0.10}$		1.83 $^{+0.01}_{-0.01}$
PSZ2 G002.77-56.16			A3856	0.141	334.6677	-38.9018	1.13	14.4	0201903001	1146 $^{+11}_{-11}$	4.92 $^{+0.14}_{-0.14}$	2.38 $^{+2.13}_{-2.13}$	2.07 $^{+0.22}_{-0.22}$	0.46 $^{+0.02}_{-0.02}$	27.56 $^{+0.09}_{-0.09}$		1.41 $^{+0.01}_{-0.01}$
PSZ2 G226.18+76.79			A1413	0.143	178.8250	23.4049	1.84	62.42	0502690201	1240 $^{+4}_{-4}$	6.24 $^{+0.06}_{-0.06}$	2.73 $^{+2.97}_{-2.97}$	2.28 $^{+0.20}_{-0.20}$	0.27 $^{+0.01}_{-0.01}$	34.50 $^{+0.06}_{-0.06}$		1.41 $^{+0.01}_{-0.01}$
PSZ2 G236.92-26.65			A3364	0.148	86.9081	-31.8727	2.07	22.14	0201900901	1151 $^{+8}_{-8}$	5.01 $^{+0.10}_{-0.10}$	2.88 $^{+2.08}_{-2.08}$	1.74 $^{+0.13}_{-0.13}$	0.87 $^{+0.04}_{-0.04}$	24.25 $^{+0.09}_{-0.09}$		0.47 $^{+0.03}_{-0.03}$
PSZ2 G008.47-56.34			A3854	0.149	334.4400	-37.7260	1.20	22.14	0201902901	1035 $^{+6}_{-6}$	3.64 $^{+0.06}_{-0.06}$	1.68 $^{+1.58}_{-1.58}$	2.17 $^{+0.14}_{-0.14}$	0.22 $^{+0.04}_{-0.04}$	30.04 $^{+0.10}_{-0.10}$		1.19 $^{+0.03}_{-0.03}$
PSZ2 G003.93-59.41			A3888	0.151	338.6125	-37.7360	1.32	20.13	0404910801	1304 $^{+8}_{-8}$	7.31 $^{+0.14}_{-0.14}$	3.90 $^{+2.70}_{-2.70}$	1.87 $^{+0.11}_{-0.11}$	1.99 $^{+0.03}_{-0.03}$	20.53 $^{+0.01}_{-0.01}$		0.92 $^{+0.02}_{-0.02}$
PSZ2 G021.10+33.24			A2204	0.152	248.1959	5.5754	6.97	14.8	0306490201	1347 $^{+13}_{-13}$	8.06 $^{+0.23}_{-0.23}$	4.05 $^{+4.53}_{-4.53}$	1.99 $^{+0.24}_{-0.24}$	0.10 $^{+0.02}_{-0.02}$	58.66 $^{+0.23}_{-0.23}$		2.91 $^{+0.01}_{-0.01}$
PSZ2 G244.71+32.50			A0868	0.153	146.3575	-8.6557	3.59	8.5	0017540101	1058 $^{+13}_{-13}$	3.91 $^{+0.15}_{-0.15}$	1.47 $^{+1.63}_{-1.63}$	2.77 $^{+0.42}_{-0.42}$	0.54 $^{+0.15}_{-0.15}$	18.19 $^{+0.01}_{-0.01}$		-1.20 $^{+0.09}_{-0.09}$
PSZ2 G049.22+30.87			RX J1220.1+2638	0.164	260.0417	26.6250	5.65	14.8	0500670401	1166 $^{+8}_{-8}$	5.30 $^{+0.11}_{-0.11}$	1.42 $^{+1.21}_{-1.21}$	2.19 $^{+0.42}_{-0.42}$	0.58 $^{+0.15}_{-0.15}$	54.33 $^{+0.20}_{-0.20}$		2.35 $^{+0.03}_{-0.03}$
PSZ2 G263.68-22.55			A3404	0.164	101.3712	-54.2273	5.60	11.7	0404910401	1302 $^{+10}_{-10}$	7.38 $^{+0.18}_{-0.18}$	2.80 $^{+2.84}_{-2.84}$	2.64 $^{+0.25}_{-0.25}$	1.05 $^{+0.05}_{-0.05}$	32.66 $^{+0.38}_{-0.38}$		0.79 $^{+0.03}_{-0.03}$
