



**HAL**  
open science

## High-resolution, 3D radiative transfer modeling - I. The grand-design spiral galaxy M 51

Ilse de Looze, Jacopo Fritz, Maarten Baes, George J. Bendo, Luca Cortese, Médéric Boquien, Alessandro Boselli, Peter Camps, Asantha Cooray, Diane Cormier, et al.

### ► To cite this version:

Ilse de Looze, Jacopo Fritz, Maarten Baes, George J. Bendo, Luca Cortese, et al.. High-resolution, 3D radiative transfer modeling - I. The grand-design spiral galaxy M 51: I. The grand-design spiral galaxy M 51. *Astronomy and Astrophysics - A&A*, 2014, 571, pp.A69. 10.1051/0004-6361/201424747 . cea-01290011

**HAL Id: cea-01290011**

**<https://cea.hal.science/cea-01290011>**

Submitted on 17 Mar 2016

**HAL** is a multi-disciplinary open access archive for the deposit and dissemination of scientific research documents, whether they are published or not. The documents may come from teaching and research institutions in France or abroad, or from public or private research centers.

L'archive ouverte pluridisciplinaire **HAL**, est destinée au dépôt et à la diffusion de documents scientifiques de niveau recherche, publiés ou non, émanant des établissements d'enseignement et de recherche français ou étrangers, des laboratoires publics ou privés.

# High-resolution, 3D radiative transfer modeling

## I. The grand-design spiral galaxy M 51<sup>★</sup>

Ilse De Looze<sup>1,2</sup>, Jacopo Fritz<sup>1</sup>, Maarten Baes<sup>1</sup>, George J. Bendo<sup>3</sup>, Luca Cortese<sup>4</sup>, Médéric Boquien<sup>2</sup>, Alessandro Boselli<sup>5</sup>, Peter Camps<sup>1</sup>, Asantha Cooray<sup>6,7</sup>, Diane Cormier<sup>8</sup>, Jon I. Davies<sup>9</sup>, Gert De Geyter<sup>1</sup>, Thomas M. Hughes<sup>1</sup>, Anthony P. Jones<sup>10,11</sup>, Oskar Ł. Karczewski<sup>12</sup>, Vianney Lebouteiller<sup>13</sup>, Nanyao Lu<sup>14</sup>, Suzanne C. Madden<sup>13</sup>, Aurélie Rémy-Ruyer<sup>10,11,13</sup>, Luigi Spinoglio<sup>15</sup>, Matthew W. L. Smith<sup>8</sup>, Sebastien Viaene<sup>1</sup>, and Christine D. Wilson<sup>16</sup>

<sup>1</sup> Sterrenkundig Observatorium, Universiteit Gent, Krijgslaan 281 S9, 9000 Gent, Belgium  
e-mail: Ilse.DeLooze@ugent.be

<sup>2</sup> Institute of Astronomy, University of Cambridge, Madingley Road, Cambridge, CB3 0HA, UK

<sup>3</sup> UK ALMA Regional Centre Node, Jodrell Bank Centre for Astrophysics, School of Physics and Astronomy, University of Manchester, Oxford Road, Manchester M13 9PL, UK

<sup>4</sup> Centre for Astrophysics & Supercomputing, Swinburne University of Technology, Mail H30 – PO Box 218, Hawthorn, VIC 3122, Australia

<sup>5</sup> Laboratoire d'Astrophysique de Marseille – LAM, Université d'Aix-Marseille & CNRS, UMR 7326, 38 rue F. Joliot-Curie, 13388 Marseille Cedex 13, France

<sup>6</sup> Department of Physics & Astronomy, University of California, Irvine, CA 92697, USA

<sup>7</sup> Division of Physics, Astronomy and Mathematics, California Institute of Technology, Pasadena, CA 91125, USA

<sup>8</sup> Zentrum für Astronomie der Universität Heidelberg, Institut für Theoretische Astrophysik, Albert-Ueberle Str. 2, 69120 Heidelberg, Germany

<sup>9</sup> School of Physics and Astronomy, Cardiff University, Queens Buildings, The Parade, Cardiff, CF24 3AA, UK

<sup>10</sup> CNRS, Institut d'Astrophysique Spatiale, UMR 8617, 91405 Orsay, France

<sup>11</sup> Université Paris Sud, Institut d'Astrophysique Spatiale, UMR 8617, 91405 Orsay, France

<sup>12</sup> Department of Physics & Astronomy, University of Sussex, Brighton, BN1 9QH, UK

<sup>13</sup> Laboratoire AIM, CEA, Université Paris VII, IRFU/Service d'Astrophysique, Bât. 709, 91191 Gif-sur-Yvette, France

<sup>14</sup> NASA Herschel Science Center, MS 100-22, California Institute of Technology, Pasadena, CA 91125, USA

<sup>15</sup> Istituto di Astrofisica e Planetologia Spaziali, INAF-IAPS, via Fosso del Cavaliere 100, 00133 Roma, Italy

<sup>16</sup> Department of Physics & Astronomy, McMaster University, Hamilton, Ontario, L8S 4M1, Canada

Received 4 August 2014 / Accepted 11 September 2014

### ABSTRACT

**Context.** Dust reprocesses about half of the stellar radiation in galaxies. The thermal re-emission by dust of absorbed energy is considered to be driven merely by young stars so is often applied to tracing the star formation rate in galaxies. Recent studies have argued that the old stellar population might be responsible for a non-negligible fraction of the radiative dust heating.

**Aims.** In this work, we aim to analyze the contribution of young ( $\lesssim 100$  Myr) and old ( $\sim 10$  Gyr) stellar populations to radiative dust heating processes in the nearby grand-design spiral galaxy M 51 using radiative transfer modeling. High-resolution 3D radiative transfer (RT) models are required to describe the complex morphologies of asymmetric spiral arms and clumpy star-forming regions and to model the propagation of light through a dusty medium.

**Methods.** In this paper, we present a new technique developed to model the radiative transfer effects in nearby face-on galaxies. We construct a high-resolution 3D radiative transfer model with the Monte-Carlo code SKIRT to account for the absorption, scattering, and non-local thermal equilibrium (NLTE) emission of dust in M 51. The 3D distribution of stars is derived from the 2D morphology observed in the IRAC 3.6  $\mu\text{m}$ , GALEX FUV, H $\alpha$ , and MIPS 24  $\mu\text{m}$  wavebands, assuming an exponential vertical distribution with an appropriate scale height. The dust geometry is constrained through the far-ultraviolet (FUV) attenuation, which is derived from the observed total-infrared-to-far-ultraviolet luminosity ratio. The stellar luminosity, star formation rate, and dust mass have been scaled to reproduce the observed stellar spectral energy distribution (SED), FUV attenuation, and infrared SED.

**Results.** The dust emission derived from RT calculations is consistent with far-infrared and submillimeter observations of M 51, implying that the absorbed stellar energy is balanced by the thermal re-emission of dust. The young stars provide 63% of the energy for heating the dust responsible for the total infrared emission (8–1000  $\mu\text{m}$ ), while 37% of the dust emission is governed through heating by the evolved stellar population. In individual wavebands, the contribution from young stars to the dust heating dominates at all infrared wavebands but gradually decreases towards longer infrared and submillimeter wavebands for which the old stellar population becomes a non-negligible source of heating. Upon extrapolation of the results for M 51, we present prescriptions for estimating the contribution of young stars to the global dust heating based on a tight correlation between the dust heating fraction and specific star formation rate.

**Key words.** radiative transfer – dust, extinction – galaxies: individual: M 51 – galaxies: ISM – infrared: galaxies

\* Appendix A is available in electronic form at  
<http://www.aanda.org>

## 1. Introduction

Although dust only makes up a small part of the interstellar material (typically 1% in mass), its impact on the other interstellar medium (ISM) constituents becomes important in several astrophysical processes and galaxy evolution. Dust particles act as a catalyst in the formation process of molecular hydrogen on the surfaces of dust particles (Hollenbach & Salpeter 1971). The smallest dust grains regulate the heating of the neutral ISM gas component through photoelectric heating (Bakes & Tielens 1994), and inelastic interactions with gas particles. The irregular shape of large dust grains supplies shielding for molecules and ions from the hard radiation of young stars, hereby promoting the congregation and recombination of several elements, molecules and ions.

Dust particles also distort our view on the other galaxy components (stars, gas, ...) through absorption and scattering processes. The location, distribution, and structure of dust clouds will control the propagation of light. An accurate description of the obscuration effects of dust is, therefore, necessary to determine the specific amount of processed light by dust grains which will be re-emitted at infrared(IR)/submillimeter(submm) wavelengths. Only a proper characterization of the effects of dust across the multi-wavelength spectrum will allow us to recover and interpret several fundamental properties such as the star formation history (SFH), star formation rate (SFR), initial mass function (IMF), etc. (e.g., Driver et al. 2007; Maller et al. 2009; da Cunha et al. 2010).

Other than corrections for dust obscuration to ultraviolet and optical data, it is also important to understand the different sources that heat the dust at infrared and submillimeter wavelengths. Especially because the far- and total-infrared emission in galaxies is often fully attributed to star formation, and, thus, used to derive an estimate of the star formation activity (e.g., Kennicutt et al. 2009). While the radiative cooling through thermal emission at infrared wavelengths is the dominant cooling mechanism for dust, the heating of dust can be powered by different mechanisms (see also Sauvage & Thuan 1994; Kewley et al. 2002; Boquien et al. 2011; Bendo et al. 2012a). Collisional heating, which dominates in hot gas where the electron velocities are high, and chemical heating, due to the formation of molecular hydrogen on the grain surfaces, are generally negligible on global galaxy scales compared to radiative heating mechanisms. Radiative heating can be driven by different stellar populations, ranging from young OB associations (Devereux & Young 1990; Devereux & Hameed 1997) over evolved stars with mass losing photospheres or circumstellar envelopes (Knapp et al. 1992; Mazzei & de Zotti 1994) to the general stellar radiation field. The latter emission spectrum is composed of energy escaping from active star-forming regions (Sauvage et al. 1990; Popescu et al. 2000; Misiriotis et al. 2001) as well as the underlying old stellar population (Helou 1986; Sauvage & Thuan 1994; Xu & Helou 1996; Li & Draine 2002; Boselli et al. 2004). The dust grains heated by the old stars and light escaping from star-forming regions shows up as large-scale diffuse dust emission, often denoted as cirrus. In high-energy environments, the dust might be heated mechanically through shocks and/or turbulence in the ISM (Miura & Nakamoto 2006), but its contribution is negligible on global galaxy scales. Other dust heating processes are powered by an active galactic nucleus (AGN) heating the surrounding dusty medium (e.g., de Grijp et al. 1985; Wu et al. 2007) or cosmic rays (e.g., Wolfire et al. 1995).

In the past, other works have employed surface brightness ratios of several infrared bands to study the main dust heating

mechanisms by looking at their correlation with several star formation rate indicators and total stellar mass tracers (Bendo et al. 2010, 2012a, 2014; Galametz et al. 2010; Boquien et al. 2011; Hughes et al. 2014). The study of a single infrared surface brightness ratio has the advantage of being mainly sensitive to dust temperature changes. The exploration of multi-waveband dust spectral energy distribution (SED) fitting techniques, however, benefits from the broader wavelength coverage to study the different heating sources in several infrared wavebands (e.g., Montalto et al. 2009; Aniano et al. 2012; Dale et al. 2012; Mentuch Cooper et al. 2012; Smith et al. 2012; Ciesla et al. 2014; Viaene et al. 2014). Panchromatic SED fitting methods are even capable of constraining the absorbed energy from star light and re-emitted photons from dust grains by postulating a dust energy balance (e.g., CIGALE, MAGPHYS and GRASIL among others).

The disadvantage of SED fitting is often the lack of stringent constraints on the 3D geometry of stars and dust in galaxies, and the ignorance on the propagation distance and direction of light within a galaxy. A complete self-consistent study of the dust heating mechanisms in galaxies requires high-resolution, 3D radiative transfer (RT) models (see Steinacker et al. 2013 for an overview on RT codes) accounting for the complex geometry of stars and dust in galaxies, and directly solving the equation describing the transfer of stellar light through a dusty medium (e.g., Popescu et al. 2000; Misiriotis et al. 2001; Law et al. 2011). Detailed RT models of spatially resolved galaxies, furthermore, provide the opportunity to analyze the spatial distribution of dust, the composition and emissivity of dust grains as well as possible variations throughout the galaxy. The main advantage of radiative transfer calculations is the non-local character of dust heating that can be addressed by tracing the propagation of stellar radiation through the dusty galaxy medium.

Up to the present day, RT modeling has been applied to a number of edge-on spiral galaxies (Misiriotis et al. 2001; Alton et al. 2004; Dasyra et al. 2005; Bianchi 2008; Baes et al. 2010; Popescu et al. 2011; De Looze et al. 2012a,b; Schechtman-Rook et al. 2012; De Geyter et al. 2013, 2014), which have the advantage that the extinction and emission of dust can easily be observed along the line-of-sight, and that the dust can be vertically resolved and traced out to large radii. One drawback of modeling edge-on galaxies is the lack of insight in the spatial distribution of star-forming regions and the clumpiness of the interstellar medium, which impedes a spatially resolved analysis of the various stellar populations and dust grains, and prevents analyzing the interplay between stars and dust to characterize the main dust heating mechanisms. To fully recover the asymmetric stellar and dust geometries, we require detailed radiative transfer calculations of well-resolved low-inclination systems, where the dust distribution can be studied into great detail, and the presence of unobscured/embedded star-forming regions can be identified. High-resolution 3D radiative transfer modeling of the Milky Way has allowed Robitaille et al. (2012) to identify a deficit in the dust distribution in the inner Galaxy and a potentially increased abundance of polycyclic aromatic hydrocarbons (PAHs). Owing to the improved spatial resolution and extended wavelength coverage down to the submillimeter regime provided by *Herschel* (Pilbratt et al. 2010), the radiative transfer modeling of stars and dust in nearby galaxies beyond our own Galaxy now becomes possible.

In this work, we aim to expand dust radiative transfer modeling of edge-on spirals galaxies to face-on counterparts. Due to the complex morphology with asymmetric spiral arms, bar and ring structures, obscured star-formation regions, and the clumpy

nature of the interstellar medium visible in face-on disk galaxies, we require a high-resolution 3D radiative transfer model that can simultaneously reproduce the multi-wavelength emission of stars and dust across the electromagnetic spectrum. The complexity of 3D radiative transfer models with their high number of unknown parameters, in combination with the computational cost of the full radiative transfer treatment, makes it a challenging task. Given that the asymmetric structures of face-on galaxies are hard to parameterize with analytic functions, and that the clumpy nature of localized sources within the galaxy's disk will be a hard nut to crack for most standard optimization techniques (e.g., genetic algorithms, Levenberg-Marquardt minimization, etc.), we present an alternative modeling technique that will benefit from high-resolution imaging observations to constrain the stellar and dust geometries in the radiative transfer model.

In this paper, we construct a 3D radiative transfer model of the grand-design spiral galaxy M 51<sup>1</sup> (NGC 5194) with a near to face-on orientation ( $i \sim 20^\circ$ ) and a huge database of multi-band ancillary data. We exploit the panchromatic RT model of stars and dust in M 51 to analyze the main dust heating sources on spatially resolved scales across the infrared spectrum. Rather than giving hints on the predominant dust heating source, the RT calculations track the emission of stellar populations of different ages and their interaction with the dusty interstellar medium, which allows to constrain the percentage of dust heated by young stars ( $\lesssim 100$  Myr) and the old stellar population ( $\sim 10$  Gyr) at every location in 3D space across multiple infrared wavebands.

The paper is organized with the description of the multi-wavelength dataset presented in Sect. 2. In Sect. 3, we introduce a new method for radiative transfer modeling to constrain the geometry of stars and dust in face-on galaxies based on observations in ultra-violet (UV), optical, near-infrared (NIR) and infrared (IR) wavebands. We discuss the strengths and possible flaws of the constructed RT model in Sect. 4. Section 5 applies the RT model to analyze the main dust heating mechanisms in M 51 and the variation in dust heating sources with infrared wavebands. The main conclusions of the paper are summed up in Sect. 6. In Appendices A.1 and A.2, we investigate the implications introduced by the model assumptions on the relative dust-to-stellar scale height and clumpiness of the dust distribution, respectively, for the RT model of M 51.

## 2. Multi-wavelength dataset

The paper aims to construct a multi-wavelength radiative transfer model for M 51 that can self-consistently explain the stellar emission, dust attenuation, and dust emission. We, hereto, present a new RT modeling technique that employs observations to characterize the geometrical distribution of stars and dust in a 3D model. To compare the model output images with observations, we require a multi-wavelength dataset. The observational dataset used for model construction (see Sect. 3.2) and model validation (see Sect. 4) is described in Sect. 2.1, while the necessary data pre-processing steps are outlined in Sect. 2.2.

<sup>1</sup> M 51 is a nearby, late-type (SABc) galaxy located at a distance of  $D = 8.4$  Mpc (Feldmeier et al. 1997), and belongs to an interacting system with the companion galaxy NGC 5195. The galaxy pair is often referred to as M 51, while the individual galaxy NGC 5194 is named M 51a. In this paper, we refer to M 51 for the single galaxy NGC 5194 belonging to the interacting system.

### 2.1. Dataset

Ultra-violet (UV) maps were obtained with the Galaxy Evolution Explorer (GALEX, Martin et al. 2005) with a total observing time of 1414 s as part of the GALEX Nearby Galaxy Survey (NGS, Bianchi et al. 2003a,b; Gil de Paz et al. 2004), a program collecting deep GALEX FUV and NUV data for about 200 nearby galaxies. We collected GALEX FUV and NUV maps provided by Gil de Paz et al. (2007) from the NASA Extragalactic Database (NED)<sup>2</sup>.

H $\alpha$  data were retrieved from Boselli & Gavazzi (2002). The galaxy was observed with the 1.20 m Newton telescope at the Observatoire de Haute Provence using a TK 1024  $\times$  1024 pixel CCD detector and narrowband filters at 6450 and 6561 Å with exposure times of 900 s each. After standard data processing steps were applied to the data, the 6450 Å image was subtracted from the 6561 Å image to produce the H $\alpha$  image. Foreground stars were identified as negative point-like residual features, and were removed by interpolating over the data. The final image has pixels with sizes of 0.69'', a calibration uncertainty of 5%, and a PSF FWHM of 2''.

M 51 was also observed as part of the Sloan Digital Sky Survey (SDSS) in two separate frames (each with an exposure time of 54 s) using the  $u$ ,  $g$ ,  $r$ ,  $i$  and  $z$  filters. We downloaded the two image tiles from the SDSS DR7 archive (Abazajian et al. 2009). The two separate image frames are background subtracted before they are combined to final mosaics based on standard iraf image processing tasks (imalign, imcombine).

The Two Micron All Sky Survey (2MASS, Jarrett et al. 2003) observed M 51 in the  $J$ ,  $H$ ,  $K_s$  filters as part of the Large Galaxy Atlas Survey. We collected 2MASS  $J$ ,  $H$ ,  $K_s$  maps provided by Jarrett et al. (2003) from NED.

The Wide-field Infrared Survey Explorer (WISE, Wright et al. 2010) observed M 51 in its four photometric bands at 3.4, 4.6, 11.6 and 22.0  $\mu\text{m}$ . WISE maps are converted from DN units to Jy based on the photometric zeropoints quoted in the User's Guide to the WISE Preliminary Data Release<sup>3</sup>. Three types of corrections are recommended for extended source photometry by Jarrett et al. (2013). The first aperture correction factor corrects for the PSF profile fitting that is used for the WISE absolute photometric calibration, with corrections of 0.034, 0.041, -0.030 and 0.029 mag applied to the WISE 3.4, 4.6, 11.6 and 22.0  $\mu\text{m}$  filters. The second correction corrects for the shape of the spectrum. We do not account for this color correction since the SKIRT output spectra are convolved with the appropriate WISE response filters. The last correction factor only applies to the WISE 4 band (0.92), correcting for the calibration discrepancy between WISE photometric standard blue stars and red galaxies.