PSZ2 G097.72+38.12			A2218	0.171	248.9597	66.2125	2.60	17.11	0112980101	1169 $^{+7}_{-7}$	5.37 $^{+0.10}_{-0.10}$	2.23 $^{+2.41}_{-2.41}$	2.41 $^{+0.21}_{-0.21}$	1.38 $^{+0.05}_{-0.05}$	20.91 $^{+0.06}_{-0.06}$		-0.21 $^{+0.03}_{-0.03}$
PSZ2 G067.17+67.46			A1914	0.171	216.5105	37.8243	1.06	15.7	0112230201	1348 $^{+10}_{-10}$	8.25 $^{+0.19}_{-0.19}$	4.05 $^{+4.38}_{-4.38}$	2.04 $^{+0.17}_{-0.17}$	1.26 $^{+0.01}_{-0.01}$	33.78 $^{+0.02}_{-0.02}$		0.72 $^{+0.01}_{-0.01}$
PSZ2 G149.75+34.68			A0665	0.182	127.7462	65.8398	4.24	5.2	0109890401	1334 $^{+23}_{-23}$	8.07 $^{+0.41}_{-0.41}$	1.81 $^{+1.45}_{-1.45}$	4.46 $^{+0.93}_{-0.93}$	4.52 $^{+0.09}_{-0.09}$	18.87 $^{+0.99}_{-0.99}$		-2.35 $^{+0.11}_{-0.11}$
PSZ2 G313.33+61.13			A1689	0.183	197.8726	-1.3417	1.52	36.27	0093030101	1339 $^{+6}_{-6}$	8.19 $^{+0.11}_{-0.11}$	4.45 $^{+0.17}_{-0.17}$	1.84 $^{+0.12}_{-0.12}$	0.63 $^{+0.01}_{-0.01}$	50.16 $^{+0.07}_{-0.07}$		2.08 $^{+0.01}_{-0.01}$
PSZ2 G195.75-24.32			A0520	0.203	73.5402	2.9212	3.74	20.9	0201510101	1336 $^{+12}_{-12}$	8.31 $^{+0.22}_{-0.22}$	2.92 $^{+3.30}_{-3.30}$	2.84 $^{+0.37}_{-0.37}$	6.46 $^{+0.05}_{-0.05}$	12.65 $^{+0.01}_{-0.01}$		-4.17 $^{+0.02}_{-0.02}$
PSZ2 G006.76+30.45			A2163	0.203	243.9399	-6.1491	16.50	10.6	0112230601	1794 $^{+18}_{-18}$	20.10 $^{+0.61}_{-0.61}$	6.12 $^{+7.11}_{-7.11}$	3.29 $^{+0.54}_{-0.54}$	3.36 $^{+0.01}_{-0.01}$	24.08 $^{+0.02}_{-0.02}$		-1.11 $^{+0.00}_{-0.00}$
PSZ2 G182.59+53.83			A0963	0.206	154.2653	39.0482	1.25	18.12	0084230701	1129 $^{+7}_{-7}$	5.02 $^{+0.10}_{-0.10}$	1.99 $^{+1.77}_{-1.77}$	2.53 $^{+0.38}_{-0.38}$	0.40 $^{+0.04}_{-0.04}$	37.99 $^{+0.05}_{-0.05}$		1.53 $^{+0.02}_{-0.02}$
PSZ2 G294.68-37.01			A2744	0.274	45.9366	-77.8784	8.73	13.14	0084230601	1276 $^{+11}_{-11}$	7.38 $^{+0.19}_{-0.19}$	1.88 $^{+2.11}_{-2.11}$	3.92 $^{+0.46}_{-0.46}$	1.14 $^{+0.05}_{-0.05}$	25.45 $^{+0.05}_{-0.05}$		0.32 $^{+0.04}_{-0.04}$
PSZ2 G166.09+43.38			A1763	0.223	203.8298	41.0001	0.94	12.9	0042341801	1194 $^{+14}_{-14}$	6.41 $^{+0.23}_{-0.23}$	3.18 $^{+3.52}_{-3.52}$	2.02 $^{+0.23}_{-0.23}$	0.51 $^{+0.06}_{-0.06}$	21.34 $^{+0.02}_{-0.02}$		0.41 $^{+0.03}_{-0.03}$