The Infrared Array Camera (IRAC, Fazio et al. 2004) and Multiband Imaging Photometer (MIPS, Rieke et al. 2004) on board the *Spitzer* Space telescope observed M 51 as part of the *Spitzer* Infrared Nearby Galaxy Survey (SINGS, Kennicutt et al. 2003). We retrieved IRAC 3.6, 4.5, 5.8 and 8.0  $\mu\text{m}$  maps from the SINGS data archive<sup>4</sup>. We, furthermore, retrieve an optical  $R$  band image from the SINGS data archive. The IRAC images were multiplied with the appropriate extended source photometrical correction coefficients (0.91, 0.94, 0.66 and 0.74 for the IRAC 3.6, 4.5, 5.8 and 8  $\mu\text{m}$  wavebands; see the IRAC

<sup>2</sup> <http://ned.ipac.caltech.edu/>

<sup>3</sup> [http://wise2.ipac.caltech.edu/docs/release/prelim/expSUP/sec2\\_3f.html](http://wise2.ipac.caltech.edu/docs/release/prelim/expSUP/sec2_3f.html)

<sup>4</sup> <http://irsa.ipac.caltech.edu/data/SPITZER/SINGS/>

instrument cookbook<sup>5</sup>). We use the reprocessed MIPS 24  $\mu\text{m}$  data from Bendo et al. (2012b).

M51 is observed by the *Herschel* Space Observatory (Pilbratt et al. 2010) as part of the Very Nearby Galaxy Survey (VNGS, PI: Christine Wilson). The PACS (70, 160  $\mu\text{m}$ , ObsID 1342188328, 1342188329) and SPIRE (250, 350, 500  $\mu\text{m}$ , ObsID 1342188589) instruments onboard *Herschel* performed continuum mapping of the interacting galaxy pair NGC 5194/NGC 5195. PACS and SPIRE data of M 51 were observed on December 20th and 26th 2010, covering an area of about  $20' \times 20'$  centered on the galaxy M 51. The same area was observed in nominal and orthogonal scan direction with four and two repetitions with PACS and SPIRE instruments, respectively, at a medium scan speed (20 arcsec  $\text{s}^{-1}$ ). The *Herschel* dataset was already presented in Mentuch Cooper et al. (2012). For the analysis in this paper, PACS data were retrieved from the *Herschel* Science Archive (HSA) and re-processed using the *Herschel* Interactive Processing Environment (HIPE, v12, Balm 2012) up to level 1 using the calibration file v56. The level 1 data are reduced in version 23 of scanamorphos (Roussel 2013). The SPIRE data reduction was performed in a similar way by processing the data up to level 1 using the standard pipeline and calibration file version `spire_cal_12_1` in HIPE v11 and reducing the level 1 data using the BRiGht Galaxy ADaptive Element method (Smith et al., in prep.). The latter uses a custom method to remove the temperature drift, and bring all bolometers to the same level (instead of the default *temperatureDriftCorrection* and the residual, median baseline subtraction). The SPIRE maps are multiplied with  $K_{\text{PtoE}}$  correction factors to convert from point source to extended source photometric calibration. The appropriate correction factors (i.e., 91.289, 51.799, 24.039 at 250, 350 and 500  $\mu\text{m}$ , respectively; see SPIRE Observers' manual) for a constant  $\nu S_\nu$  spectrum convert the maps from Jy/beam to MJy/sr units, accounting for the most up-to-date measured SPIRE beam areas of 465, 823, 1769 arcsec<sup>2</sup> at 250, 350 and 500  $\mu\text{m}$ , respectively. No color corrections are applied since the SKIRT output SED and images are convolved with the appropriate response curve to derive modeled band fluxes.

The full-width-half-maximum (FWHM) of the point-spread-function (PSF) in each of the different bands as well as the calibration uncertainties are summarized in Table 1 along with the appropriate references reporting on these numbers.

## 2.2. Data pre-processing

Foreground stars are identified from the 2MASS Point Source Catalog. Similar to the correction techniques applied in Mentuch Cooper et al. (2012), we only identify point sources as stars (and not HII regions) if the surface brightness ratio of the IRAC 8  $\mu\text{m}$  band versus  $R$  band emission,  $I_8/I_R$ , is lower than 1.5. Point sources satisfying this criterion are removed, i.e., their emission is replaced by the median pixel value in the background annulus with inner radius 8'' and outer radius 16'' around the star's position.

The background in each of the maps has been determined by measuring the background in a sufficiently large number (depending on the size of the background region available) of apertures of size  $R = 4 \times \text{FWHM}$ , randomly distributed outside the bright emission regions of M 51. The final background value is chosen as the median value of those measurements, with the standard deviation being representative for the uncertainty on

**Table 1.** Overview of the characteristics (full-width half maximum of the beam, calibration uncertainties) for the multi-wavelength dataset used in the RT analysis of M 51.

Band	$\lambda$ [ $\mu\text{m}$ ]	FWHM [arcsec]	Ref <sup>a</sup>	$\sigma_{\text{cal}}$	Ref <sup>a</sup>
GALEX FUV	0.15	4.2	1	0.05 mag	1
GALEX NUV	0.23	5.3	1	0.03 mag	1
SDSS <i>u</i>	0.35	2 <sup>b</sup>	2	2%	3
SDSS <i>g</i>	0.48	2 <sup>b</sup>	2	2%	3
SDSS <i>r</i>	0.62	2 <sup>b</sup>	2	2%	3
H $\alpha$	0.66	2	4	5%	4
SDSS <i>i</i>	0.76	2 <sup>b</sup>	2	2%	3
SDSS <i>z</i>	0.91	2 <sup>b</sup>	2	2%	3
2MASS <i>J</i>	1.25	3.3 <sup>c</sup>	2	0.03 mag	5
2MASS <i>H</i>	1.65	3.1 <sup>c</sup>	2	0.03 mag	5
2MASS <i>K</i>	2.16	3.3 <sup>c</sup>	2	0.03 mag	5
WISE 1	3.4	6.1	6	2.4% <sup>d</sup>	7
IRAC 3.6 $\mu\text{m}$	3.6	1.7	8	1.8%	9
IRAC 4.5 $\mu\text{m}$	4.5	1.7	8	1.9%	9
WISE 2	4.6	6.4	6	2.8%	7
IRAC 5.8 $\mu\text{m}$	5.8	1.9	8	2.0%	9
IRAC 8.0 $\mu\text{m}$	8.0	2.0	8	2.1%	9
WISE 3	12	6.5	6	4.5%	7
WISE 4	22	12	6	5.7%	7
MIPS 24	24	6	10	4%	10
PACS 70	70	5.8	11	5%	12
PACS 160	160	12.1	11	5%	12
SPIRE 250	250	18.2	13	4% <sup>d</sup>	13
SPIRE 350	350	24.9	13	4% <sup>d</sup>	13
SPIRE 500	500	36.3	13	4% <sup>d</sup>	13

**Notes.** <sup>(a)</sup> References: (1) Morrissey et al. (2007); (2) this paper; (3) Padmanabhan et al. (2008); (4) Boselli & Gavazzi (2002); (5) Skrutskie et al. (2006); (6) Wright et al. (2010); (7) Jarrett et al. (2011); (8) Fazio et al. (2004); (9) Reach et al. (2005); (10) Engelbracht et al. (2007); (11) PACS Observers' Manual; (12) Balog et al. (2013); (13) SPIRE Observers' Manual. <sup>(b)</sup> We estimate the FWHM of the SDSS data by fitting a 2D Gaussian (IDL task `starfit`) to several stellar objects in the field of M 51. We apply the sum of two Gaussian kernels with FWHM of 2'' to convolve the SDSS data, which should be a good representation of the telescope's PSF (Aniano et al. 2011). <sup>(c)</sup> We estimate the FWHM of the 2MASS data by fitting a 2D Gaussian (IDL task `starfit`) to several stellar objects in the field of M 51. We apply Gaussian kernels with FWHM of 3'' (*H*) and 3.5'' (*J, K*) to convolve the 2MASS data to the PACS 160  $\mu\text{m}$  resolution (Aniano et al. 2011). <sup>(d)</sup> Calibration uncertainties for the *Herschel* SPIRE instrument are assumed to be around 4% in each band, adding in quadrature the 4% absolute calibration error from the assumed models used for Neptune (SPIRE Observers' manual), and a random uncertainty of 1.5% accounting for the repetitive measurements of Neptune (see Bendo et al. 2013).

the background subtraction procedure. The background value has been subtracted from the frames in case the median background value exceeded the standard deviation.

All maps are corrected for Galactic extinction according to the re-calibrated  $A_\lambda$  in Schlafly & Finkbeiner (2011) from Schlegel et al. (1998), as reported on NED. The attenuation in the ultraviolet wavebands is determined based on the *V*-band attenuation  $A_V$  and assuming an extinction law with  $R_V$  derived in Fitzpatrick (1999).

To homogenize the observational dataset, and reconcile the resolution of the input geometries in SKIRT, we convolve the entire dataset to the resolution of the PACS 160  $\mu\text{m}$  image ( $\sim 12.1''$  or 493 pc at the distance of M 51), which is the lowest resolution map that will be used to constrain stellar and dust geometries in

<sup>5</sup> [http://irsa.ipac.caltech.edu/data/SPITZER/docs/irac/iracinstrument-handbook/IRAC\\_Instrument\\_Handbook.pdf](http://irsa.ipac.caltech.edu/data/SPITZER/docs/irac/iracinstrument-handbook/IRAC_Instrument_Handbook.pdf)

SKIRT<sup>6</sup>. The appropriate convolution kernels are retrieved from Aniano et al. (2011). All convolved images are rebinned to maps with pixel size of 2". With a pixel size six times smaller compared to the FWHM of the PSF in the PACS 160 μm waveband (~12.1"), adjacent pixels will not be independent in the input images describing the stellar and dust geometries in SKIRT. Given that the non-local interactions between stellar photons and dust particles are inherent to radiative transfer calculations, the possible dependence of adjacent pixels in the input geometries will not severely affect the RT models.

To make the size of each input image consistent, we restrict the emission in the PACS 160 μm to pixels detected at a signal-to-noise level of at least five, resulting in 35 375 pixels of size 2 × 2 arcsec<sup>2</sup>. The emission in the other input images is cut at the same contour level to prevent that the extent of the SKIRT input images, and, thus, the radial distribution differs among RT components.

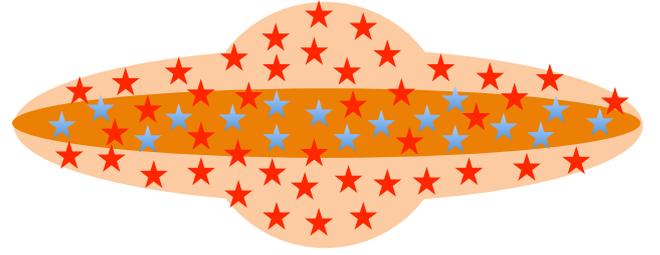
### 3. 3D radiative transfer model

#### 3.1. Radiative transfer code: SKIRT

SKIRT (Baes et al. 2003, 2011) is a 3D Monte Carlo radiative transfer code designed to model the absorption, scattering, and thermal re-emission of dust in a variety of environments: circumstellar disks (Vidal & Baes 2007), clumpy tori around active galactic nuclei (Stalevski et al. 2012), and different galaxy types (Baes et al. 2010; Gadotti et al. 2010; Gomez et al. 2010; De Looze et al. 2012a,b; De Geyter et al. 2013). SKIRT can furthermore be applied to model the emission of simulated galaxies from hydrodynamical simulations (Camps et al. 2013a). An important extension of SKIRT in view of high-resolution radiative transfer models involves the implementation of adaptive grid structures such as the hierarchical Octree (Kurosawa & Hillier 2001; Niccolini & Alcolea 2006), *k*-d grid structures, and Voronoi tessellations (Voronoi 1908), which are described in more detail in Saffly et al. (2013, 2014) and Camps et al. (2013b), respectively. Such adaptive grid structures allow to increase the resolution in the dust grid where dust clumps or asymmetric spiral arm structures are present, and at the same time keep the computational cost limited.

#### 3.2. Model construction

The construction of a 3D radiative transfer model composed of stars and dust includes the characterization of the *geometrical distribution*, *spectrum*, and *normalization* of each of the different stellar and dust components. The geometrical distribution of SKIRT components is described through 3D geometries that render the probability density distribution for stars or dust (see Sect. 3.2.1). These 3D density distributions give the probability that a given location in 3D space will be populated by dust particles or stars of a certain age. The description of the emission spectra of stars characterizes the variation of stellar emission with wavelength (see Sect. 3.2.2), while the normalization of the spectra encompasses the scaling of this emission spectrum for a given bolometric luminosity, or a luminosity specified in a given waveband (see Sect. 3.2.3). The emission spectrum for the dust component in our RT calculations is constrained through



**Fig. 1.** A schematic diagram of our standard model observed from an edge-on view, consisting of a bulge and thick disk component (represented with a light orange color) of old stars (red stars), and a thin star-forming disk (dark orange) containing a dust component and young stars (blue stars).

the grain composition, with the grain size distribution, optical dust properties, and dust emissivities specified for each individual grain population (see Sect. 3.2.2), and normalized by the total dust mass (see Sect. 3.2.3).

#### 3.2.1. SKIRT model distribution: 3D stellar and dust geometries

The model construction for M 51 is based on the standard model represented in Fig. 1. The standard model is composed of a bulge, thick and thin disk, with the thin disk harboring the dust and young stars, while the bulge and thick disk consist merely of old stars. The standard model represented in Fig. 1 has been applied, whether or not with some small modifications, in several other radiative transfer studies (e.g., Xilouris et al. 1999; Popescu et al. 2000, 2011; Pierini et al. 2004; Tuffs et al. 2004; Bianchi 2008), and has been proven to give a realistic representation of the geometry of stars and dust in galaxies, capable of reproducing the observed UV/optical and IR/submm emission. While the standard model also applies to model edge-on galaxies, the galaxies observed from a face-on perspective require 3D asymmetric geometries to describe the observed stellar and dust structures, which is different from the 2D analytic prescriptions used to describe the stars and dust in edge-on galaxies.

The geometrical distribution of the stellar and dust components in the 3D radiative transfer model of M 51 is constrained based on a set of multi-wavelength observations. Although observations only depict a 2D projection of the 3D galaxy, we are able to recover the 3D geometry through de-projection following an Euler-like transformation described by the equation:

$$\begin{pmatrix} x_p \\ y_p \\ z_p \end{pmatrix} = \begin{pmatrix} \cos \omega & -\sin \omega & 0 \\ \sin \omega & \cos \omega & 0 \\ 0 & 0 & 1 \end{pmatrix} \begin{pmatrix} 0 & 1 & 0 \\ -1 & 0 & 0 \\ 0 & 0 & 1 \end{pmatrix} \times \begin{pmatrix} \cos \theta & 0 & -\sin \theta \\ 0 & 1 & 0 \\ \sin \theta & 0 & \cos \theta \end{pmatrix} \begin{pmatrix} \cos \phi & -\sin \phi & 0 \\ \sin \phi & \cos \phi & 0 \\ 0 & 0 & 1 \end{pmatrix} \begin{pmatrix} x \\ y \\ z \end{pmatrix} \quad (1)$$

with an inclination angle  $\theta \sim 20^\circ$ , position angle  $\omega \sim 170^\circ$  (Tully 1974), azimuth  $\phi \sim 0^\circ$ , and the projected *z* coordinate,  $z_p = 0$ . To describe the vertical *z* distribution, we assume an exponential distribution  $\exp(-|z|/h_z)$  with a certain scale height,  $h_z$ .

Due to the nearly face-on view on M 51 ( $i = 20^\circ$ ), we lack any constraints on the vertical scale height of stars and dust. Therefore, we need to rely on studies of the vertical extent of stars and dust in highly-inclined disk galaxies, under the assumption that the relative star-dust geometry in M 51 is similar to the observed vertical structures in high-inclination systems. The scale heights have been shown to be comparable for stars

<sup>6</sup> We prefer to convolve all data to the resolution of the PACS 160 μm map to maintain the spatial structures at scales of ~500 pc rather than convolving all data to the resolution in the SPIRE wavebands. For comparison between SPIRE modeled and observed data, we convolve the model images to the appropriate SPIRE resolution.

and dust in low mass galaxies (Seth et al. 2005), but for high mass galaxies the ISM is considered to collapse into dust lanes distributed in thin disks (Dalcanton et al. 2004). Also model fitting to several edge-on galaxies have resulted in smaller dust scale heights compared to the vertical extent of the thick disk composed of old stars (Jansen et al. 1994; Xilouris et al. 1999; Bianchi 2007; De Geyter et al. 2013, 2014; Verstappen et al. 2013). The scale height of dust in galaxies is, however, still debated with several galaxies showing evidence of extra-planar dust (e.g., Howk & Savage 1999; Alton et al. 2000; Thompson et al. 2004). While this extra-planar dust component is interpreted as infalling or outflowing material, the presence of small dust grains and PAHs traced up to several kpc above the plane of several edge-on spiral galaxies suggests that a diffuse dust component with scale heights exceeding the vertical scale height of more commonly observed stellar and dust disks might be present in galaxies, and, thus, not of external origin (Irwin & Madden 2006; Irwin et al. 2007; Kamphuis et al. 2007; Whaley et al. 2009). The divergent angular resolution of observations at various IR/submm wavebands furthermore complicates the determination of dust scale heights. On top of the resolution issues, the intrinsic dust distribution is tough to define from the emission in a single waveband due to the different heating sources that might contribute at individual wavebands.

Consistent with observational constraints and modeling results obtained for edge-on spiral galaxies, we assume that the old stars and dust are distributed in a disk with an exponential vertical profile with scale heights  $h_{z,*} \sim 450$  pc and  $h_{z,d} \sim 225$  pc, respectively. The scale height of old stars in M 51 ( $h_{z,*} \sim 450$  pc) is based on the scale length fitted to the observed  $R$ ,  $I$  and  $K$  band images ( $h_{R,*} \sim 90''$  or 3665 pc, Beckman et al. 1996), and the mean intrinsic flattening of the stellar disk as obtained from 2D model decomposition ( $h_{R,*}/h_{z,*} \sim 8.21$ , Kregel et al. 2002) and 3D RT modeling ( $h_{R,*}/h_{z,*} 8.26$ , De Geyter et al. 2014). We assume a relative stellar-to-dust scale height ratio of 2:1, consistent with the RT model predictions for edge-on galaxies from Xilouris et al. (1999); Bianchi (2007) and De Geyter et al. (2014), resulting in a dust scale height of 225 pc. We study the effect of variations in the relative dust-to-stellar scale heights in Sect. A.1, which shows that changing the relative dust-to-stellar scale height results in an inconsistency with the observed FUV attenuation and IR emission. We are, therefore, confident that the assumed relative dust-to-stellar scale height  $h_{z,d}/h_{z,*} = 0.5$  provides a good representation of the vertical distribution of stars and dust in M 51. The comparison of 3D radiative transfer models with variable dust-to-stellar scale heights suggests that the observed FUV attenuation is able to put stringent constraints on the relative dust-to-stellar vertical distribution in galaxies. This new technique of determining the relative dust-to-stellar scale height will be exploited for a larger sample of nearby galaxies in future work.

The young stars (both ionizing and non-ionizing) are assumed to be distributed in a thin disk with scale height  $h_z \sim 100$  pc, which is compatible with the scale height of star-forming disks in our Galaxy (Bahcall & Soneira 1980) and NGC 891 (Schechtman-Rook & Bershadsky 2013). A power spectrum analysis of the dust emission at 24 and  $100 \mu\text{m}$  in M 33, furthermore, results in similar dust scale heights of 100 pc for dust components heated by young stars (Combes et al. 2012).

The scale heights of stellar and dust disks are assumed to be invariant throughout the disk of M 51. The stellar scale height has been shown to increase with galactocentric distance in some galaxy disks (e.g., de Grijs & Peletier 1997; Narayan & Jog 2002), but this variation should mainly occur in early-type

**Table 2.** Overview of the different stellar and dust components in the RT model of M 51.

Component	Parameter	Value
Old stars		
Bulge <sup>a</sup>	$n$	0.67
	$R_e$ [pc]	635.3
	$q$	0.88
	$L_V$ [ $L_{\odot,V}$ ]	$3.2 \times 10^9$
Thick disk <sup>b</sup>	2D geometry	IRAC 3.6 $\mu\text{m}$ <sup>c</sup>
	$h_z$ [pc]	450
	$L_V$ [ $L_{\odot,V}$ ]	$2.0 \times 10^{10}$
Young stars (non-ionizing)		
Thin disk <sup>d</sup>	2D geometry	GALEX FUV <sup>e</sup>
	$h_z$ [pc]	100
	SFR [ $M_{\odot} \text{ yr}^{-1}$ ]	3
Young stars (ionizing)		
Thin disk <sup>f</sup>	2D geometry	$\text{H}\alpha + 0.031 \times \text{MIPS } 24 \mu\text{m}^g$
	$h_z$ [pc]	100
	SFR [ $M_{\odot} \text{ yr}^{-1}$ ]	3
	$M_d$ [ $M_{\odot}$ ]	$4.5 \times 10^6$
Dust		
Thin disk <sup>h</sup>	2D geometry	$A_{\text{FUV}}^i$
	$h_z$ [pc]	225
	$M_d$ [ $M_{\odot}$ ]	$7.3 \times 10^7$

**Notes.** <sup>(a)</sup> The bulge component is modeled as a flattened Sérsic parametrized through the effective radius  $R_e$ , Sérsic index  $n$ , intrinsic flattening  $q$  and  $V$  band luminosity  $L_V$ . <sup>(b)</sup> The thick disk is modeled using observations to constrain the 2D geometry, and assuming an exponential distribution with scale height  $h_z$  in the vertical direction. The  $V$  band luminosity,  $L_V$ , determines the normalization of the stellar emission spectrum. <sup>(c)</sup> Corrected for the bulge emission. <sup>(d)</sup> The thin disk of non-ionizing stars is modeled using observations to constrain the 2D geometry, and assuming an exponential distribution with scale height  $h_z$  in the vertical direction. The star formation rate, SFR, determines the normalization of the stellar emission spectrum. <sup>(e)</sup> Observed FUV emission, not corrected for internal dust attenuation. <sup>(f)</sup> The thin disk of ionizing stars is modeled using observations to constrain the 2D geometry, and assuming an exponential distribution with scale height  $h_z$  in the vertical direction. The star formation rate, SFR, determines the normalization of the stellar+PDR emission spectrum. The dust mass,  $M_d$ , signifies the amount of grains present in the shells around young stellar clusters, associated with the PDR models. <sup>(g)</sup> The MIPS 24  $\mu\text{m}$  image is corrected for the emission of a diffuse dust component. <sup>(h)</sup> The thin dust disk is modeled using observations to constrain the 2D geometry, and assuming an exponential distribution with scale height  $h_z$  in the vertical direction. The dust mass,  $M_d$ , is used to normalize the amount of grains in the RT model. <sup>(i)</sup> Constructed based on the prescriptions from Cortese et al. (2008).

galaxies, and become negligible in late-type spirals (de Grijs & Peletier 1997). The recent interaction of NGC 5194 with companion NGC 5195 (Salo & Laurikainen 2000; Dobbs et al. 2010) might, however, have caused twisting or warping of the disk (e.g., Shetty et al. 2007), resulting in a changing position angle and/or inclination with radius.

Hereafter, we specify the observational constraints used to describe the geometrical distribution of the different stellar and dust components in the standard model (see also Table 2).

#### Old stellar population

The old stellar component is constrained by the observed IRAC 3.6  $\mu\text{m}$  emission. However, we need to separate the emission originating from the bulge and disk components in this IRAC 3.6  $\mu\text{m}$  map. Hereto, we model the stellar bulge using

our radiative transfer code SKIRT, and assuming the best fitting bulge parameters reported by Peng et al. (2010), based on their 2D model decomposition of M51 derived with the galaxy fitting algorithm galfit (Peng et al. 2002). More specifically, the emission of the bulge is modeled using a flattened Sersic profile with a Sersic index  $n = 0.67$ , effective radius  $R_e = 635.3$  pc and flattening parameter  $q = 0.88$ , assuming an inclination angle of  $20^\circ$  and position angle of  $170^\circ$ . We scale the luminosity of the bulge until we reproduce the bulge-to-disk ratio (0.16) of the model for M51 presented by Peng et al. (2010). After subtraction of the bulge component, the IRAC  $3.6\ \mu\text{m}$  map is considered to only include the emission of old stars distributed in the thick disk.

#### Young non-ionizing stars

We rely on the observed GALEX FUV image, dominated by emission from B and more massive A stars tracing the unobscured star formation over time scales of 10–100 Myr (Meurer et al. 1999; Boselli et al. 2009), to constrain the geometrical distribution of the young non-ionizing stellar component in M51.

#### Young ionizing stars

We rely on the observed  $H\alpha$  and MIPS  $24\ \mu\text{m}$  maps to describe the 2D distribution of young ionizing stars. Given that part of the  $24\ \mu\text{m}$  emission might result from diffuse dust heated by the interstellar radiation field rather than star-forming regions, we need to correct for this diffuse dust component contributing to the  $24\ \mu\text{m}$  emission. We estimate that the global emission of diffuse regions contributing to the  $24\ \mu\text{m}$  emission amounts to about 20% – assuming that the contribution is similar to the diffuse dust component observed in the Scd galaxy M33 (see Xilouris et al. 2012)<sup>7</sup>. We model an exponential disk with scale length  $h_R = 2000$  pc (following Xilouris et al. 2012) with the radiative transfer code SKIRT, scale it to 20% of the global  $24\ \mu\text{m}$  flux of M51, and subtract the diffuse disk component from the observed  $24\ \mu\text{m}$  map. A map of young ionizing stars is constructed based on the prescriptions from Calzetti et al. (2007)<sup>8</sup> to correct the  $H\alpha$  emission for internal extinction, i.e.,  $L_{H\alpha, \text{corr}} = L_{H\alpha} + 0.031 \times L_{24}$ .

#### Dust component

The geometry of the dust component in M51 is constrained through the FUV attenuation. The FUV attenuation is computed based on the calibrations presented by Cortese et al. (2008) through polynomials expressing a relationship between  $A_{\text{FUV}}$ , and the ratio of the total-infrared (TIR) and the FUV luminosity. The advantage of using the TIR-to-FUV ratio to determine the FUV attenuation is the independence of the method on the relative dust-star geometry and the assumed extinction law (Buat & Xu 1996; Meurer et al. 1999; Gordon et al. 2000; Witt & Gordon 2000). The total-infrared luminosity ( $L_{\text{TIR}}$ ) on spatially resolved scales in M51 is computed from the MIPS  $24\ \mu\text{m}$ , PACS  $70\ \mu\text{m}$  and  $160\ \mu\text{m}$  maps (convolved to the PACS  $160\ \mu\text{m}$  resolution, and rebinned to a grid with pixel size of  $2''$ ) and the calibrations presented by Galametz et al. (2013). The regression coefficients from Cortese et al. (2008) are chosen as a function of the specific star formation rate (sSFR), as traced by different colors,

<sup>7</sup> Based on the contribution of old stars to the  $24\ \mu\text{m}$  (48%) emission derived for the Sab galaxy M81 (Lu et al. 2014), the diffuse emission component emitting at  $24\ \mu\text{m}$  might become higher in spiral galaxies with an earlier-type classification and, thus, possibly be also higher in M51 compared to the 20% derived for M33.

<sup>8</sup> We prefer to rely on the calibration coefficient of 0.031 reported by Calzetti et al. (2007) in combination with the diffuse emission-subtracted MIPS  $24\ \mu\text{m}$  map rather than the factor 0.02 given by Kennicutt et al. (2009) which accounts for the contribution from diffuse emission heated by old stellar populations.

to account for differences in the mean age of the stellar population, which will affect the calculation of  $A_{\text{FUV}}$ <sup>9</sup>. For M51, we constrain the age dependence based on FUV –  $H$  colors (see Fig. 2, left panel). The lowest FUV –  $H$  colors correspond to high sSFRs, which are mainly found in the outer spiral arm regions located on the nearside of the interaction with NGC 5195. Similarly, Kaleida & Scowen (2010) found an enhancement in the number of stellar associations in the northern spiral arm closest to the companion, which suggest the triggering of star formation by the interaction. The right panel of Fig. 2 shows the  $A_{\text{FUV}}$  map derived following the procedures in Cortese et al. (2008). The FUV attenuation clearly peaks in the center of M51, but also the spiral arm regions can be identified as regions of high extinction. The PACS  $70\ \mu\text{m}$  image contours correspond reasonably well with the observed features in the FUV attenuation map. However, most of the extinction occurs in the inner regions of the spiral arms, which either suggests that most of the material is located on the inside of the spiral arm, or, alternatively, that the star-forming regions are more obscured there.

This FUV attenuation map is used to constrain the geometrical distribution of dust by de-projection to a 3D geometry, and assuming an exponential distribution in the  $z$  direction. The exponential distribution of the dust has been proven adequate to describe the vertical distribution of dust in a galaxy’s disk (e.g., Verstaappen et al. 2013; Hughes et al. 2014). The effect of a two-phase dust medium with clumpy dust features is studied and discussed in Sect. A.2, which shows that including the clumping of dust (in addition to the 2D asymmetric geometry derived from the FUV attenuation map) causes an underestimation of the FUV attenuation. We, therefore, argue that an additional component of quiescent compact dust clouds or clumps is not necessary to describe the dust component in the RT model of M51.

The resolution of the radiative transfer simulations is set by the size of individual grid cells, i.e., the physical area with constant dust and stellar mass fractions, and properties. We, therefore, need to make sure to sufficiently sample the different regions within the galaxy to obtain reliable results from the RT calculations. We use a k-d tree grid structure, following the prescriptions from Saftly et al. (2014). The  $x \times y \times z$  dimensions of the grid are  $20 \times 20 \times 5\ \text{kpc}^3$ , and the size of individual cells ranges from  $40^3$  to  $312^3\ \text{pc}^3$ . With an average optical depth  $\tau_V \sim 0.09$  and maximum value  $\tau_V \sim 0.33$  for the total number of grid cells (1 550 066), we are confident that the dust grid is well sampled, and will produce reliable results simulating the interaction between stellar light and dust.

### 3.2.2. SKIRT model spectra

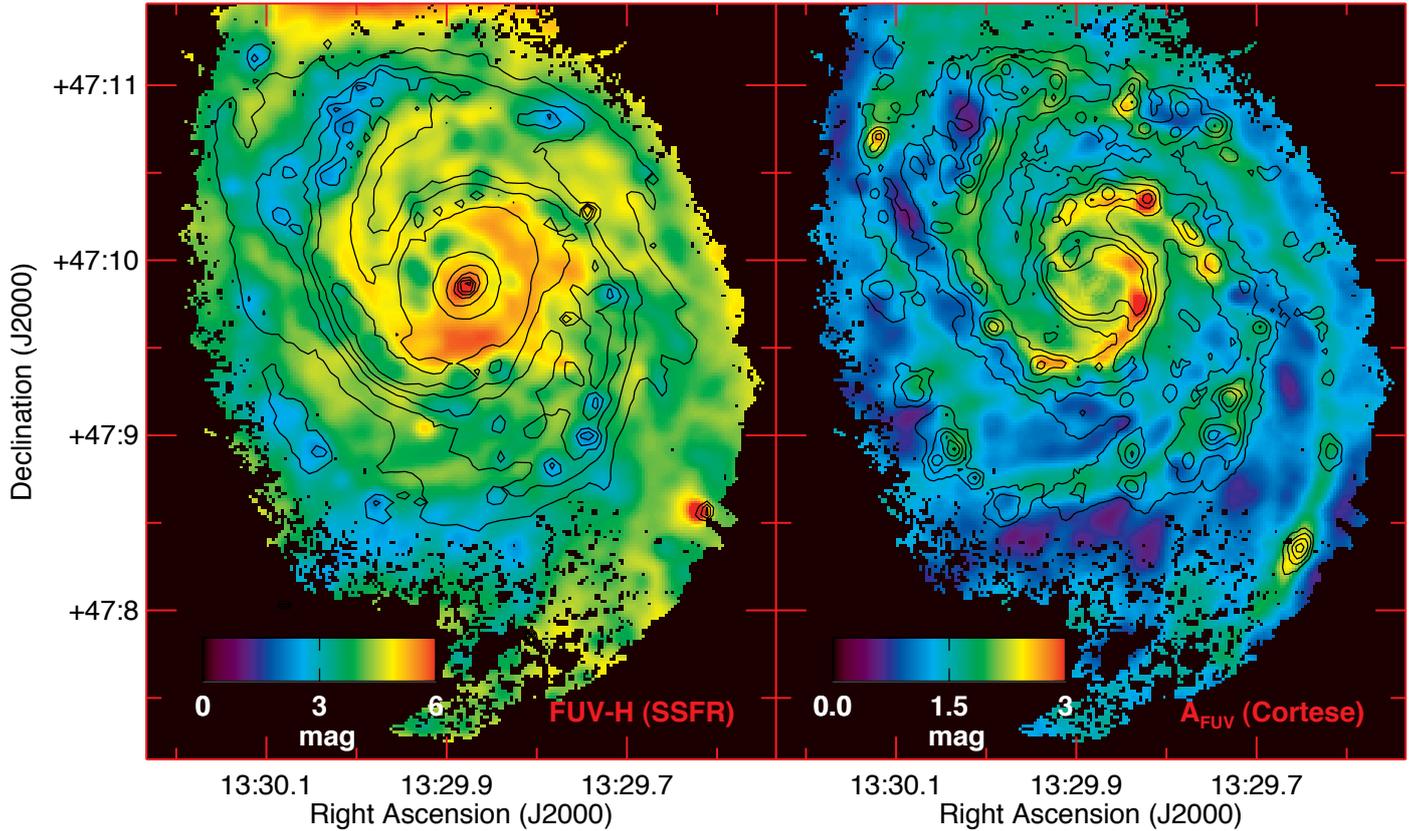
#### Old stellar population

We model the stellar emission spectrum of the old stars based on Maraston (1998, 2005) single stellar populations (SSP). The age and metallicity of the SSP are fixed to 10 Gyr old and  $Z = 0.02$ , consistent with the age and metallicity estimates in M51 reported by Bresolin et al. (2004); Moustakas et al. (2010) based on observational constraints, and the fitting of stellar SEDs by Mentuch Cooper et al. (2012).

#### Young non-ionizing stars

The stellar emission spectrum of FUV-emitting stars is described through Starburst 99 (Leitherer et al. 1999) templates for an

<sup>9</sup> The  $A_{\text{FUV}}$  calibrations presented in Cortese et al. (2008) were derived for a Salpeter (1955) IMF, but have been shown to be appropriate for other initial mass functions (e.g., Chabrier 2003) only resulting in small variations of  $A_{\text{FUV}}$  (0.1–0.2 dex) within the error bars.



**Fig. 2.** The FUV –  $H$  color map (left), which is considered a proxy of the specific star formation rate, and the FUV dust attenuation (right) in M51. The  $R$  band contours have been overlaid as black contours in the left panel, while the black contours in the right panel trace the dust distribution as observed from the PACS  $70\ \mu\text{m}$  image.

instantaneous starburst with solar metallicity ( $Z = 0.02$ ) and Kroupa (2002) IMF, normalized to a total mass of  $10^6 M_{\odot}$ , and averaged over stellar ages between 10 and 100 Myr.

#### Young ionizing stars

We describe the emission of ionizing stars based on the starburst templates from the library of pan-spectral SED models for young star clusters with ages  $< 10$  Myr presented by Groves et al. (2008). The SED templates of Groves et al. (2008) are based on a one-dimensional dynamical evolution model of HII regions around massive clusters of young stars, combined with the stellar spectral synthesis code Starburst 99 (Leitherer et al. 1999), and the nebular modeling code MAPPINGS III (Groves et al. 2004), assuming a Kroupa (2002) broken power-law IMF between 0.1 and  $120 M_{\odot}$ . The important parameters controlling the shape of the emission spectrum are the metallicity ( $Z$ ) of the gas, mean cluster mass ( $M_{\text{cl}}$ ), age and compactness ( $C$ ) of the stellar clusters, the pressure of the surrounding ISM ( $P_0$ ), and the cloud covering fraction ( $f_{\text{PDR}}$ ). We choose a solar metallicity ( $Z_{\odot}$ ) consistent with the metallicity of the old stellar component in the RT model of M51. The age parameter can be eliminated by averaging the spectra for all 21 cluster ages from 0.01 to 10 Myr. We assume fixed values for the mean cluster mass ( $M_{\text{cl}} \sim 10^5 M_{\odot}$ ) and ISM pressure ( $P_0/k \sim 10^6 \text{ cm}^{-3} \text{ K}$ ). The latter approximations are based on the fact that more massive star clusters can be modeled as the superposition of several individual clusters, and that the variation in ISM pressure mainly affects the nebular emission lines rather than the shape of the emission spectrum (Groves et al. 2008). We assume a cloud covering fraction  $f_{\text{PDR}} = 0.2$ , which is consistent with the assumptions of Jonsson et al. (2010) applied for the modeling of star-forming

regions in nearby galaxy samples, and in line with observations in M51 and other nearby galaxies finding negligible fractions of completely obscured star-forming regions (Calzetti et al. 2005; Prescott et al. 2007). Given our choices of the mean cluster mass  $M_{\text{cl}}$  and ISM pressure  $P_0/k$ , we can choose values for the compactness parameter ranging from  $\log C = 4.5$  to 6.5. The compactness of the stellar clusters will mainly affect the temperature of dust grains with star clusters of lower densities having their peak at longer far-infrared wavelengths, and being, thus, characterized by colder dust temperatures. We assume a moderate compactness factor of  $\log C = 5.5$  (see also De Looze et al. 2012b), corresponding more or less to a warm dust component with dust temperature  $T_{\text{d}} = 30$  K.

#### Dust component

We assume an uniform composition of dust particles with a fixed grain size distribution throughout the galaxy, with the abundances of the grains (graphites, silicates and PAHs), extinction, and emissivity of the dust mixture taken from the Draine & Li (2007) model derived for the dust in the Milky Way.

#### 3.2.3. SKIRT model normalization

Sections 3.2.1 and 3.2.2 have described our assumptions on the geometrical distribution of stars and dust, and the emission spectra for different stellar populations and dust properties, respectively. The normalization factors of stellar and dust components are the only free parameters in the model of M51, which are not fixed a priori but determined through minimization procedures. The normalization of stellar components includes the scaling of the stellar emission spectra calculated for a single star to

a certain bolometric or monochromatic luminosity representative for the stellar population. The luminosity of a single stellar photon will be determined as the total luminosity of the stellar component divided by the number of photon packages. The RT calculations presented in this paper are based on  $5 \times 10^6$  photon packages. The normalization of the dust component sets the amount of dust particles in the RT calculations through the total dust mass. Given that the density variations in the stellar and dust geometries are based on the relative intensity variations as perceived from observations, we rely on global photometry measurements to determine the normalization factors.