PSZ2 G092.71+73.46			A2219	0.228	250.0837	46.7107	1.76	12.6	0603000501	1481 $^{+14}_{-14}$	11.62 $^{+0.33}_{-0.33}$	4.54 $^{+4.13}_{-4.13}$	2.56 $^{+0.24}_{-0.24}$	1.02 $^{+0.05}_{-0.05}$	21.93 $^{+0.02}_{-0.02}$		0.13 $^{+0.04}_{-0.04}$
PSZ2 G072.62+41.46			A2390	0.231	328.4031	17.6949	8.66	10.8	0111270101	1438 $^{+13}_{-13}$	10.67 $^{+0.24}_{-0.24}$	4.91 $^{+4.24}_{-4.24}$	2.17 $^{+0.31}_{-0.31}$	0.48 $^{+0.04}_{-0.04}$	37.27 $^{+0.14}_{-0.14}$		1.43 $^{+0.03}_{-0.03}$
PSZ2 G073.97-27.82			RX J0303.8-7752	0.274	45.9366	-77.8784	8.73	11.8	0203330101	1261 $^{+17}_{-17}$	7.54 $^{+0.31}_{-0.31}$	4.13 $^{+3.46}_{-3.46}$	1.82 $^{+0.26}_{-0.26}$	2.34 $^{+0.09}_{-0.09}$	20.90 $^{+0.09}_{-0.09}$		-0.80 $^{+0.06}_{-0.06}$
PSZ2 G294.68-37.01			RXC J0532.9-3701	0.275	83.2323	-37.0270	2.90	11.6	0042341801	1194 $^{+14}_{-14}$	6.41 $^{+0.23}_{-0.23}$	3.18 $^{+3.52}_{-3.52}$	2.02 $^{+0.23}_{-0.23}$	0.51 $^{+0.06}_{-0.06}$	21.34 $^{+0.02}_{-0.02}$		0.32 $^{+0.04}_{-0.04}$
PSZ2 G241.76-30.88			RXC J0232.2-4421	0.284	38.0772	-44.3464	2.49	12.7	0042340301	1225 $^{+12}_{-12}$	6.99 $^{+0.21}_{-0.21}$	2.22 $^{+2.50}_{-2.50}$	3.14 $^{+0.40}_{-0.40}$	1.80 $^{+0.08}_{-0.08}$	33.96 $^{+0.06}_{-0.06}$		0.29 $^{+0.03}_{-0.03}$
PSZ2 G259.98-63.43			RXC J0232.2-4421	0.284	82.2211	-39.4714	2.13	7.4	0042340801	1218 $^{+20}_{-20}$	6.86 $^{+0.33}_{-0.33}$	1.28 $^{+1.98}_{-1.98}$	5.35 $^{+0.73}_{-0.73}$	1.87 $^{+0.04}_{-0.04}$	30.43 $^{+0.55}_{-0.55}$		0.19 $^{+0.05}_{-0.05}$
PSZ2 G106.87-83.23			A2813	0.292	10.8519	-20.6229	1.54	11.5	0042340201	1155 $^{+16}_{-16}$	5.91 $^{+0.25}_{-0.25}$	2.37 $^{+2.68}_{-2.68}$	2.50 $^{+0.34}_{-0.34}$	0.77 $^{+0.10}_{-0.10}$	22.92 $^{+0.05}_{-0.05}$		0.40 $^{+0.09}_{-0.09}$
PSZ2 G266.04-21.25			1ES0657-558	0.296	104.6277	-55.9434	4.17	21.14	0112980201	1534 $^{+11}_{-11}$	13.92 $^{+0.29}_{-0.29}$	7.41 $^{+8.09}_{-8.09}$	1.88 $^{+0.18}_{-0.18}$	7.30 $^{+0.03}_{-0.03}$	18.82 $^{+0.08}_{-0.08}$		-4.06 $^{+0.02}_{-0.02}$
PSZ2 G195.60+44.06			A0781	0.298	140.1018	30.5028	1.94	57.47	040170101	1116 $^{+6}_{-6}$	5.36 $^{+0.09}_{-0.09}$	0.07 $^{+0.10}_{-0.10}$	8.13 $^{+36.44}_{-36.44}$	7.11 $^{+0.05}_{-0.05}$	11.24 $^{+0.01}_{-0.01}$		-4.72 $^{+0.02}_{-0.02}$