Since the observed images used as input geometries for the RT calculations were cut at a certain signal-to-noise level, we might miss some of the extended stellar and dust emission observed in M 51. Global photometry measurements for M 51 published in the literature will, therefore, not be applicable for the RT model. We determine flux measurements adjusted to the coverage in the RT model of M 51 by summing the emission within the areas of observed images that were retained as input geometries to the RT simulations. The flux measurements derived in this manner are on average about  $\sim 10\text{--}20\%$  lower compared to published values in the literature that account for more extended emission features. The previously published photometry of M 51 is likely to have been derived through different aperture photometry techniques from images to which different data reduction techniques and calibrations were applied. Given that our fluxes in all bands are determined in a homogeneous way, small differences between our photometry and the published photometry are expected.

The uncertainties on the flux measurements include calibration uncertainties and stochastic uncertainties which reflect the uncertainties due to the method of flux extraction (see Ciesla et al. 2012 for a more detailed description on error determination). The stochastic uncertainties include the instrumental uncertainties (related to the data coverage for each pixel), the background uncertainties, and confusion noise. The background uncertainties are determined as the standard deviation of the mean values of several background regions (see Sect. 2.2). The confusion errors only become significant in SPIRE wavebands, for which we rely on the confusion noise measurements from Nguyen et al. (2010). The different error contributions are added in quadrature. The obtained fluxes are not corrected for the variation in beam size according to the shape of the spectrum. To allow a direct comparison between the observed band fluxes and the radiative transfer models, we convolve the output SED from the RT simulations with the appropriate response curves to obtain the model flux densities for a certain waveband.

The normalization constraints used for every RT component are briefly outlined. All model parameters are optimized in such a way that they reproduce the observed global flux densities within the error bars. Table 2 provides an overview of the normalization factors ( $V$  band luminosity,  $L_V$ , for the old stellar populations in the bulge and thick disk, star formation rates, SFR, for the (non-)ionizing young stars in the thin disk, and dust mass,  $M_d$ , for the dust component in the thin disk) obtained from the model fitting described below.

#### Old stellar population

We constrain the luminosity of the old stars in the bulge and thick disk through the observed stellar SED. The global intensity of the old stellar component is scaled linearly until our models achieve the best fit with the stellar SED using a  $\chi^2$  minimization procedure. To constrain the observed stellar SED, we only account for the 2MASS  $JHK$  and IRAC  $3.6\mu\text{m}$  wavebands since their emission is dominated by old stars, and the effect of dust

attenuation can safely be considered negligible. We scale the relative contributions to the global luminosity of the bulge and thick disk stellar components assuming a bulge-to-disk ratio of 0.16 (see fitting results from Peng et al. 2010).

#### Young stars

The luminosity of the non-ionizing and ionizing stars is scaled to reproduce the observed global FUV and MIPS  $24\mu\text{m}$  emission, respectively. We, hereby, postulate that the scaling of the input emission spectra of non-ionizing and ionizing stars correspond to the same star formation rate.

We obtain a SFR of  $\sim 3 M_\odot \text{ yr}^{-1}$  for young stars with ages younger than 100 Myr. The level of star formation obtained from the RT simulations corresponds well to the SFR estimates  $SFR(24\mu\text{m}) \sim 3.4 M_\odot \text{ yr}^{-1}$  and  $SFR(\text{FUV}) \sim 4.3 M_\odot \text{ yr}^{-1}$  derived by Calzetti et al. (2005) for a Kroupa (2001) IMF (consistent with a Kroupa (2002) IMF assumed in this work).

#### Dust component

The dust mass is scaled to reproduce the FUV attenuation. Since the  $A_{\text{FUV}}$  map derived from the SSP fitting procedure varies throughout M 51, and, thus, sets multiple constraints to the dust component in the radiative transfer model, we perform a least squares fitting procedure that minimizes the function:

$$f(\alpha) = \sum_{i=0, N_{\text{pix}}} \left( \frac{A_{\text{FUV,obs}}[i] - A_{\text{FUV,RT}}[i]}{\sigma_{A_{\text{FUV,obs}}}[i]} \right)^2 \quad (2)$$

which sums over all the pixels that resulted in reliable  $A_{\text{FUV}}$  estimates. Each term in the sum is weighted by the uncertainty on the FUV attenuation. The model FUV attenuation is derived based on the output FUV images for the RT simulations with and without dust:  $A_{\text{FUV}} = -2.5 \log(I_{\text{total}}/I_{\text{transparent}})$ . To this, we add the FUV attenuation derived for the young star clusters with ages  $< 10$  Myr based on the model properties for the emission spectra presented in Groves et al. (2008). Given that the emission of photo-dissociation regions heated by young stars is modeled as a dusty foreground screen by Groves et al. (2008), we convert the dust mass of young stellar clusters to a FUV attenuation relying on the  $V$  band attenuation-to-hydrogen column density ratio,  $A_V/N_H = 5.34 \times 10^{-22} \text{ mag cm}^2 \text{ H}^{-1}$ , and the dust-to-hydrogen ratio,  $\Sigma M_d/(N_H m_H) = 0.01$ , derived for the Draine & Li (2007) dust model. The dust attenuation curve applied by Groves et al. (2008) corresponds to the  $R_V = A_V/E_{B-V} = 3.1$  curve presented by Fischera & Dopita (2005) with  $A_{\text{FUV}}/A_V \sim 3$ , which is used to convert the obtained  $A_V$  values into FUV band attenuation.

The optimization of the normalization of young stellar emission and dust mass is performed simultaneously because the youngest stars are most obscured being still (partly) embedded in their birth clouds. The youngest stars are also hotter and emit the bulk of their energy at shorter wavelengths, which are more easily absorbed, and scattered by dust grains. Based on the global FUV and  $24\mu\text{m}$  flux density of M 51, and the FUV band attenuation obtained from the procedure outlined in Cortese et al. (2008; see Sect. 3.2.1), we are able to simultaneously constrain the luminosity of young stars and the dust mass. A possible degeneracy between both parameters is non-existent because an increase in dust mass will lower the FUV emission, and, therefore, require an additional scaling of the luminosity of young stars. The dust mass can, however, only be scaled to a maximum value that still reproduces the total infrared emission observed in M 51. Within the uncertainties of the assumed dust properties, the resulting dust masses and stellar luminosities might differ depending on the variation in dust opacities and grain composition in the dust model. We assume an uniform dust population with grain composition, extinction, and emissivity properties for the

dust mixture taken from the [Draine & Li \(2007\)](#) model for the dust in the Milky Way, which has been shown to reproduce the observed IR emission of several nearby galaxy samples (e.g., [Dale et al. 2012](#); [Ciesla et al. 2014](#)).

Other than the dust mass used in the radiative transfer simulations, we have to account for the dust mass residing in the photo-dissociation regions surrounding HII regions. With the emission of individual stars clusters and their embedding dusty envelopes occurring on scales smaller than the resolution of our dust grid (i.e., 40 pc in the highest resolution dust grid cells), we have approximated their emission based on the library of SEDs presented by [Groves et al. \(2008\)](#) to model an ensemble of HII regions surrounded by photo-dissociation regions with a certain covering factor (here  $f_{\text{PDR}} \sim 0.2$ ). The total dust mass associated with the young star clusters of ages  $<10$  Myr, however, only accounts for 6% of the total mass budget in M51. The two-phase dust component in the disk accounts for  $7.3 \times 10^7 M_{\odot}$ , while  $0.4 \times 10^7 M_{\odot}$  is distributed in the outer shells of young star clusters. We find a total dust mass of  $7.7 \times 10^7 M_{\odot}$  which is about  $\sim 35\%$  lower compared to the dust mass  $M_d \sim 1.19 \times 10^8 M_{\odot}$  obtained from the dust SED modeling with the [Draine & Li \(2007\)](#) models reported by [Mentuch Cooper et al. \(2012\)](#).

The 35% difference in dust mass for M51 derived from the RT model (this paper) and a single modified-blackbody (MBB) fitting procedure ([Mentuch Cooper et al. 2012](#)) can likely be attributed to the different model assumptions. The RT model, first of all, accounts for a wide range of dust temperatures along the line-of-sight (with single dust temperatures derived for every dust grain species within every 3D dust grid cell) as opposed to the single dust temperature approximation assumed for MBB fitting. The RT model is, furthermore, capable of self-consistently explaining the UV and mid-infrared emission from young stars, the dust obscuration and emission, which provides strong constraints on the sources of dust heating (see Sect. 5) and relative dust-star geometry (see Sect. A.1). Based on the dust mass derivation from the RT model ( $M_d = 7.7 \times 10^7 M_{\odot}$ ), and the gas mass reported in [Mentuch Cooper et al. \(2012\)](#), we obtain a gas-to-dust mass ratio of 142, which is consistent with the gas-to-dust mass fractions obtained for our Milky Way (160, [Zubko et al. 2004](#)), and other nearby late-type spiral galaxies from the *Herschel* Reference Survey ([Cortese et al. 2012](#)).

## 4. RT model validation

### 4.1. Global galaxy scale

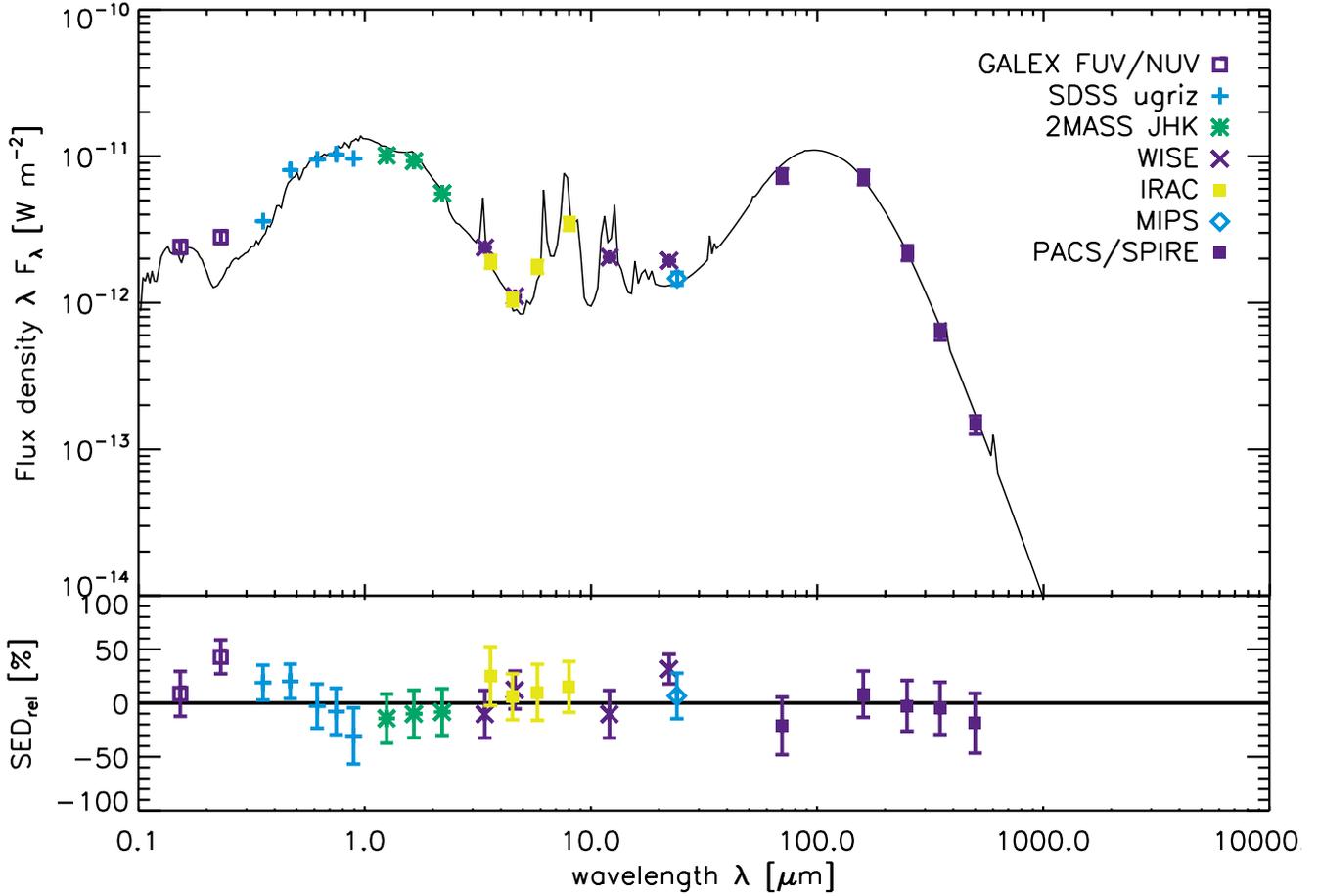
Figure 3 shows the multi-wavelength SED in our radiative transfer model (black solid line) overlaid with the observed multi-wavelength constraints, derived through summing the emission within the high signal-to-noise regions in M51 (see Sect. 3.2.3). Figure 3 indicates that the RT model generally reproduces the observed FIR/submm emission in M51 within the error bars. The amount of absorbed stellar energy (i.e., 10% of the total stellar bolometric luminosity) is, thus, shown to be in perfect balance with the thermal energy re-emitted by dust grains. The dust energy balance in M51 is, hereby, in contrast with conclusions of many similar studies on edge-on spiral galaxies ([Popescu et al. 2000, 2011](#); [Misiriotis et al. 2001](#); [Alton et al. 2004](#); [Dasyra et al. 2005](#); [Baes et al. 2010](#); [De Looze et al. 2012a,b](#); [Holwerda et al. 2012](#)), reporting the underestimation of the observed FIR/submm emission by a factor of 3–4 in RT models, which were constrained based on optical observations of the dust attenuation. Different scenarios were invoked to reconcile the model calculations with the observations, among

which higher FIR/submm dust emissivities ([Alton et al. 2004](#); [Dasyra et al. 2005](#); [MacLachlan et al. 2011](#)), or the distribution of a major dust fraction in geometries (second thin inner dust disk, two-phase clumpy dust) that have a negligible effect on the dust attenuation ([Popescu et al. 2000, 2011](#); [Misiriotis et al. 2001](#); [Bianchi 2008](#); [MacLachlan et al. 2011](#)), are the most favorable explanations.

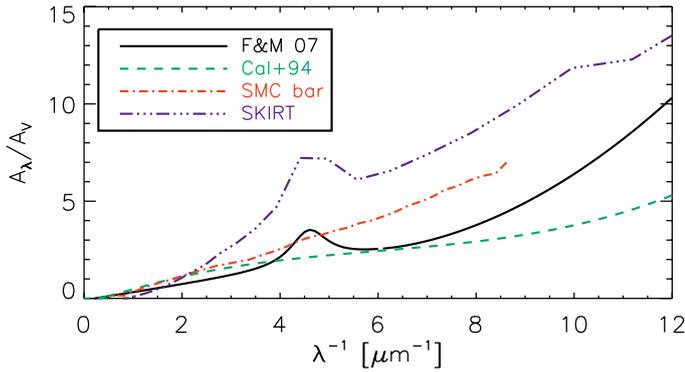
We believe that the reason for the dust energy balance in M51 compared to the inconsistency in the dust energy budget for edge-on galaxies can be looked for in the different aspects of our modeling approach. The main difference in construction of our RT model depends on the viewing angle of M51, which allows to discern the position of star-forming regions, and asymmetries in stellar and dust geometrical distributions. The adequate description of these sites of young star clusters, and the observed asymmetric structures in the RT model allows to make a careful prediction of the radiative heating of dust by young and old stellar populations, resulting in a more reliable estimate of the thermal dust emission at IR/submm wavelengths. Another difference lies in the dust geometry, and the constraints used to scale the total dust mass in the RT model. Rather than constraining the dust attenuation from optical observations, we have calculated the FUV attenuation based on the [Cortese et al. \(2008\)](#) recipes which simultaneously rely on FUV and infrared observations to derive  $A_{\text{FUV}}$  estimates. The compatibility between the absorbed energy of stars and the re-emitted dust emission in M51, diverging from what is observed for edge-on galaxies, suggests that the origin of the dust energy balance problem in edge-ons results from the incapability of constraining the location and emission of star-forming regions, and determining the amount of dust attenuation solely from UV/optical observations (see also [Kreckel et al. 2013](#)). The resolution of optical data in the determination of dust attenuation has also been shown to be of great importance (e.g., [Liu et al. 2013](#)) due to the small spatial scales (few parsecs) at which most of the attenuation processes take place.

To better highlight the differences between the observed and model emission in several wavebands, we show the residual flux densities, i.e., relative difference between observations and model flux densities or  $\text{SED}_{\text{rel}}[\%] = (F_{\text{obs}} - F_{\text{model}})/F_{\text{obs}}$ , in the bottom panel of Fig. 3. The uncertainties on the residuals include the uncertainties on the observed flux densities as well as the Monte Carlo noise from the radiative transfer calculations. The modeled SKIRT fluxes – distributed in a wavelength grid with 500 steps ranging from 0.05 to 5000  $\mu\text{m}$  – are convolved with the response curves of the respective bands to obtain the model flux.

In most wavebands, the observed and modeled fluxes agree within the error bars. One clear exception is the observed emission in the NUV band which is significantly lower in the RT model compared to the observations. We argue that the discrepancy can be attributed, at least in part, to the dust bump at 2175 Å characterizing the dust attenuation curve of our RT model for M51. The output dust attenuation curve from the RT model (see Fig. 4) shows a similar dust absorption feature as observed in the Galactic dust attenuation curve with selective-to-total extinction  $R_V = 3.0$  ([Fitzpatrick & Massa 2007](#)). Based on the analysis of [Calzetti et al. \(2005\)](#), the dust attenuation in M51 seems characterized by a weak 2175 Å absorption feature, as compared to the clear 2175 Å dust bump in the Milky Way (MW) extinction curve, and to better resemble the Small Magellanic Clouds (SMC) extinction curve ([Gordon et al. 2003](#)). With the 2175 Å dust feature being absent in the lower metallicity environments of the SMC and LMC ([Gordon et al. 2003](#)), the metallicity



**Fig. 3.** *Top:* multi-wavelength SED of M51 where the black solid line represents the emission derived from the radiative transfer model for M51 (including dust, old and young stars) and the legend indicates the origin of the different data measurements. *Bottom:* residual SED obtained as the relative difference between observations and model flux densities or  $SED_{rel}[\%] = (F_{obs} - F_{model})/F_{obs}$ .



**Fig. 4.** Dust attenuation curve as derived from our radiative transfer calculations (triple dotted-dashed purple line). The dust extinction curves derived for the Milky Way with  $R_V = 3.0$  (Fitzpatrick & Massa 2007) and the Small Magellanic Clouds (Gordon et al. 2003), and the Calzetti et al. (1994) attenuation law for starburst galaxies are added for comparison as black solid, red dotted-dashed and green dashed lines, respectively.

of galaxies has been argued to play a role in the strength of the bump. With the metallicity of M51 being solar to even super-solar locally, we would expect a prominent 2175 Å absorption feature if the metallicity plays a role in shaping the

dust composition, and, thus, the shape of the attenuation curve. Other than metallicity, the absorption bump at 2175 Å appears to be sensitive to the UV radiation field (e.g., Wild et al. 2011), which can alter the dust grain size distribution through destruction/coagulation processes, and/or change the ionization state of dust grains (Boulanger et al. 1988; Cesarsky et al. 1996). Also radiative transfer effects coupled with age-selective attenuation have been shown to affect the bump strength (Silva et al. 1998; Granato et al. 2000). The strength of the attenuation bump, furthermore, depends on the carriers of the extinction feature (Gordon et al. 1997; Witt & Gordon 2000) which could be related to dust grains composed of aromatic carbon (Draine 1989) such as graphite grains or PAHs. We, therefore, argue that the dust population in M51 might differ from the standard Draine & Li (2007) dust composition for the Milky Way. In Sect. A.2, we also show that the strength of the 2175 Å absorption feature might be influenced by the clumping fraction of dust.