PSZ2 G125.71+53.86			A1576	0.302	189.2441	63.1871	1.68	6.1	0402250101	1132 $^{+27}_{-27}$	5.67 $^{+0.39}_{-0.39}$	2.24 $^{+2.97}_{-2.97}$	2.53 $^{+0.35}_{-0.35}$	1.44 $^{+0.15}_{-0.15}$	23.77 $^{+0.33}_{-0.33}$		0.04 $^{+0.01}_{-0.01}$
PSZ2 G008.94-81.22			A2744	0.307	3.5775	-30.3863	1.60	14.10	0042340101	1360 $^{+11}_{-11}$	9.81 $^{+0.24}_{-0.24}$	2.56 $^{+3.01}_{-3.01}$	3.84 $^{+0.69}_{-0.69}$	5.74 $^{+0.03}_{-0.03}$	14.80 $^{+0.05}_{-0.05}$		-3.51 $^{+0.02}_{-0.02}$
PSZ2 G278.58+39.16			A1300	0.308	172.9775	-19.9285	4.50	11.9	0042341001	1302 $^{+16}_{-16}$	8.61 $^{+0.33}_{-0.33}$	3.78 $^{+3.29}_{-3.29}$	2.28 $^{+0.32}_{-0.32}$	3.51 $^{+0.08}_{-0.08}$	24.58 $^{+0.12}_{-0.12}$		-1.15 $^{+0.03}_{-0.03}$
PSZ2 G349.46-59.95			AS1063	0.347	342.1824	-44.5305	1.84	25.15	0504630101	1458 $^{+10}_{-10}$	12.66 $^{+0.27}_{-0.27}$	5.41 $^{+4.80}_{-4.80}$	2.34 $^{+0.27}_{-0.27}$	0.88 $^{+0.02}_{-0.02}$	34.45 $^{+0.17}_{-0.17}$		1.01 $^{+0.02}_{-0.02}$
PSZ2 G083.29-31.03			RXC J2228+2037	0.412	337.1405	20.6204	4.26	24.15	0147890101	1255 $^{+12}_{-12}$	8.70 $^{+0.30}_{-0.30}$	2.76 $^{+2.34}_{-2.34}$	3.15 $^{+0.49}_{-0.49}$	2.16 $^{+0.09}_{-0.09}$	27.32 $^{+0.53}_{-0.53}$		-0.21 $^{+0.05}_{-0.05}$
PSZ2 G284.41+52.45																	

Table B.1. continued.

<i>Planck</i> name	SPT name	ACT name	Alt. name	<i>z</i>	RA	Dec	N_{H} [10^{20} cm^{-2}]	Exp. MOS1-2, PN [ks]	Obs. Id	R_{500}^{X} [kpc]	M_{500}^{X} [$10^{14} M_{\odot}$]	M_{500}^{HE} [$10^{14} M_{\odot}$]	<i>S</i>	$\langle\sigma\rangle$ [10^{-2}]	C_{SB} [10^{-2}]	κC_{SB} [10^{-2}]	<i>M</i>
PSZ2 G265.10–59.50	SPT-CL J0243–4833		RXC J0243.6–4834	0.500	40.9129	–48.5011	2.15	15.6	0672090501, 0723780801	1079 $^{+36}_{-21}$	6.13 $^{+0.63}_{-0.42}$	2.11 $^{+2.69}_{-0.82}$	2.91 $^{+0.87}_{-0.41}$	1.02 $^{+0.20}_{-0.14}$	45.52 $^{+1.04}_{-0.82}$	44.77 $^{+1.03}_{-0.82}$	1.63 $^{+0.13}_{-0.12}$
PSZ2 G044.77–51.30			MACS J0214.9–1359	0.503	33.7383	–48.0045	2.88	17.6	0693661901	1175 $^{+21}_{-22}$	7.95 $^{+0.64}_{-0.44}$	3.29 $^{+2.75}_{-0.81}$	2.42 $^{+0.41}_{-0.27}$	0.98 $^{+0.17}_{-0.10}$	26.93 $^{+0.60}_{-0.46}$	26.55 $^{+0.59}_{-0.46}$	0.45 $^{+0.07}_{-0.07}$
PSZ2 G211.21+38.66			MACS J0911.2+1746	0.505	137.7970	17.7760	3.28	33.25	0693662501	1038 $^{+14}_{-14}$	5.49 $^{+0.44}_{-0.31}$	1.90 $^{+1.62}_{-0.51}$	2.89 $^{+0.44}_{-0.31}$	2.23 $^{+0.10}_{-0.10}$	17.92 $^{+0.46}_{-0.46}$	17.77 $^{+0.46}_{-0.46}$	–1.29 $^{+0.12}_{-0.12}$
PSZ2 G004.45–19.55				0.516	289.2692	–33.5228	5.90	14.5	0656201001	1207 $^{+22}_{-22}$	8.73 $^{+0.48}_{-0.48}$	4.34 $^{+5.06}_{-1.06}$	2.01 $^{+0.35}_{-0.35}$	2.16 $^{+0.14}_{-0.14}$	29.93 $^{+0.75}_{-0.75}$	29.56 $^{+0.74}_{-0.74}$	–0.12 $^{+0.15}_{-0.15}$
PSZ2 G10.28–87.48				0.520	12.2959	–24.6792	1.50	26.4	0693662101, 0723780201	989 $^{+30}_{-30}$	4.83 $^{+0.46}_{-0.46}$	2.63 $^{+3.17}_{-0.82}$	1.84 $^{+0.41}_{-0.41}$	0.47 $^{+0.09}_{-0.09}$	26.27 $^{+1.11}_{-1.11}$	26.14 $^{+1.11}_{-1.11}$	0.66 $^{+0.08}_{-0.08}$
PSZ2 G212.44+63.19				0.529	163.2159	24.2584	1.89	34.20	0693660701, 0723780701	937 $^{+17}_{-17}$	4.15 $^{+0.23}_{-0.23}$	0.34 $^{+0.09}_{-0.09}$	12.06 $^{+5.06}_{-5.06}$	1.81 $^{+0.17}_{-0.17}$	11.51 $^{+0.60}_{-0.60}$	11.23 $^{+0.58}_{-0.58}$	–1.78 $^{+0.13}_{-0.13}$
PSZ2 G201.50–27.31			MACS J0454.1–0300	0.538	73.5471	–3.0162	3.92	22.17	0205670101	1159 $^{+14}_{-14}$	7.94 $^{+0.29}_{-0.29}$	4.05 $^{+3.51}_{-1.11}$	1.96 $^{+0.27}_{-0.27}$	1.87 $^{+0.10}_{-0.10}$	22.29 $^{+0.38}_{-0.38}$	22.29 $^{+0.38}_{-0.38}$	–0.38 $^{+0.09}_{-0.09}$
PSZ2 G094.56+51.03			WHL J227.050+57.90	0.539	227.0821	57.9164	1.50	26.20	0693660101, 0723780501	1064 $^{+14}_{-14}$	6.15 $^{+0.25}_{-0.25}$	0.44 $^{+0.58}_{-0.58}$	13.98 $^{+4.66}_{-4.66}$	6.57 $^{+0.18}_{-0.18}$	13.72 $^{+0.42}_{-0.42}$	13.49 $^{+0.41}_{-0.41}$	–5.12 $^{+0.14}_{-0.14}$
PSZ2 G228.16+75.20			MACS J1149.5+2223	0.544	177.3976	22.4011	1.92	13.2	0693661701	121 $^{+27}_{-27}$	9.36 $^{+0.64}_{-0.64}$	2.63 $^{+3.15}_{-1.11}$	3.56 $^{+0.74}_{-0.74}$	2.39 $^{+0.20}_{-0.20}$	17.37 $^{+0.87}_{-0.87}$	17.28 $^{+0.87}_{-0.87}$	–1.48 $^{+0.18}_{-0.18}$
PSZ2 G111.61–45.71			CL0016+16	0.546	4.6399	16.4362	3.99	33.24	0111000101, 0111000201	1214 $^{+11}_{-11}$	9.21 $^{+0.24}_{-0.24}$	3.22 $^{+3.61}_{-1.11}$	2.86 $^{+0.37}_{-0.37}$	0.95 $^{+0.08}_{-0.08}$	19.11 $^{+0.30}_{-0.30}$	18.98 $^{+0.30}_{-0.30}$	–0.30 $^{+0.08}_{-0.08}$