Other than the dust 2175 Å absorption feature, the UV slope in the output attenuation curve from the RT model appears significantly steeper compared to the extinction curves derived for the MW (Fitzpatrick & Massa 2007), SMC (Gordon et al. 2003) and a sample of starburst galaxies (Calzetti et al. 1994). We do need to caution that the uncertainty on the  $V$  band attenuation of the RT model is large due to the relatively fewer  $V$  band photons that are scattered or absorbed compared to the shorter

wavelength UV photons (e.g., see the grainy appearance of the inter-arm regions in the  $\tau_V$  map in Fig. 7 due to Poisson noise). The output attenuation curve from SKIRT, furthermore, depends on the intrinsic extinction curve for the dust grains (which is similar to the MW extinction curve), but also reflects any 3D geometrical and/or scattering effects. If the curve is representative for the attenuation law of the RT model of M 51, the strong UV slope might be telling us that the total-to-selective extinction ratio,  $R_V$ , is smaller than the average  $R_V = 3.1$  observed in the Milky Way (Cardelli et al. 1989). The curve with  $R_V = 3.1$  for the Milky Way is, however, an extinction law, which does not include the effects of scattering and geometry. The attenuation curve from SKIRT does account for these effects, which could explain part of the difference. Scattering and the relative dust-star geometry have indeed been shown to play an important role in characterizing the attenuation law in face-on galaxies (Baes & Dejonghe 2001). A further investigation of the SKIRT attenuation curve with different dust models and relative dust-star geometries is beyond the scope of the present study.

#### 4.2. Spatially resolved galaxy scales

While the normalization of the SKIRT RT stellar and dust components (see Sect. 3.2.3) is optimized for global galaxy measurements (except for the FUV attenuation), deviations from the observed multi-waveband images of M 51 might occur in our radiative transfer model on spatially resolved scales. Even though the observed 2D structures in M 51 are used to constrain the dust and stellar geometries in the model, we might have disregarded spatial variations in the age of stellar populations, and/or deviations from a uniform dust composition throughout the galaxy. We need to investigate whether local dissimilarities occur between the radiative transfer model and observations, and, more importantly, understand the shortcomings of our model that result in its inability to reproduce the observed characteristics. Therefore, we have convolved all observed and model images to the working resolution of  $\sim 12.1''$  (or 493 pc) in the PACS  $160\ \mu\text{m}$  waveband.

##### 4.2.1. Stellar and dust emission

Figure 5 compares a handful of observed images of M 51 (1st column) with the model results from the RT calculations (2nd column). The third column shows the relative residuals, i.e.,  $(I_{\text{obs}} - I_{\text{model}})/I_{\text{obs}}$ , with the histogram of residuals (normalized to 1 at the peak) shown in the last column. From top to bottom, we compare the images observed in GALEX FUV, IRAC  $3.6\ \mu\text{m}$ , IRAC  $8\ \mu\text{m}$ , MIPS  $24\ \mu\text{m}$  and PACS  $160\ \mu\text{m}$  wavebands.

The FUV emission in the model matches reasonably well to the observed FUV emission in M 51, with the exception of some regions in the spiral arms, where the model overestimates the observed FUV emission up to 60%, and the underestimation of the FUV emission in the inter-arm regions. These small variations between the model and observations are most likely a reflection of the stellar clusters of different ages that populate the regions of M 51 throughout the center and spiral arms. While the radiative transfer model assumes an averaging of ages for stars younger than 100 Myr throughout the galaxy plane, the stellar clusters of specific ages are more likely clumped together in different regions, and not homogeneously spread throughout the galaxy (see for instance Bianchi et al. 2005; Kaleida & Scowen 2010). Darker regions in the residual map indicate areas where the emission of young ionizing stars ( $<10$  Myr)

is overestimated, and might, thus, merely consist of non-ionizing stars. Alternatively, the escape fraction of hard UV photons from their birth clouds might be overestimated in those regions. The bright regions in the residual FUV map indicate that the RT model lacks a more diffuse emission component in the inter-arm regions, which could hint at the presence of a thick disk component with dust that scatters UV photons into the line of sight. Alternatively, we might underestimate the escape fraction of young stellar photons which could produce a diffuse emission component.

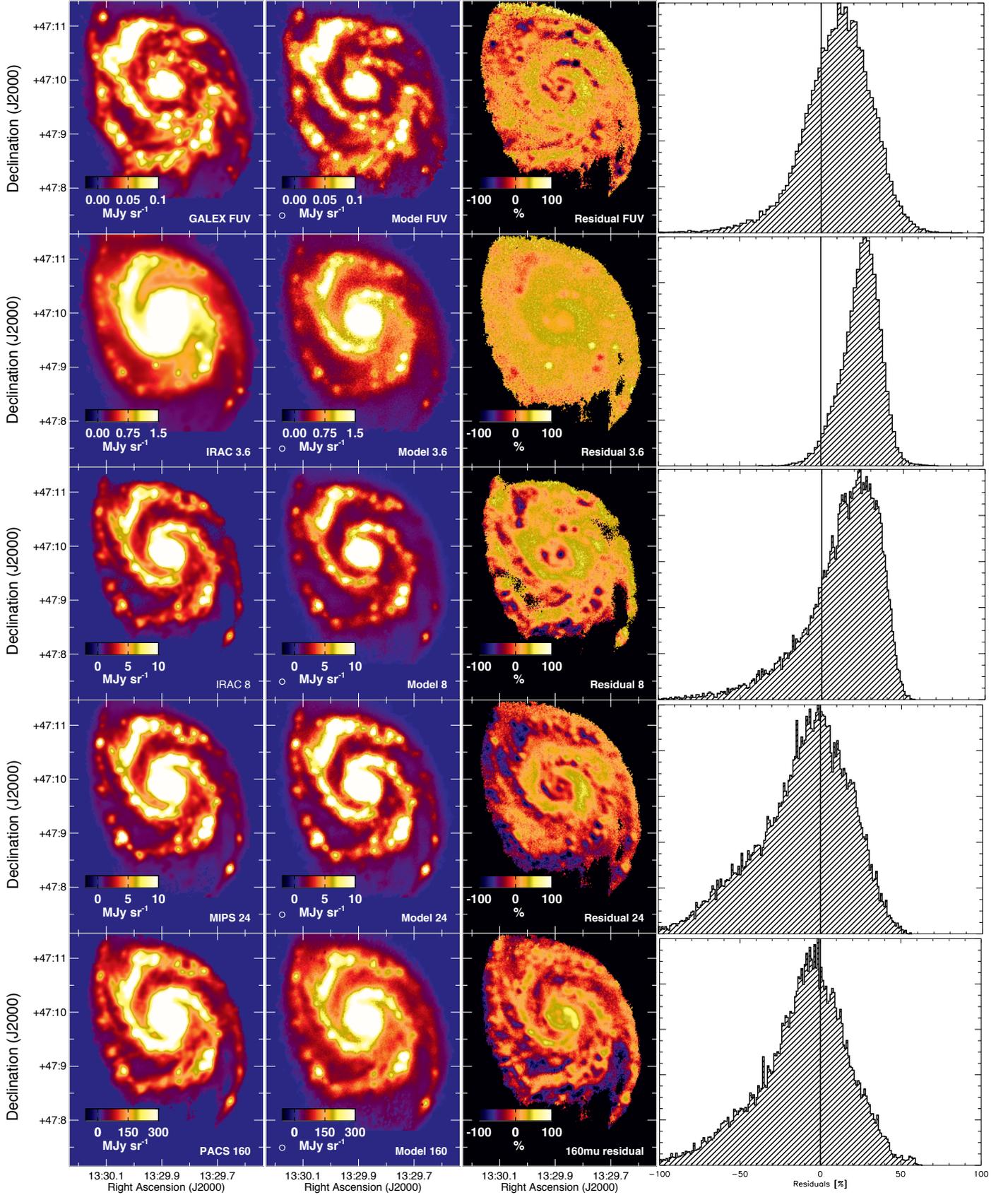
The good correspondence of the modeled  $3.6\ \mu\text{m}$  map with the observed IRAC  $3.6\ \mu\text{m}$  image (i.e., model and observations agree within 50%) suggests that the old stellar component in the RT model is representative for the population of old stars in M 51. Small deviations can be perceived in the inter-arm regions of M 51, where the model underestimates the diffuse stellar emission. The  $3.6\ \mu\text{m}$  emission seems, furthermore, overestimated in some regions of the spiral arms, which suggests that the contribution of young stars to the near-infrared wavebands is somewhat too high in the model.

The IRAC  $8\ \mu\text{m}$  emission is overestimated in the spiral arms by at most 50%, whereas the model also underpredicts the PAH emission observed in some inter-arm regions of M 51. The absence of bright  $8\ \mu\text{m}$  emission in the spiral arms might, partly, be linked to the underestimation of the  $24\ \mu\text{m}$  emission and, thus, star formation activity (see later). Alternatively, the improper characterization of the PAH abundance in the RT model (assuming a constant PAH fraction throughout the galaxy), which does not account for the destruction of PAHs exposed to hard radiation fields (e.g., Boselli et al. 2004; Helou et al. 2004; Calzetti et al. 2005; Draine et al. 2007; Prescott et al. 2007; Bendo et al. 2008; Gordon et al. 2008), might explain the differences between model and observations.

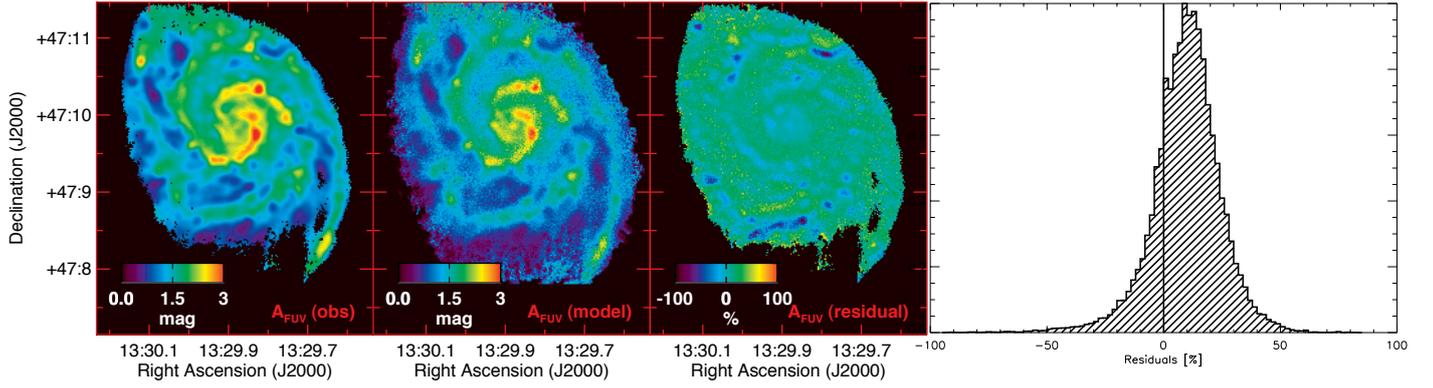
We observe a similar pattern in the  $24\ \mu\text{m}$  emission of M 51, with the model overestimating the emission up to 50% in the spiral arm structures, i.e., the sites of young stellar activity, and underestimating the diffuse  $24\ \mu\text{m}$  in the inter-arm regions and disk of M 51. We argue that the former discrepancy could again originate from the uniform age range assumed for young stars throughout M 51 (averaging all ages younger than 100 Myr), while radial differences have been shown to occur with a decrease in age of the young stellar populations towards larger radial distances from M 51's center (Calzetti et al. 2005). Alternatively, the deviation from an uniform dust grain composition in some parts of the galaxy could contribute to the spatial inaccuracies in our modeled  $24\ \mu\text{m}$  image. Since the abundance of very small grains (VSGs) has been shown to trace the star formation activity (e.g., Calzetti et al. 2005; Bernard et al. 2008; Paradis et al. 2009), it would not be surprising if the abundances of very small grains would be higher in localized star-forming regions due to the shattering of larger grains. The increase of the PAH  $8\ \mu\text{m}$ -to-MIPS  $24\ \mu\text{m}$  ratio with galactocentric radius (Vlahakis et al. 2013), indeed, suggests that the excitation and/or abundance of PAHs, and very small grains is not uniform throughout the galaxy's disk.

At longer infrared wavebands, we observe the opposite trend with the model underestimating the emission from the spiral arms, and overestimating the diffuse dust emission in the disk of M 51 up to 50%.

In summary, we find local variations in the RT model emission deviating from the observed stellar and dust emission in M 51. The local deviations suggest that the model assumptions of an uniform dust composition, and the averaging of stellar ages throughout the galaxy are only adequate to first order. The



**Fig. 5.** Observed (left column), modeled (2nd column), and residual (3rd column) images of M51 for the GALEX FUV (top row), IRAC 3.6  $\mu\text{m}$  (2nd row), IRAC 8  $\mu\text{m}$  (3rd row), MIPS 24  $\mu\text{m}$  (4th row), and PACS 160  $\mu\text{m}$  (5th row) wavebands. The observed images have been convolved to the resolution of the PACS 160  $\mu\text{m}$  filter, i.e., the resolution of all input parameters maps in the radiative transfer code. The size of the PACS beam ( $\sim 12$  arcsec) is indicated in the bottom left corner of the model images. The last column shows the histogram of residuals (normalized to 1 at the peak).



**Fig. 6.** From left to right: observed, modeled, and residual image of the FUV dust attenuation,  $A_{FUV}$ . The last panel shows the histogram of residuals (normalized to 1 at the peak).

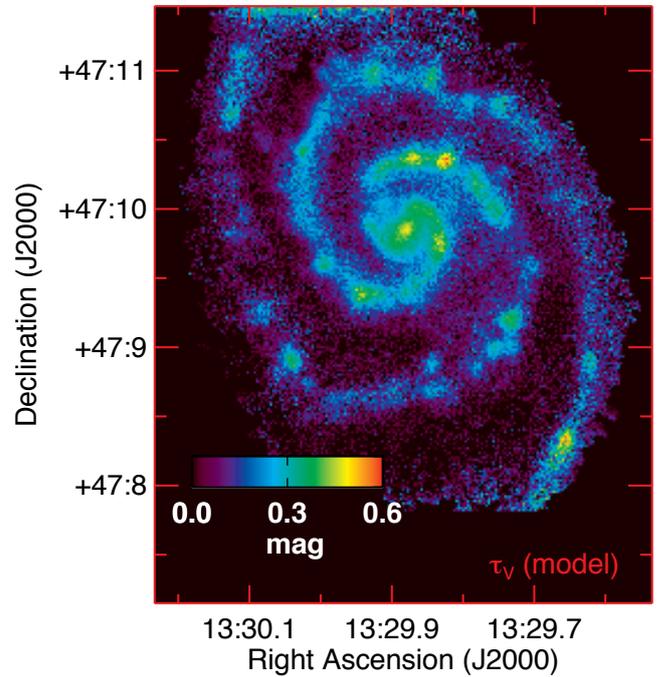
characterization of the dust grain size distribution, abundances, and grain emissivities, in combination with the stellar age variation, requires several additional datasets that are sensitive to the emission and absorption features of individual grain populations (e.g., *Spitzer* IRS, *SWIFT*, ...), and is, therefore, beyond the scope of this current work. The underestimation of the diffuse emission in M 51, furthermore, suggests that we might need to consider the presence of an additional dust component distributed in a disk with larger scale height.

#### 4.2.2. Dust attenuation

Figure 6 presents the observed, model, and residual images of the FUV dust attenuation in M51 (see first three panels) with the histogram of residuals (normalized to 1 at the peak) shown in the last column. The observed FUV attenuation map derived from the procedures outlined by Cortese et al. (2008) is well reproduced by our model, which suggests that the relative orientation of stars and dust in our model, and, thus, the scale heights of dust and stars are modeled in a realistic way. Small deviations occur in some localized sources in the spiral arms, which suggests that our model overestimates the attenuation in parts of the galaxy, mainly located in the outer regions of the spiral arms. With those outer spiral arms being mostly affected by the interaction with NGC 5195, the material in the outer arm regions might have been re-arranged due to the intervention of NGC 5195, making UV photons capable of escaping more easily from their birth clouds. A metallicity gradient, which has not been accounted for in the RT model assuming a constant metallicity, might also cause the FUV attenuation to be lower at longer galactocentric distances. The variations between model and observations, however, generally remain within the uncertainties of the FUV attenuation.

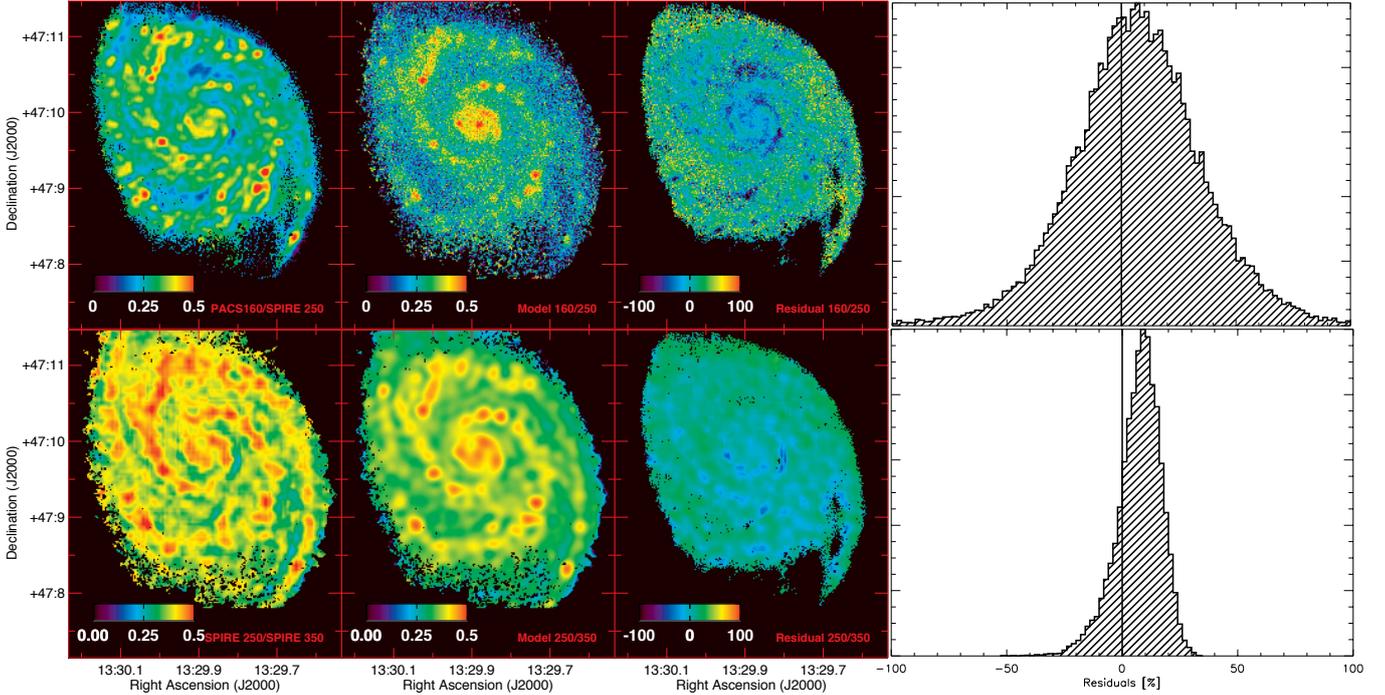
#### 4.2.3. Optical depth

To analyze the transparency of the stellar disk at the inclination of M51, we compare the RT model with and without dust by constructing a map of the optical depth,  $\tau_V = \exp(-I_{V,\text{total}}/I_{V,\text{transparent}})$ , from the model (see Fig. 7), with  $I_{V,\text{total}}$  representing the total  $V$  band intensity and  $I_{V,\text{transparent}}$  the intrinsic  $V$  band intensity in the absence of any dust (i.e., a transparent medium). For values  $\tau_V < 1$ , the galaxy is considered nearly transparent, implying that most of the stellar radiation can be observed directly without the obscuring effect of dust. Higher



**Fig. 7.** Optical depth in the  $V$  band,  $\tau_V$ , derived from the RT calculations for M51. The  $\tau_V$  map has a grainy structure in the inter-arm regions due to Poisson noise, driven by the relatively few number of  $V$  band photons that are absorbed and/or scattered in the inter-arm regions.

values of  $\tau_V$  would imply that a significant portion of the stellar  $V$  band light is absorbed, and reprocessed by dust grains. Based on radiative transfer simulations of edge-on spiral galaxies, the face-on optical depth in most optical wavebands is generally concluded to be lower than one, making the disk nearly transparent in the optical wavelength regime (Xilouris et al. 1997, 1998, 1999; Alton et al. 2004; Bianchi 2007; Popescu et al. 2011; De Geyter et al. 2013, 2014). Some exceptions were found for the secondary dust disk used to explain the missing FIR emission in RT models of edge-on galaxies (Popescu et al. 2000, 2011; Misiriotis et al. 2001; MacLachlan et al. 2011). The model fitting procedures of edge-on galaxies have, however, been shown to be influenced by model degeneracies between dust scale length and face-on optical depth (De Geyter et al. 2014), which could hamper the derivation of a reliable  $\tau_V$  estimate.