PSZ2 G180.25+21.03			MACS J0717.5+3745	0.546	109.3800	37.7587	6.63	156.116	0672420101, 0672420201, 0672420301	1356 $^{+6}_{-6}$	12.84 $^{+0.17}_{-0.17}$	3.36 $^{+3.96}_{-1.11}$	3.82 $^{+0.45}_{-0.45}$	1.76 $^{+0.04}_{-0.04}$	16.98 $^{+0.11}_{-0.11}$	16.85 $^{+0.11}_{-0.11}$	–0.96 $^{+0.03}_{-0.03}$
PSZ2 G183.90+42.99			WHL J137.713+38.83	0.559	137.7032	38.8357	1.63	14.8	0723780101	1173 $^{+27}_{-27}$	8.44 $^{+0.60}_{-0.60}$	4.39 $^{+5.07}_{-1.11}$	1.92 $^{+0.33}_{-0.33}$	1.08 $^{+0.11}_{-0.11}$	25.65 $^{+0.73}_{-0.73}$	25.48 $^{+0.72}_{-0.72}$	0.34 $^{+0.06}_{-0.06}$
PSZ2 G155.27–68.42			WHL J24.3324–8.477	0.567	24.3536	–8.4557	3.52	28.18	0693662801, 0700180201	1149 $^{+22}_{-22}$	8.01 $^{+0.46}_{-0.46}$	2.03 $^{+2.56}_{-0.80}$	3.95 $^{+1.05}_{-1.05}$	2.72 $^{+0.13}_{-0.13}$	21.26 $^{+0.52}_{-0.52}$	21.02 $^{+0.51}_{-0.51}$	–1.11 $^{+0.12}_{-0.12}$
PSZ2 G046.13+30.72			WHL J17705.5+24042.4	0.569	259.2742	24.0737	5.18	26.2	0693661401	843 $^{+20}_{-20}$	3.17 $^{+0.22}_{-0.22}$	1.84 $^{+2.11}_{-0.80}$	1.72 $^{+0.34}_{-0.34}$	0.59 $^{+0.14}_{-0.14}$	24.92 $^{+1.25}_{-1.25}$	24.57 $^{+1.24}_{-1.24}$	0.51 $^{+0.09}_{-0.09}$
PSZ2 G239.93–39.97				0.580	71.6966	–37.0625	1.44	34.23	0679181001, 0693661201	1022 $^{+15}_{-15}$	5.73 $^{+0.23}_{-0.23}$	2.31 $^{+1.50}_{-0.80}$	2.49 $^{+0.45}_{-0.45}$	2.25 $^{+0.14}_{-0.14}$	20.65 $^{+0.55}_{-0.55}$	20.47 $^{+0.54}_{-0.54}$	–0.95 $^{+0.10}_{-0.10}$
PSZ2 G254.64–45.20	SPT-CL J0417–4748		MACS J20647.7+7015	0.581	64.3464	–47.8134	1.34	20.9	0700182401	974 $^{+17}_{-17}$	4.97 $^{+0.27}_{-0.27}$	1.95 $^{+2.34}_{-0.80}$	2.54 $^{+0.52}_{-0.52}$	0.74 $^{+0.11}_{-0.11}$	48.76 $^{+0.88}_{-0.88}$	47.76 $^{+0.86}_{-0.86}$	1.97 $^{+0.09}_{-0.09}$
PSZ2 G144.83+25.11			MACS J2129.4–0741	0.584	101.9590	70.2481	5.40	72.46	0551850401, 0551851301	1130 $^{+10}_{-10}$	7.78 $^{+0.20}_{-0.20}$	3.51 $^{+3.89}_{-1.12}$	2.22 $^{+0.25}_{-0.25}$	0.33 $^{+0.04}_{-0.04}$	30.83 $^{+0.32}_{-0.32}$	30.66 $^{+0.32}_{-0.32}$	1.02 $^{+0.03}_{-0.03}$
PSZ2 G045.32–38.46			WHL J155625.2+444042	0.589	322.3591	–7.6913	4.32	9.3	0700182001	1109 $^{+10}_{-10}$	7.39 $^{+0.20}_{-0.20}$	3.08 $^{+2.46}_{-0.80}$	2.40 $^{+0.33}_{-0.33}$	0.25 $^{+0.06}_{-0.06}$	30.05 $^{+0.31}_{-0.31}$	29.85 $^{+0.31}_{-0.31}$	–0.34 $^{+0.17}_{-0.17}$
PSZ2 G070.89+49.26			WHL J151820.6+292740	0.602	239.1098	44.6772	1.23	47.25	0693661301	970 $^{+13}_{-13}$	5.02 $^{+0.21}_{-0.21}$	2.28 $^{+2.00}_{-0.80}$	2.20 $^{+0.29}_{-0.29}$	0.77 $^{+0.06}_{-0.06}$	14.68 $^{+0.40}_{-0.40}$	14.62 $^{+0.40}_{-0.40}$	1.36 $^{+0.05}_{-0.05}$
PSZ2 G045.87+57.70			WHL J215.168+39.91	0.609	215.1709	39.9187	0.82	26.19	0693661001	1016 $^{+13}_{-13}$	5.82 $^{+0.22}_{-0.22}$	2.13 $^{+2.43}_{-0.80}$	2.73 $^{+0.39}_{-0.39}$	0.25 $^{+0.08}_{-0.08}$	39.38 $^{+0.62}_{-0.62}$	37.97 $^{+0.60}_{-0.60}$	–0.52 $^{+0.07}_{-0.07}$
PSZ2 G073.31+67.52			WHL J213.697+54.78	0.609	215.1709	39.9187	0.82	26.19	0693661001	1035 $^{+14}_{-14}$	6.15 $^{+0.26}_{-0.26}$	1.45 $^{+1.75}_{-0.80}$	4.25 $^{+0.89}_{-0.89}$	0.33 $^{+0.08}_{-0.08}$	18.95 $^{+0.52}_{-0.52}$	18.71 $^{+0.51}_{-0.51}$	–0.02 $^{+0.07}_{-0.07}$
PSZ2 G099.86+58.45				0.615	213.6952	54.7840	1.50	21.9	0693660601, 0693662701, 0723780301	1082 $^{+21}_{-21}$	7.09 $^{+0.42}_{-0.42}$	2.45 $^{+1.92}_{-0.80}$	2.89 $^{+0.68}_{-0.68}$	1.43 $^{+0.10}_{-0.10}$	17.15 $^{+0.53}_{-0.53}$	17.02 $^{+0.52}_{-0.52}$	–0.78 $^{+0.12}_{-0.12}$
PLCK G193.31–46.13				0.634	53.9644	–6.9758	4.15	67.48	0658200401, 0693661501	940 $^{+15}_{-15}$	4.76 $^{+0.23}_{-0.23}$	0.00 $^{+0.00}_{-0.00}$	0.00 $^{+0.00}_{-0.00}$	11.95 $^{+0.20}_{-0.20}$	7.06 $^{+0.20}_{-0.20}$	6.63 $^{+0.19}_{-0.19}$	–10.04 $^{+0.15}_{-0.15}$
PLCK G147.3–16.6				0.645	44.1056	40.2885	8.29	59.39	0679181301, 0693661601	1033 $^{+16}_{-16}$	6.39 $^{+0.30}_{-0.30}$	3.13 $^{+3.60}_{-1.12}$	2.04 $^{+0.32}_{-0.32}$	2.59 $^{+0.12}_{-0.12}$	10.78 $^{+0.30}_{-0.30}$	10.64 $^{+0.29}_{-0.29}$	–2.46 $^{+0.07}_{-0.07}$
PSZ2 G219.89–34.39				0.680	94.1429	–52.4518	4.25	35.20	0693662301	910 $^{+22}_{-22}$	4.56 $^{+0.34}_{-0.34}$	1.70 $^{+1.94}_{-0.80}$	2.68 $^{+0.43}_{-0.43}$	2.90 $^{+0.20}_{-0.20}$	28.90 $^{+0.30}_{-0.30}$	28.52 $^{+0.29}_{-0.29}$	–0.75 $^{+0.14}_{-0.14}$