**Fig. 8.** Surface brightness intensity ratio maps PACS 160  $\mu\text{m}$ -to-SPIRE 250  $\mu\text{m}$  (*top*) and SPIRE 250  $\mu\text{m}$ -to-SPIRE 350  $\mu\text{m}$  (*bottom*) as derived from observations (*left panel*) and radiative transfer calculations (*second panel*). The panels in the third and last row show the residual images and the histogram of residuals (normalized to 1 at the peak), respectively.

The RT model for M51 shows a variation in optical depth from  $\tau_V < 0.1$  in the inter-arm regions,  $\tau_V \sim 0.3$  on average in the spiral arms, with peaks up to  $\tau_V \sim 0.6$  in localized star-forming regions in the spiral arms and the center of M51. The results for M51 support the earlier analyses of edge-on galaxies, indicating that the optical stellar light of galaxies is hardly influenced by extinction when viewed face-on. The relatively small effects of dust in optical wavebands is re-assuring for the determination of intrinsic galaxy properties such as stellar mass, star formation history, star formation rate, and global galaxy population diagnostics such as initial mass function, luminosity functions, and color-magnitude diagrams (e.g., Driver et al. 2007; Maller et al. 2009; da Cunha et al. 2010). We should, however, extend the same analysis to larger samples of galaxies with wide spread in star formation rate to confirm that also the optical disks of more actively star forming disks hardly experience effects of dust extinction. We should, furthermore, caution that the  $\tau_V$  values are averaged on scales of 12.1'' (or 493 pc at the distance of M51), which does not exclude the possibility that higher optical depths occur on scales smaller than 500 pc (i.e., the size of typical giant molecular clouds). Averaged over larger areas, the peaks of high optical depth will be smoothed out (e.g., Liu et al. 2013).

#### 4.2.4. Infrared colors

We verify whether our RT model is sensitive to the dust temperature variation on spatially resolved scales, and, thus, capable to reproduce the observed infrared colors. We compare the flux density ratios, PACS 160  $\mu\text{m}$ -to-SPIRE 250  $\mu\text{m}$ , and SPIRE 250  $\mu\text{m}$ -to-SPIRE 350  $\mu\text{m}$ , as derived from observations to the output from radiative transfer calculations (see Fig. 8). If we assume that a single modified black-body provides an adequate description of the dust SED at wavelengths long wards

of about 100  $\mu\text{m}$  (e.g., Smith et al. 2012; Galametz et al. 2012; Auld et al. 2013; Rémy-Ruyer et al. 2013; Cortese et al. 2014), i.e.,

$$F_\nu = \frac{M_d}{D^2} \kappa_{\nu_0} \left( \frac{\nu}{\nu_0} \right)^\beta B_\nu(T), \quad (3)$$

then the ratio of the flux densities at different infrared wavebands will be independent of the dust absorption coefficient,  $\kappa_{\nu_0}$ , defined at a reference wavelength  $\nu_0$ . Indeed, the flux ratios or infrared colors, i.e.,

$$\frac{F_\nu(\nu_1)}{F_\nu(\nu_2)} = \left( \frac{\nu_1}{\nu_2} \right)^{\beta+3} \left( \frac{e^{h\nu_2/kT} - 1}{e^{h\nu_1/kT} - 1} \right), \quad (4)$$

will only depend on the dust emissivity index,  $\beta$ , and the dust temperature,  $T_d$ . Figure 8 shows the observed (1st column), modeled (2nd column), and residual (3rd column) flux density ratios of PACS 160  $\mu\text{m}$ -vs.-SPIRE 250  $\mu\text{m}$  (top) and SPIRE 250  $\mu\text{m}$ -vs.-SPIRE 350  $\mu\text{m}$  (bottom). The last column shows the histogram of residuals, normalized to 1 at the peak. The spatial resemblance of peaks and hollows between the observed and modeled surface brightness intensity ratios makes us confident that the RT model is capable of giving an unbiased view on the importance and variance in dust heating sources across the galaxy. The PACS 160  $\mu\text{m}$ -vs.-SPIRE 250  $\mu\text{m}$  color shows good global agreement with some local deviations from the observed ratios. The PACS 160  $\mu\text{m}$ -vs.-SPIRE 250  $\mu\text{m}$  colors are too warm in some localized regions in the inter-arms of M51, which could indicate that the escape fraction of young stars from star-forming regions in the spiral arms is too high, and capable of heating the diffuse dust in the inter-arm regions. The agreement between model and observations is better for the SPIRE 250  $\mu\text{m}$ -to-SPIRE 350  $\mu\text{m}$  color with a relatively small scatter in the distribution of residuals around the mean value, which suggests that

the heating of the colder dust component in M51 is accounted for in a most reliable way by the RT model.

## 5. Dust heating in M51

### 5.1. Wavelength-dependent dust heating

The 3D high-resolution radiative transfer model of M51 allows a self-consistent study of the dust heating mechanisms on spatial scales of 500 pc. By tracing the individual photons of the different stellar populations in our radiative transfer calculations, we can investigate how much of the dust emitting at certain wavelengths is heated by the young and old stellar population. In the case of M51, we ignore possible AGN activity as a significant source of mid-infrared emission, and, thus, dust heating based on the results from [Mentuch Cooper et al. \(2012\)](#). We make a distinction here between the old stellar population, and the radiation from young stars with ages  $\lesssim 100$  Myr. A clear distinction between the two sources of dust heating is hampered by the non-linearity of the relation between the shape of the absorbed emission spectrum, and the re-emitted emission at specific wavelengths. The sum of the energy absorbed from young and old stars is equal to the total absorbed energy (i.e.,  $L_{\text{abs,total}} = L_{\text{abs,old}} + L_{\text{abs,young}}$ ), which can be translated into the total energy re-emitted at infrared and (sub)millimeter wavebands (i.e.,  $L_{\text{TIR,total}} = L_{\text{TIR,old}} + L_{\text{TIR,young}}$ ). The absorbed energy of young and old stars, separately, can, however, not simply be translated in the dust emission at specific monochromatic wavebands. More specifically, the dust emission in a monochromatic waveband depends on the global embedding radiation field, and, therefore, can not be separated into a young and old emission spectrum, i.e.:  $F_{X,\text{total}} \neq F_{X,\text{old}} + F_{X,\text{young}}$ , for a monochromatic waveband X. Since the non-linearity in the dependence of dust heating on the radiation field can not be eliminated from the RT calculations, the heating due to old and young stars cannot strictly be separated. To account for this non-linear behavior, we approximate the heating by old and young stars through

$$F'_{X,\text{old}} = \frac{0.5F_{X,\text{old}} + 0.5(F_{X,\text{total}} - F_{X,\text{young}})}{F_{X,\text{total}}} \quad (5)$$

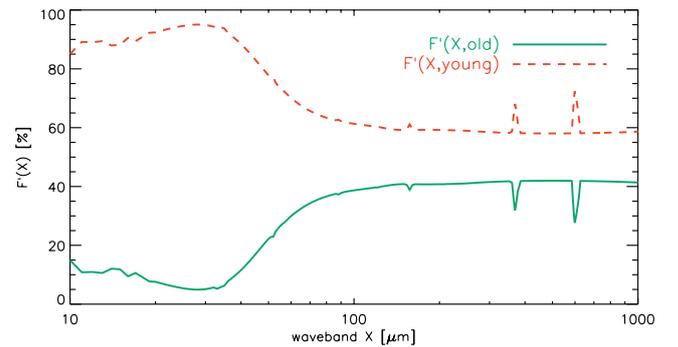
and,

$$F'_{X,\text{young}} = \frac{0.5F_{X,\text{young}} + 0.5(F_{X,\text{total}} - F_{X,\text{old}})}{F_{X,\text{total}}}, \quad (6)$$

respectively, for which  $F'_{X,\text{total}} = F'_{X,\text{old}} + F'_{X,\text{young}}$  does hold. The flux densities  $F_{X,\text{total}}$ ,  $F_{X,\text{old}}$ , and  $F_{X,\text{young}}$  are derived from repeated radiative transfer simulations with a stellar population consisting of old+young, old, and young stars, respectively. Figure 9 shows the relative fraction of energy from old and young stars contributing to the heating of dust emitting at infrared wavelengths from 10 to 1000  $\mu\text{m}$ . Table 3 indicates the fraction (in %) of dust heating by old stars (middle column), and young stars (right column) at several wavebands. Based on the radiative transfer calculations, we can conclude that the dust heating by the young stars in M51 clearly dominates in mid-infrared wavebands, and gradually decreases towards longer wavelengths. The evolved stellar population has a non-negligible contribution to the dust heating in far-infrared and submillimeter wavebands, which is consistent with previous multi-band studies of the dust heating instigators (e.g., [Bianchi 2008](#); [Bendo et al. 2010, 2012a, 2014](#); [Boquien et al. 2011](#); [Popescu et al. 2011](#); [Xilouris et al. 2012](#)). Nonetheless,

**Table 3.** Dust heating fractions in several monochromatic infrared wavebands X, with  $F_{X,\text{old}}$  and  $F_{X,\text{young}}$  indicating the fraction of energy for radiative dust heating provided by old and young stars ( $<100$  Myr), respectively.

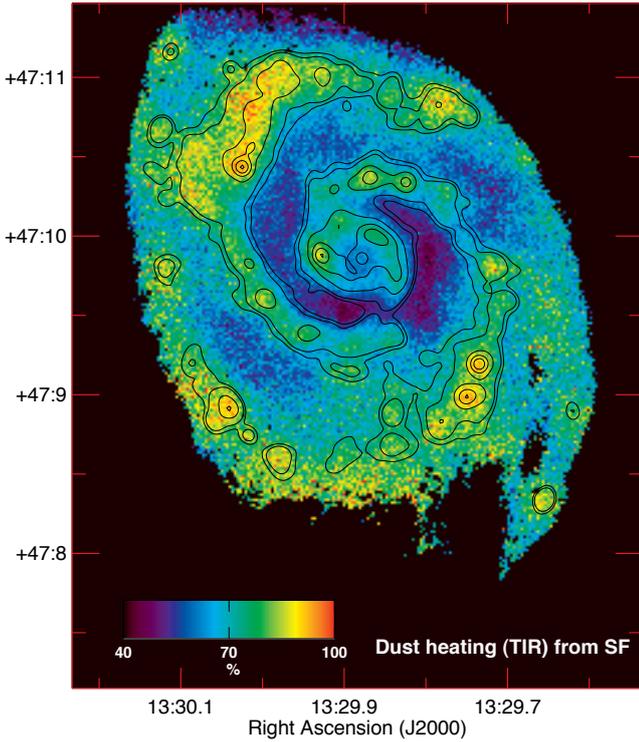
$\lambda$ [ $\mu\text{m}$ ]	$F_{X,\text{old}}$ [%]	$F_{X,\text{young}}$ [%]
24	6	94
70	34	66
100	39	61
160	40	60
250	41	59
350	42	58
500	42	58
1000	41	59
TIR	37	63



**Fig. 9.** Contribution from old (green solid line) and young stars (red dashed line) to the radiative heating of dust across the multi-wavelength infrared spectrum, as determined from the radiative transfer model of M51.

the stellar energy provided by young stars remains the dominant heating source also at longer wavelengths. The maximum contribution of  $\sim 40\%$  from the old stellar population to the dust heating in M51 is consistent with the upper limit found by [Popescu et al. \(2000\)](#); [Misiriotis et al. \(2001\)](#) for five edge-on spiral galaxies. [Law et al. \(2011\)](#) report an even lower contribution with less than 20% of energy to the global dust heating provided by old stars. The importance of young stars contributing to the heating of the diffuse, cold dust component is also evident from the tight correlation between the effective SFR ( $SFR/L_{500}$ ), and the cold dust temperature ( $T_c$ ) for 20 local star forming galaxies of the Key Insights on Nearby Galaxies: A Far-IR Survey with *Herschel* (KINGFISH) sample ([Kirkpatrick et al. 2014](#)). The scatter in the latter relation, and a trend between  $L_{3.6}/L_{500}$  and the dust emissivity index  $\beta$ , however, points at a non-negligible (although non-dominant) contribution from the old stellar population to the dust heating at FIR/submm wavelengths ([Kirkpatrick et al. 2014](#)).

For the entire dust energy budget emitted between 8 and 1000  $\mu\text{m}$ , the dust heating analysis is not affected by non-linear effects, and the total-infrared emission by dust heated through young and old stellar photons simply adds up to the total-infrared dust emission. We find that about 63% of the dust emission is provided through heating by young stars while the remaining 37% is heated through the more evolved stellar population. With M51 being located along the main sequence of star-forming galaxies in the local Universe ([Noeske et al. 2007](#)), we might need to consider a significant contribution of dust heating from evolved stars for nearby spiral galaxies with similar loci in the SFR-vs-stellar mass diagram (or thus similar sSFR).



**Fig. 10.** Contribution from young stars ( $\leq 100$  Myr) to the radiative dust heating responsible for the total-infrared emission of M 51. The black contours indicate the SFR obtained from a combination of FUV+ $24\ \mu\text{m}$  SFR calibrators (Hao et al. 2011) with contour levels of 0.075, 0.1, 0.25, 0.5, 0.75,  $1.0 \times 10^{-3} M_{\odot} \text{yr}^{-1}$ .

### 5.2. Spatially resolved heating of the total-infrared dust emission

Figure 10 shows the spatially resolved contribution to the dust heating from star formation to the total infrared emission ranging from about 40% to 100%. More specifically, we compute the fraction  $F'_{\text{TIR,young}}$  of the TIR emission that is due to heating by young stars. In the inter-arm regions of M 51, the dust heating appears mostly powered by evolved stars. With the black contours indicating the observed SFR in M 51, we can clearly see that the dust heating follows the morphology of star-forming regions located throughout the spiral arms. Especially, the northern spiral arm region located near the interacting galaxy NGC 5195 seems dominated by young stellar photons, that heat the dust in the spiral arms and well beyond. The peak of FUV emission in this part of M 51 indeed suggests that star formation has been triggered in the outer regions of the spiral arms in M 51, with many UV photons capable of escaping from their birth clouds heating up the surrounding dust.

### 5.3. Implications of our dust heating analysis

The non-negligible contribution from the older stellar population to the heating of dust emitting at IR/submm wavebands has been suggested in the past based on similar dust energy balance studies and/or optical extinction studies (Charlot & Fall 2000; Witt & Gordon 2000; Bianchi 2008; Cortese et al. 2008; Bendo et al. 2010, 2012a; Boquien et al. 2011; Popescu et al. 2011; De Looze et al. 2012b), and will have important implications for several extragalactic studies of the ISM properties in galaxies. More specifically, we caution the applicability of the total

infrared emission observed in galaxies as star formation rate tracer given that a significant fraction of the total infrared emission might not be entirely related to the recent star formation activity – at least for galaxies with low to moderate levels of star formation. The main driver of the total infrared energy budget in dusty starburst galaxies might, however, merely be attributed to the young stellar population.

Besides the questionable usability of  $L_{\text{TIR}}$  as reliable star formation rate diagnostic, the luminosities of far-infrared cooling lines are often compared to  $L_{\text{TIR}}$  (or  $L_{\text{FIR}}$ , 42.5–122 or 40–500  $\mu\text{m}$ ) to trace the photo-electric efficiency in the neutral gas phase, with the cooling lines tracing the gas heating and  $L_{\text{TIR}}$  probing the total dust heating. While the photo-electric effect on dust grains requires the energetic photons of young stars, the line-to-continuum ratios will be diminished by a possible contribution from the old stellar population to the continuum radiation (e.g., Parkin et al. 2013). Similarly, the total observed  $L_{\text{TIR}}$  of galaxies is often thought to entirely originate from star-forming regions, and used to constrain the input parameters of PDR models, neglecting a possible diffuse dust component heated by old stars but inefficient at heating the gas, which will result in an incorrect characterization of the PDR conditions as represented by the scaling factor of the interstellar radiation field,  $G_0$ , and the gas density,  $n_{\text{H}}$ .

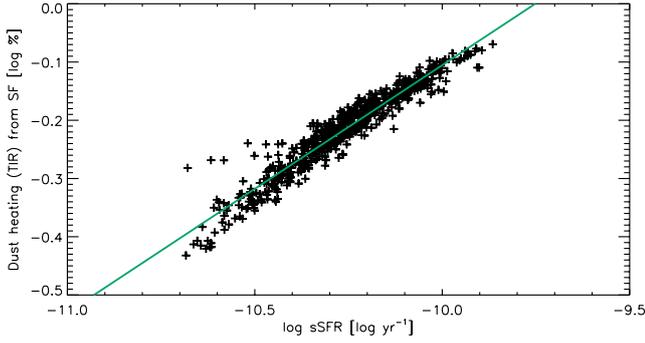
### 5.4. Dust heating prescriptions

Based on our radiative transfer model for M 51, we are able to quantify the contribution from different stellar populations to the dust heating at every location in 3D space in several infrared wavebands. To characterize the dust heating mechanisms, we look at the relation between the dust heating fraction of SF, i.e., the fraction of the energy that is provided by young stars ( $< 100$  Myr) to heat the dust component emitting at IR wavelengths, and the specific star formation rate (sSFR), i.e., the ratio of the current SFR divided by the stellar mass, which has been shown to be correlated to the FIR/submm colors of galaxies (e.g., Boselli et al. 2010, 2012; Boquien et al. 2011; Cortese et al. 2014). Since the sSFR traces the hardness of the UV radiation field, it is sensitive to the different heating sources in galaxies, and, thus, linked to the dust mass fraction heated by star formation (Ciesla et al. 2014).

We calculate the specific star formation rate from the output SFR map of the RT calculations, combined with the stellar mass derived from the 3.6  $\mu\text{m}$  SKIRT model image, adopting the conversion factor to stellar mass reported by Oliver et al. (2010). Maps of the sSFR and dust heating fraction have been rebinned to a grid with individual pixel sizes of 12.1", which corresponds to the resolution of the input images to the RT model of M 51, and, thus, the physical scale on which we can be confident of sampling independent regions. The image pixel size, however, does not guarantee that some of the stars heat the dust in neighboring pixels.

Figure 11 shows a tight relation between the sSFR and the fraction of TIR dust heating that can be attributed to star formation (i.e., the relative energy input of young stars to the dust heating). With a correlation coefficient<sup>10</sup> of  $\rho = 0.95$ , the importance of star formation in regulating the TIR emission in M 51 is clearly shown. A similar plot of the relative dust heating fraction

<sup>10</sup> The Spearman's rank correlation coefficient,  $\rho$ , is computed from the IDL procedure `r_correlate`. Values of  $\rho$  close to +1 and -1 are indicative of a strong correlation or anti-correlation, respectively, while values approaching 0 imply the absence of any correlation.



**Fig. 11.** Relation between the specific star formation rate and the relative contribution from young stars ( $\leq 100$  Myr) to the radiative dust heating responsible for the total-infrared emission in the RT model of M51. The best fitting line through the observed data points is indicated in green.

governed by old stars as a function of stellar mass (not shown here) does not reveal any particular trend with a correlation coefficient  $\rho = 0.05$ , suggesting that the dust heating in M51 is dominated by young stars as opposed to the old stellar population.

The best fitting line is fit to the data points in Fig. 11 to evaluate the mean trend between the sSFR and the TIR dust heating fraction, i.e.,  $\log \text{dust heating (SF)} = 0.42 \times \log \text{sSFR} + 4.14$ . Extrapolating the observed trend for M51, we deduce that the total infrared dust heating will be driven solely by young stellar sources for specific star formation rates  $\log \text{sSFR} \gtrsim -9.75$  (or  $\text{SFR} \gtrsim 0.2 M_{\odot} \text{ yr}^{-1}$ ,  $1.8 M_{\odot} \text{ yr}^{-1}$  and  $18 M_{\odot} \text{ yr}^{-1}$  for galaxies with stellar masses of  $10^9$ ,  $10^{10}$  and  $10^{11} M_{\odot}$ , respectively). More than half of the total infrared dust heating could be powered by the main interstellar radiation field for regions with  $\log \text{sSFR} \lesssim -10.46$  (or  $\text{SFR} \lesssim 0.03 M_{\odot} \text{ yr}^{-1}$ ,  $0.3 M_{\odot} \text{ yr}^{-1}$  and  $3 M_{\odot} \text{ yr}^{-1}$  for galaxies with stellar masses of  $10^9$ ,  $10^{10}$  and  $10^{11} M_{\odot}$ , respectively). Although the extrapolation of the dust heating mechanisms in M51 seems to guarantee that the dust heating in actively star-forming galaxies and starbursts is driven by the young stellar population, we need to be cautious to interpret the TIR emission from more quiescent galaxies with less vigorous star formation activity.

Since the dust heating conditions might differ from one galaxy to another, depending on metallicity, relative dust-star geometry, and other heating mechanisms unrelated to SF (e.g., AGN, shocks), we will, in future work, extend the dust heating analysis to a sample of star-forming galaxies representative for the local Universe to get a better handle on the link between the sSFR, and the importance of SF to the dust heating processes in galaxies. It will also be of specific interest to compare the dust heating fraction derived from the radiative transfer modeling to the fractions from diffuse versus clumpy media derived from multi-component dust models and SED fitting tools (e.g., Draine et al. 2007; Noll et al. 2009; da Cunha et al. 2010), and empirical color analyses (e.g., Bendo et al. 2010, 2012a, 2014; Galametz et al. 2010; Boquien et al. 2011; Hughes et al. 2014).

## 6. Conclusions

We present a panchromatic radiative transfer analysis of the grand-design spiral galaxy M51, that accounts for the absorption and scattering processes of stellar light by dust, and the thermal re-emission of dust in infrared wavebands. To reproduce the 3D asymmetric structures of stars and dust in M51, we rely on a multi-wavelength data set of observations to constrain the

geometrical distribution of stars and dust in our RT model. De-projection of these observed morphologies, in combination with the assumption of an exponential distribution in the vertical direction, defines the 3D structure of stars. The 3D dust distribution is characterized by the FUV dust attenuation map,  $A_{\text{FUV}}$ , derived from the TIR-to-FUV ratio, and following the prescriptions from Cortese et al. (2008), assuming an exponential  $z$  distribution.

Contrary to the dust energy balance problem encountered in RT models of edge-on spiral galaxies, the RT model for M51 is capable of simultaneously explaining the UV/optical dust extinction and FIR/submm thermal re-emission. We argue that the discrepancy in the dust energy balance in highly inclined systems is mainly related to the inability to probe the dust attenuating effects from UV/optical data, and the ignorance on the position of star-forming regions along the line-of-sight.

We benefit from the high-resolution 3D radiative transfer model, simulating the interaction of stellar light of various energies with the residing dust at every location in 3D space, to get insight in the different dust heating mechanisms that contribute to the dust emission in infrared wavebands. On global scales, we find that the young stars account for the heating of dust responsible for 63% of the total infrared emission (8–1000  $\mu\text{m}$ ), while 37% is governed through heating by the evolved stellar population. In individual wavebands, the contribution from young stars to the dust heating dominates at all infrared wavebands, but gradually decreases towards longer infrared and submillimeter wavebands, for which the old stellar population becomes a non-negligible source of heating. The tight correlation between the specific star formation rate (sSFR) and the TIR dust heating fraction from young stars confirms that young stars dominate the dust heating in M51. Upon extrapolation of the observed trends for M51, we conclude that the total infrared dust heating will be driven solely by young stellar sources for specific star formation rates  $\log \text{sSFR} \gtrsim -9.75$ , while more than half of the total infrared dust heating could be powered by the main interstellar radiation field for regions with  $\log \text{sSFR} \lesssim -10.46$ .

The first application of a new high-resolution 3D RT modeling technique reveals the importance of old stars contributing to the dust emission at all IR wavebands, which will affect the reliability of the total-infrared emission to trace the star formation rate in galaxies with moderate star formation activity. The 3D RT modeling technique presented in this work will be exploited in the future to derive possible variations in dust heating sources, dust geometry, and dust composition among different galaxy populations in the local Universe.

*Acknowledgements.* We thank the referee for useful comments and suggestions. We thank Richard Tuffs, Cristina Popescu, Simone Bianchi, Manolis Xilouris, and Giovanni Natale for fruitful discussions on the analysis of the dust heating fractions and the clumpiness of dust. We thank Robert C. Kennicutt Jr. for useful suggestions. We also thank Brent Groves for kindly providing us the dust masses associated with the emission spectra of the young HII regions with surrounding PDR envelopes. IDL is a postdoctoral researcher of the FWO-Vlaanderen (Belgium). M.B., J.F. and T.H. acknowledge the financial support of the Belgian Science Policy Office (BELSPO) through the PRODEX project “Herschel-PACS Guaranteed Time and Open Time Programs: Science Exploitation” (C90370). P.C. acknowledges the financial support of the Belgian Science Policy Office (BELSPO) through the CHARM framework (Contemporary physical challenges in Heliospheric and Astrophysical Models), a phase VII Interuniversity Attraction Pole (IAP) program. L.C. acknowledges support under the Australian Research Council’s Discovery Projects funding scheme (project number 130100664). PACS has been developed by a consortium of institutes led by MPE (Germany) and including UVIE (Austria); KU Leuven, CSL, IMEC (Belgium); CEA, LAM (France); MPIA (Germany); INAF/FISI/OAA/OAP/OAT, LENS, SISSA (Italy); IAC (Spain). This development has been supported by the funding agencies BMVIT (Austria), ESA-PRODEX (Belgium),

CEA/CNES (France), DLR (Germany), ASI/INAF (Italy), and CICYT/MCYT (Spain). SPIRE has been developed by a consortium of institutes led by Cardiff University (UK) and including Univ. Lethbridge (Canada); NAOC (China); CEA, LAM (France); IFSI, Univ. Padua (Italy); IAC (Spain); Stockholm Observatory (Sweden); Imperial College London, RAL, UCL-MSSL, UKATC, Univ. Sussex (UK); and Caltech, JPL, NHSC, Univ. Colorado (USA). This development has been supported by national funding agencies: CSA (Canada); NAOC (China); CEA, CNES, CNRS (France); ASI (Italy); MCINN (Spain); SNSB (Sweden); STFC and UKSA (UK); and NASA (USA).

## References

- Abazajian, K. N., Adelman-McCarthy, J. K., Agüeros, M. A., et al. 2009, *ApJS*, 182, 543
- Alton, P. B., Xilouris, E. M., Bianchi, S., Davies, J., & Kylafis, N. 2000, *A&A*, 356, 795
- Alton, P. B., Xilouris, E. M., Misiriotis, A., Dasyra, K. M., & Dumke, M. 2004, *A&A*, 425, 109
- Aniano, G., Draine, B. T., Gordon, K. D., & Sandstrom, K. 2011, *PASP*, 123, 1218
- Aniano, G., Draine, B. T., Calzetti, D., et al. 2012, *ApJ*, 756, 138
- Auld, R., Bianchi, S., Smith, M. W. L., et al. 2013, *MNRAS*, 428, 1880
- Baes, M., & Dejonghe, H. 2001, *MNRAS*, 326, 733
- Baes, M., Davies, J. I., Dejonghe, H., et al. 2003, *MNRAS*, 343, 1081
- Baes, M., Fritz, J., Gadotti, D. A., et al. 2010, *A&A*, 518, L39
- Baes, M., Verstappen, J., De Looze, I., et al. 2011, *ApJS*, 196, 22
- Bahcall, J. N., & Soneira, R. M. 1980, *ApJS*, 44, 73
- Bakes, E. L. O., & Tielens, A. G. G. M. 1994, *ApJ*, 427, 822
- Balm, P. 2012, *Astronomical Data Analysis Software and Systems XXI*, 461, 733
- Balog, Z., Müller, T., Nielbock, M., et al. 2013, *Experiment. Astron.*, 38
- Beckman, J. E., Peletier, R. F., Knapen, J. H., Corradi, R. L. M., & Gentet, L. J. 1996, *ApJ*, 467, 175
- Bendo, G. J., Draine, B. T., Engelbracht, C. W., et al. 2008, *MNRAS*, 389, 629
- Bendo, G. J., Wilson, C. D., Pohlen, M., et al. 2010, *A&A*, 518, L65
- Bendo, G. J., Boselli, A., Dariush, A., et al. 2012a, *MNRAS*, 419, 1833
- Bendo, G. J., Galliano, F., & Madden, S. C. 2012b, *MNRAS*, 423, 197
- Bendo, G. J., Griffin, M. J., Bock, J. J., et al. 2013, *MNRAS*, 433, 3062
- Bendo, G. J., Baes, M., Bianchi, S., et al. 2014, *MNRAS*, in press  
[arXiv:1409.1815]
- Bernard, J.-P., Reach, W. T., Paradis, D., et al. 2008, *AJ*, 136, 919
- Bianchi, L., Madore, B., Thilker, D., et al. 2003a, *The Local Group as an Astrophysical Laboratory*, 10
- Bianchi, L., Madore, B., Thilker, D., Gil de Paz, A., & GALEX Science Team 2003b, *BAAS*, 35, 1354
- Bianchi, L., Thilker, D. A., Burgarella, D., et al. 2005, *ApJ*, 619, L71
- Bianchi, S. 2007, *A&A*, 471, 765
- Bianchi, S. 2008, *A&A*, 490, 461
- Bianchi, S., Ferrara, A., Davies, J. I., & Alton, P. B. 2000, *MNRAS*, 311, 601
- Blitz, L., Fukui, Y., Kawamura, A., et al. 2007, *Protostars and Planets V*, 81
- Boquien, M., Calzetti, D., Combes, F., et al. 2011, *AJ*, 142, 111
- Boselli, A., & Gavazzi, G. 2002, *A&A*, 386, 124
- Boselli, A., Lequeux, J., & Gavazzi, G. 2004, *A&A*, 428, 409
- Boselli, A., Boissier, S., Cortese, L., et al. 2009, *ApJ*, 706, 1527
- Boselli, A., Ciesla, L., Buat, V., et al. 2010, *A&A*, 518, L61
- Boselli, A., Ciesla, L., Cortese, L., et al. 2012, *A&A*, 540, A54
- Boulanger, F., Beichman, C., Desert, F. X., et al. 1988, *ApJ*, 332, 328
- Bresolin, F., Garnett, D. R., & Kennicutt, R. C., Jr. 2004, *ApJ*, 615, 228
- Buat, V., & Xu, C. 1996, *A&A*, 306, 61
- Calzetti, D., Kinney, A. L., & Storchi-Bergmann, T. 1994, *ApJ*, 429, 582
- Calzetti, D., Kennicutt, R. C., Jr., Bianchi, L., et al. 2005, *ApJ*, 633, 871
- Calzetti, D., Kennicutt, R. C., Engelbracht, C. W., et al. 2007, *ApJ*, 666, 870
- Camps, D., Baes, M., De Geyter, G., Saitly, W., & De Looze, I. 2013a, *ASP Conf. Ser.*, 474, 53
- Camps, P., Baes, M., & Saitly, W. 2013b, *A&A*, 560, A35
- Cardelli, J. A., Clayton, G. C., & Mathis, J. S. 1989, *ApJ*, 345, 245
- Cesarsky, D., Lequeux, J., Abergel, A., et al. 1996, *A&A*, 315, L305
- Chabrier, G. 2003, *PASP*, 115, 763
- Ciesla, L., Boselli, A., Smith, M. W. L., et al. 2012, *A&A*, 543, A161
- Ciesla, L., Boquien, M., Boselli, A., et al. 2014, *A&A*, 565, A128
- Charlot, S., & Fall, S. M. 2000, *ApJ*, 539, 718
- Combes, F., Boquien, M., Kramer, C., et al. 2012, *A&A*, 539, A67
- Cortese, L., Boselli, A., Franzetti, P., et al. 2008, *MNRAS*, 386, 1157
- Cortese, L., Ciesla, L., Boselli, A., et al. 2012, *A&A*, 540, A52
- Cortese, L., Fritz, J., Bianchi, S., et al. 2014, *MNRAS*, 440, 942
- da Cunha, E., Eminian, C., Charlot, S., & Blaizot, J. 2010, *MNRAS*, 403, 1894
- Dalcanton, J. J., Yoachim, P., & Bernstein, R. A. 2004, *ApJ*, 608, 189
- Dale, D. A., Aniano, G., Engelbracht, C. W., et al. 2012, *ApJ*, 745, 95
- Dasyra, K. M., Xilouris, E. M., Misiriotis, A., & Kylafis, N. D. 2005, *A&A*, 437, 447
- De Geyter, G., Baes, M., Fritz, J., & Camps, P. 2013, *A&A*, 550, A74
- De Geyter, G., Baes, M., Camps, P., et al. 2014, *MNRAS*, 441, 869
- de Grijs, M. H. K., Miley, G. K., Lub, J., & de Jong, T. 1985, *Nature*, 314, 240
- de Grijs, R., & Peletier, R. F. 1997, *A&A*, 320, L21
- De Looze, I., Baes, M., Fritz, J., & Verstappen, J. 2012a, *MNRAS*, 419, 8
- De Looze, I., Baes, M., Bendo, G. J., et al. 2012b, *MNRAS*, 427, 2797
- Devereux, N. A., & Hameed, S. 1997, *AJ*, 113, 599
- Devereux, N. A., & Young, J. S. 1990, *ApJ*, 350, L25
- Disney, M., Davies, J., & Phillipps, S. 1989, *MNRAS*, 239, 939
- Dobbs, C. L., Theis, C., Pringle, J. E., & Bate, M. R. 2010, *MNRAS*, 403, 625
- Draine, B. 1989, *Interstellar Dust*, 135, 313
- Draine, B. T., & Li, A. 2007, *ApJ*, 657, 810
- Draine, B. T., Dale, D. A., Bendo, G., et al. 2007, *ApJ*, 663, 866
- Driver, S. P., Popescu, C. C., Tuffs, R. J., et al. 2007, *MNRAS*, 379, 1022
- Engelbracht, C. W., Blaylock, M., Su, K. Y. L., et al. 2007, *PASP*, 119, 994
- Fazio, G. G., Hora, J. L., Allen, L. E., et al. 2004, *ApJS*, 154, 10
- Feldmeier, J. J., Ciardullo, R., & Jacoby, G. H. 1997, *ApJ*, 479, 231
- Ferrara, A., Bianchi, S., Cimatti, A., & Giovanardi, C. 1999, *ApJS*, 123, 437
- Fischera, J., & Dopita, M. 2005, *ApJ*, 619, 340
- Fitzpatrick, E. L. 1999, *PASP*, 111, 63
- Fitzpatrick, E. L., & Massa, D. 2007, *ApJ*, 663, 320
- Gadotti, D. A., Baes, M., & Falony, S. 2010, *MNRAS*, 403, 2053
- Galametz, M., Madden, S. C., Galliano, F., et al. 2010, *A&A*, 518, L55
- Galametz, M., Kennicutt, R. C., Albrecht, M., et al. 2012, *MNRAS*, 425, 763
- Galametz, M., Kennicutt, R. C., Calzetti, D., et al. 2013, *MNRAS*, 431, 1956
- Gil de Paz, A., Madore, B. F., Boissier, S., & GALEX Science Team 2004, *BAAS*, 36, 1410
- Gil de Paz, A., Boissier, S., Madore, B. F., et al. 2007, *ApJS*, 173, 185
- Gomez, H. L., Baes, M., Cortese, L., et al. 2010, *A&A*, 518, L45
- Gordon, K. D., Calzetti, D., & Witt, A. N. 1997, *ApJ*, 487, 625
- Gordon, K. D., Clayton, G. C., Witt, A. N., & Misselt, K. A. 2000, *ApJ*, 533, 236
- Gordon, K. D., Clayton, G. C., Misselt, K. A., Landolt, A. U., & Wolff, M. J. 2003, *ApJ*, 594, 279
- Gordon, K. D., Engelbracht, C. W., Rieke, G. H., et al. 2008, *ApJ*, 682, 336
- Gordon, K. D., Roman-Duval, J., Bot, C., et al. 2014, *ApJ*, submitted  
[arXiv:1406.6066]
- Granato, G. L., Lacey, C. G., Silva, L., et al. 2000, *ApJ*, 542, 710
- Groves, B. 2004, Ph.D. Thesis, Australian National University
- Groves, B., Dopita, M. A., Sutherland, R. S., et al. 2008, *ApJS*, 176, 438
- Hao, C.-N., Kennicutt, R. C., Johnson, B. D., et al. 2011, *ApJ*, 741, 124
- Helou, G. 1986, *ApJ*, 311, L33
- Helou, G., Roussel, H., Appleton, P., et al. 2004, *ApJS*, 154, 253
- Hollenbach, D., & Salpeter, E. E. 1971, *ApJ*, 163, 155
- Holwerda, B. W., Bianchi, S., Böker, T., et al. 2012, *A&A*, 541, L5
- Houk, J. C., & Savage, B. D. 1999, *AJ*, 117, 2077
- Hughes, T. M., Baes, M., Fritz, J., et al. 2014, *A&A*, 565, A4
- Irwin, J. A., & Madden, S. C. 2006, *A&A*, 445, 123
- Irwin, J. A., Kennedy, H., Parkin, T., & Madden, S. 2007, *A&A*, 474, 461
- Jansen, R. A., Knapen, J. H., Beckman, J. E., Peletier, R. F., & Hes, R. 1994, *MNRAS*, 270, 373
- Jarrett, T. H., Chester, T., Cutri, R., Schneider, S. E., & Huchra, J. P. 2003, *AJ*, 125, 525
- Jarrett, T. H., Cohen, M., Masci, F., et al. 2011, *ApJ*, 735, 112
- Jarrett, T. H., Masci, F., Tsai, C. W., et al. 2013, *AJ*, 145, 6
- Jonsson, P., Groves, B. A., & Cox, T. J. 2010, *MNRAS*, 403, 17
- Kaleida, C., & Scowen, P. A. 2010, *AJ*, 140, 379
- Kennicutt, R. C., Jr., Armus, L., Bendo, G., et al. 2003, *PASP*, 115, 928
- Kamphuis, P., Holwerda, B. W., Allen, R. J., Peletier, R. F., & van der Kruit, P. C. 2007, *A&A*, 471, L1
- Kennicutt, R. C., Jr., Hao, C.-N., Calzetti, D., et al. 2009, *ApJ*, 703, 1672
- Kewley, L. J., Geller, M. J., Jansen, R. A., & Dopita, M. A. 2002, *AJ*, 124, 3135
- Kirkpatrick, A., Calzetti, D., Kennicutt, R., et al. 2014, *ApJ*, 789, 130
- Knapp, G. R., Gunn, J. E., & Wynn-Williams, C. G. 1992, *ApJ*, 399, 76
- Kregel, M., van der Kruit, P. C., & de Grijs, R. 2002, *MNRAS*, 334, 646
- Kreckel, K., Groves, B., Schinnerer, E., et al. 2013, *ApJ*, 771, 62
- Kroupa, P. 2001, *MNRAS*, 322, 231
- Kroupa, P. 2002, *Science*, 295, 82
- Kurosawa, R., & Hillier, D. J. 2001, *A&A*, 379, 336
- Law, K.-H., Gordon, K. D., & Misselt, K. A. 2011, *ApJ*, 738, 124
- Leitherer, C., Schaerer, D., Goldader, J. D., et al. 1999, *ApJS*, 123, 3
- Li, A., & Draine, B. T. 2002, *ApJ*, 572, 232
- Liu, G., Calzetti, D., Hong, S., et al. 2013, *ApJ*, 778, L41
- Lu, N., Bendo, G. J., Boselli, A., et al. 2014, *ApJ*, submitted