PSZ2 G219.89–34.39				0.700	73.6894	–20.2851	3.25	51.34	0679180501, 0693660301	1030 $^{+15}_{-15}$	6.77 $^{+0.29}_{-0.29}$	3.27 $^{+2.69}_{-1.12}$	2.07 $^{+0.37}_{-0.37}$	3.29 $^{+0.13}_{-0.13}$	18.70 $^{+0.51}_{-0.51}$	18.50 $^{+0.51}_{-0.51}$	–1.99 $^{+0.10}_{-0.10}$
PSZ2 G208.61–74.39				0.711	30.0695	–24.9132	1.37	47.28	0693662901, 0723780601	941 $^{+14}_{-14}$	5.23 $^{+0.23}_{-0.23}$	1.83 $^{+1.57}_{-0.80}$	2.85 $^{+0.43}_{-0.43}$	1.93 $^{+0.14}_{-0.14}$	15.97 $^{+0.58}_{-0.58}$	15.64 $^{+0.56}_{-0.56}$	–1.31 $^{+0.12}_{-0.12}$
PSZ2 G352.05–24.01				0.798	290.2490	–45.8500	6.03	52.28	0679180201, 0693660401	925 $^{+18}_{-18}$	5.50 $^{+0.32}_{-0.32}$	2.23 $^{+2.55}_{-0.80}$	2.46 $^{+0.38}_{-0.38}$	0.31 $^{+0.07}_{-0.07}$	28.21 $^{+0.85}_{-0.85}$	28.00 $^{+0.84}_{-0.84}$	0.86 $^{+0.08}_{-0.08}$
SPT-CL J2146–4633				0.933	326.6447	–46.5475	1.64	153.102	0744400501, 0744401301	728 $^{+11}_{-11}$	3.15 $^{+0.14}_{-0.14}$	0.34 $^{+0.24}_{-0.24}$	9.39 $^{+2.73}_{-2.73}$	2.49 $^{+0.17}_{-0.17}$	9.72 $^{+1.09}_{-1.09}$	9.44 $^{+1.06}_{-1.06}$	–1.65 $^{+0.12}_{-0.12}$
PSZ2 G266.54–27.31				0.972	93.9660	–57.7796	4.32	12.3	0658200101	993 $^{+14}_{-14}$	8.38 $^{+0.35}_{-0.35}$	3.18 $^{+3.69}_{-1.12}$	2.63 $^{+0.44}_{-0.44}$	0.68 $^{+0.18}_{-0.18}$	34.06 $^{+0.37}_{-0.37}$	33.43 $^{+0.37}_{-0.37}$	1.13 $^{+0.15}_{-0.15}$
SPT-CL J2341–5119				1.003	355.3010	–51.3286	1.21	91.45	0744400401, 0763670201	777 $^{+11}_{-11}$	4.16 $^{+0.17}_{-0.17}$	1.76 $^{+2.10}_{-0.80}$	2.10 $^{+0.27}_{-0.27}$	1.96 $^{+0.16}_{-0.16}$	27.19 $^{+1.17}_{-1.17}$	26.82 $^{+1.15}_{-1.15}$	–0.20 $^{+0.19}_{-0.19}$
SPT-CL J0546–5345				1.066	86.6551	–53.7596	6.79	127.113	0744400201, 0744400301	762 $^{+10}_{-10}$	4.21 $^{+0.18}_{-0.18}$	2.81 $^{+3.32}_{-1.12}$	1.50 $^{+0.24}_{-0.24}$	20.46 $^{+1.08}_{-1.08}$	20.46 $^{+1.08}_{-1.08}$	19.90 $^{+1.08}_{-1.08}$	0.04 $^{+0.12}_{-0.12}$
SPT-CL J2106–5844				1.132	316.5221	–58.7421	4.33	26.16	0763670301	880 $^{+19}_{-19}$	7.00 $^{+0.43}_{-0.43}$	0.67 $^{+0.47}_{-0.47}$	10.42 $^{+1.26}_{-1.26}$	1.55 $^{+0.21}_{-0.21}$	15.56 $^{+0.70}_{-0.70}$	15.31 $^{+0.69}_{-0.69}$	–0.58 $^{+0.17}_{-0.17}$