- MacLachlan, J. M., Matthews, L. D., Wood, K., & Gallagher, J. S. 2011, *ApJ*, 741, 6
- Maller, A. H., Berlind, A. A., Blanton, M. R., & Hogg, D. W. 2009, *ApJ*, 691, 394
- Maraston, C. 1998, *MNRAS*, 300, 872
- Maraston, C. 2005, *MNRAS*, 362, 799
- Martin, D. C., Fanson, J., Schiminovich, D., et al. 2005, *ApJ*, 619, L1
- Mazzei, P., & de Zotti, G. 1994, *ApJ*, 426, 97
- Mentuch Cooper, E., Wilson, C. D., Foyle, K., et al. 2012, *ApJ*, 755, 165
- Meurer, G. R., Heckman, T. M., & Calzetti, D. 1999, *ApJ*, 521, 64
- Misiriotis, A., Popescu, C. C., Tuffs, R., & Kylafis, N. D. 2001, *A&A*, 372, 775
- Miura, H., & Nakamoto, T. 2006, *ApJ*, 651, 1272
- Montalto, M., Seitz, S., Riffeser, A., et al. 2009, *A&A*, 507, 283
- Morrissey, P., Conrow, T., Barlow, T. A., et al. 2007, *ApJS*, 173, 682
- Moustakas, J., Kennicutt, R. C., Jr., Tremonti, C. A., et al. 2010, *ApJS*, 190, 233
- Narayan, C. A., & Jog, C. J. 2002, *A&A*, 390, L35
- Niccolini, G., & Alcolea, J. 2006, *A&A*, 456, 1
- Noeske, K. G., Weiner, B. J., Faber, S. M., et al. 2007, *ApJ*, 660, L43
- Noll, S., Burgarella, D., Giovannoli, E., et al. 2009, *A&A*, 507, 1793
- Nguyen, H. T., Schulz, B., Levenson, L., et al. 2010, *A&A*, 518, L5
- Oliver, S., Frost, M., Farrah, D., et al. 2010, *MNRAS*, 405, 2279
- Padmanabhan, N., Schlegel, D. J., Finkbeiner, D. P., et al. 2008, *ApJ*, 674, 1217
- Paradis, D., Reach, W. T., Bernard, J.-P., et al. 2009, *AJ*, 138, 196
- Parkin, T. J., Wilson, C. D., Schirm, M. R. P., et al. 2013, *ApJ*, 776, 65
- Peng, C. Y., Ho, L. C., Impey, C. D., & Rix, H.-W. 2002, *AJ*, 124, 266
- Peng, C. Y., Ho, L. C., Impey, C. D., & Rix, H.-W. 2010, *AJ*, 139, 2097
- Pierini, D., Gordon, K. D., Witt, A. N., & Madsen, G. J. 2004, *ApJ*, 617, 1022
- Pilbratt, G. L., Riedinger, J. R., Passvogel, T., et al. 2010, *A&A*, 518, L1
- Popescu, C. C., Misiriotis, A., Kylafis, N. D., Tuffs, R. J., & Fischera, J. 2000, *A&A*, 362, 138
- Popescu, C. C., Tuffs, R. J., Dopita, M. A., et al. 2011, *A&A*, 527, A109
- Prescott, M. K. M., Kennicutt, R. C., Jr., Bendo, G. J., et al. 2007, *ApJ*, 668, 182
- Reach, W. T., Megeath, S. T., Cohen, M., et al. 2005, *PASP*, 117, 978
- Rémy-Ruyer, A., Madden, S. C., Galliano, F., et al. 2013, *A&A*, 557, A95
- Rieke, G. H., Young, E. T., Engelbracht, C. W., et al. 2004, *ApJS*, 154, 25
- Robitaille, T. P., Churchwell, E., Benjamin, R. A., et al. 2012, *A&A*, 545, A39
- Rosolowsky, E., Engargiola, G., Plambeck, R., & Blitz, L. 2003, *ApJ*, 599, 258
- Roussel, H. 2013, *PASP*, 125, 1126
- Saftly, W., Camps, P., Baes, M., et al. 2013, *A&A*, 554, A10
- Saftly, W., Baes, M., & Camps, P. 2014, *A&A*, 561, A77
- Salo, H., & Laurikainen, E. 2000, *MNRAS*, 319, 377
- Salpeter, E. E. 1955, *ApJ*, 121, 161
- Sauvage, M., & Thuan, T. X. 1994, *ApJ*, 429, 153
- Sauvage, M., Vigroux, L., & Thuan, T. X. 1990, *A&A*, 237, 296
- Schechtman-Rook, A., & Bershad, M. A. 2013, *ApJ*, 773, 45
- Schechtman-Rook, A., Bershad, M. A., & Wood, K. 2012, *ApJ*, 746, 70
- Seth, A. C., Dalcanton, J. J., & de Jong, R. S. 2005, *AJ*, 130, 1574
- Schlafly, E. F., & Finkbeiner, D. P. 2011, *ApJ*, 737, 103
- Schlegel, D. J., Finkbeiner, D. P., & Davis, M. 1998, *ApJ*, 500, 525
- Shetty, R., Vogel, S. N., Ostriker, E. C., & Teuben, P. J. 2007, *ApJ*, 665, 1138
- Silva, L., Granato, G. L., Bressan, A., & Danese, L. 1998, *ApJ*, 509, 103
- Skrutskie, M. F., Cutri, R. M., Stiening, R., et al. 2006, *AJ*, 131, 1163
- Smith, M. W. L., Eales, S. A., Gomez, H. L., et al. 2012, *ApJ*, 756, 40
- Stalevski, M., Fritz, J., Baes, M., Nakos, T., & Popović, L. Č. 2012, *MNRAS*, 420, 2756
- Steinacker, J., Baes, M., & Gordon, K. D. 2013, *ARA&A*, 51, 63
- Tabatabaei, F. S., Braine, J., Xilouris, E. M., et al. 2014, *A&A*, 561, A95
- Thompson, T. W. J., Howk, J. C., & Savage, B. D. 2004, *AJ*, 128, 662
- Tuffs, R. J., Popescu, C. C., Völk, H. J., Kylafis, N. D., & Dopita, M. A. 2004, *A&A*, 419, 821
- Tully, R. B. 1974, *ApJS*, 27, 437
- Verstappen, J., Fritz, J., Baes, M., et al. 2013, *A&A*, 556, A54
- Viaene, S., Fritz, J., Baes, M., et al. 2014, *A&A*, 567, A71
- Vidal, E., & Baes, M. 2007, *Baltic Astron.*, 16, 101
- Vlahakis, C., van der Werf, P., Israel, F. P., & Tilanus, R. P. J. 2013, *MNRAS*, 433, 1837
- Voronoi, G. 1908, *J. reine und angewandte Mathematik*, 134, 1980
- Whaley, C. H., Irwin, J. A., Madden, S. C., Galliano, F., & Bendo, G. J. 2009, *MNRAS*, 395, 97
- Wild, V., Charlot, S., Brinchmann, J., et al. 2011, *MNRAS*, 417, 1760
- Witt, A. N., & Gordon, K. D. 2000, *ApJ*, 528, 799
- Wolfire, M. G., Hollenbach, D., McKee, C. F., Tielens, A. G. G. M., & Bakes, E. L. O. 1995, *ApJ*, 443, 152
- Wright, E. L., Eisenhardt, P. R. M., Mainzer, A. K., et al. 2010, *AJ*, 140, 1868
- Wu, H., Zhu, Y.-N., Cao, C., & Qin, B. 2007, *ApJ*, 668, 87
- Xilouris, E. M., Kylafis, N. D., Papamastorakis, J., Paleologou, E. V., & Haerendel, G. 1997, *A&A*, 325, 135
- Xilouris, E. M., Alton, P. B., Davies, J. I., et al. 1998, *A&A*, 331, 894
- Xilouris, E. M., Byun, Y. I., Kylafis, N. D., Paleologou, E. V., & Papamastorakis, J. 1999, *A&A*, 344, 868
- Xilouris, E. M., Tabatabaei, F. S., Boquien, M., et al. 2012, *A&A*, 543, A74
- Xu, C., & Helou, G. 1996, *ApJ*, 456, 163
- Zubko, V., Dwek, E., & Arendt, R. G. 2004, *ApJS*, 152, 211

## Appendix A: The effect of model degeneracies

### A.1. Relative dust-stellar scale heights

In Sect. 3.2.1, we describe our model assumptions for the relative star-dust geometry in M51 based on observations of edge-on galaxies for which the stellar and dust geometries can be vertically resolved. Since the relative geometry of stars and dust has been shown to affect the derived attenuation (e.g., Disney et al. 1989; Ferrara et al. 1999), we aim to analyze the model implications regarding our assumptions on the relative dust-star geometry. Figure A.1 shows the model SEDs (top) and relative difference between the model SEDs (bottom) for a RT model of M51 with dust scale height  $h_{z,d} = 150, 225$  and  $450$  pc, or a dust-to-stellar scale height of  $\frac{1}{3}, \frac{1}{2}$  and  $1$ , respectively.

The difference between the three RT models with different dust scale heights is limited to the UV and far-infrared wavelength domain, while the optical wavebands do not reveal a clear dissimilarity between the observed SEDs. The FUV emission decreases for RT models with increasing dust scale heights, i.e., dust scale heights that approach the vertical scale height of stars, suggesting that more FUV emission is absorbed for dust components distributed in a thicker disk. Comparing the three models, the UV absorption increases from  $\sim 75\%$  to  $78\%$  and  $81\%$  for dust scale heights of  $h_{z,d} = 150, 225$  and  $450$  pc, respectively. The scattered fraction of UV light, on the other hand, slightly decreases from  $12\%$  to  $11\%$ . The total amount of absorbed stellar energy (i.e.,  $10\%$  of the total stellar bolometric luminosity) remains invariable for the three RT models, which confirms earlier results reported by Beckman et al. (1996) and Tuffs et al. (2004) that the relative dust-star geometry does not severely affect the attenuation (except in UV wavebands). The increased absorption of UV light for RT models with longer dust scale heights translates in brighter PAH features and increased mid-infrared emission. The differences between the three RT models with variation in dust scale heights at infrared wavebands are limited to  $\sim 10\%$ , and, therefore, will not have a significant effect on the RT model parameters.

Figure A.2 shows the observed (first panel) and model images of the FUV dust attenuation in M51 for RT models with dust scale height  $h_{z,d} = 150$  pc (second panel),  $h_{z,d} = 225$  pc (third panel), and  $h_{z,d} = 450$  pc (last panel). From the FUV attenuation maps, we can derive that the absorbed FUV fraction is indeed higher/lower for larger/smaller dust scale heights. Since an adjustment of the dust mass would be necessary to reproduce the observed FUV attenuation map for RT models with dust scale heights  $h_{z,d} = 150$  and  $450$  pc, respectively, we will not longer reproduce the observed FIR/submm observations of M51. We, therefore, argue that the assumed relative dust-to-stellar scale height  $h_{z,d}/h_{z,*} = 0.5$  provides a good representation of the true vertical distribution of stars and dust in M51. We do need to be cautious regarding uncertainties on the assumed dust model parameters, i.e., deviations from the assumed dust opacity  $\kappa$ , and its variation with wavelength (e.g., Tabatabaei et al. 2014; Gordon et al. 2014), which will reflect in a similar uncertainty on the dust mass determination, FUV attenuation, and dust emission in IR/submm wavebands.

### A.2. Clumpiness of dust

The RT model of M51 does not include any effects of clumpy dust in addition to the 2D asymmetry derived from the FUV attenuation map to describe the dust geometry. To analyze the effects of dust clumping on the RT model results for M51, we

introduce a two-phase geometry in our radiative transfer code SKIRT characterized by a diffuse and clumpy phase. The two-phase component is conditioned through three parameters: the number of clumps, the size of individual clumps, and the clump fraction, i.e., the fraction of the dust mass that is locked up in clumps with the remaining mass distributed in the diffuse medium. Rather than using a density ratio between the clumps and diffuse medium, we model the clumpy regions using a Cubic Spline Smoothing kernel  $W(r, h)$  to describe the density of the two-phase medium:

$$\rho(r) = (1 - f_{cl})\rho_{orig}(r) + \frac{f_{cl}}{N} \sum_{i=1}^N W(r - r_i, h), \quad (\text{A.1})$$

with  $N$  = the number of clumps,  $\rho_{orig}$  = the original density,  $f_{cl}$  = the fraction of mass locked up in clumps,  $r$  = the 3D position in the dust grid, and  $h$  = the scale radius of a single clump. The clumps with their centers at position  $r_i$  are distributed stochastically throughout the diffuse medium relying on the probability density distribution function representative of the dust component.

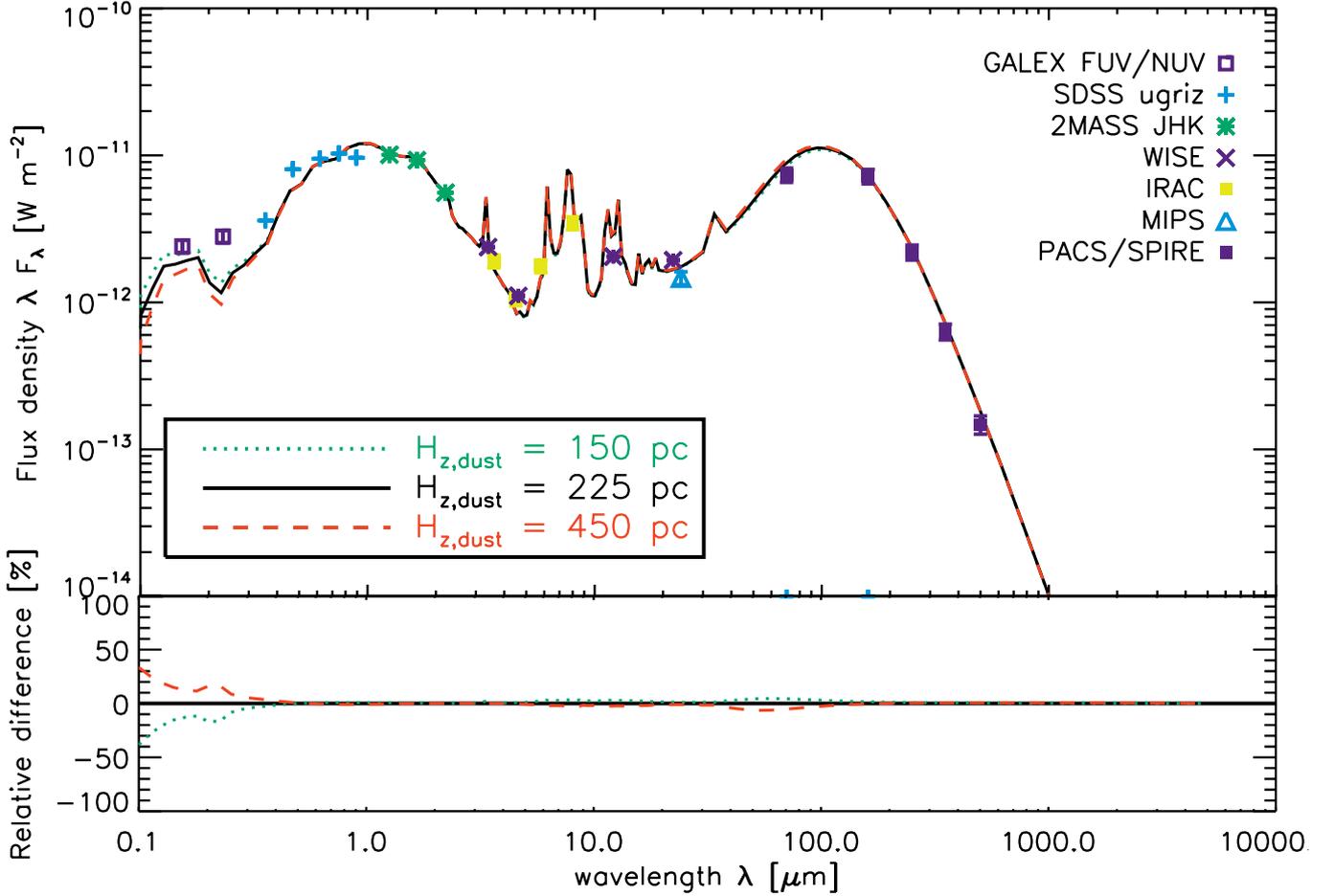
Similar to the two-phase clumpy medium applied by Bianchi (2008), the clumps in our model are assumed to have sizes similar to the giant molecular clouds observed in our Milky Way. Based on the largest cloud mass of  $10^{6.5} M_{\odot}$  (Reach et al. 2005) and assuming a typical surface density  $\Sigma = 100 M_{\odot} \text{pc}^{-2}$  observed in Local Group galaxies (Rosolowsky et al. 2003; Blitz et al. 2007), we derive a corresponding radius of  $R \sim 100$  pc. The number of dust clumps of a single mass depends on the gas-to-dust ratio (G/D) and the clump fraction  $f_{cl}$  through:

$$(G/D)^{-1} \sum_{i=1}^{N_c} M_i = f_{cl} M_{dust}. \quad (\text{A.2})$$

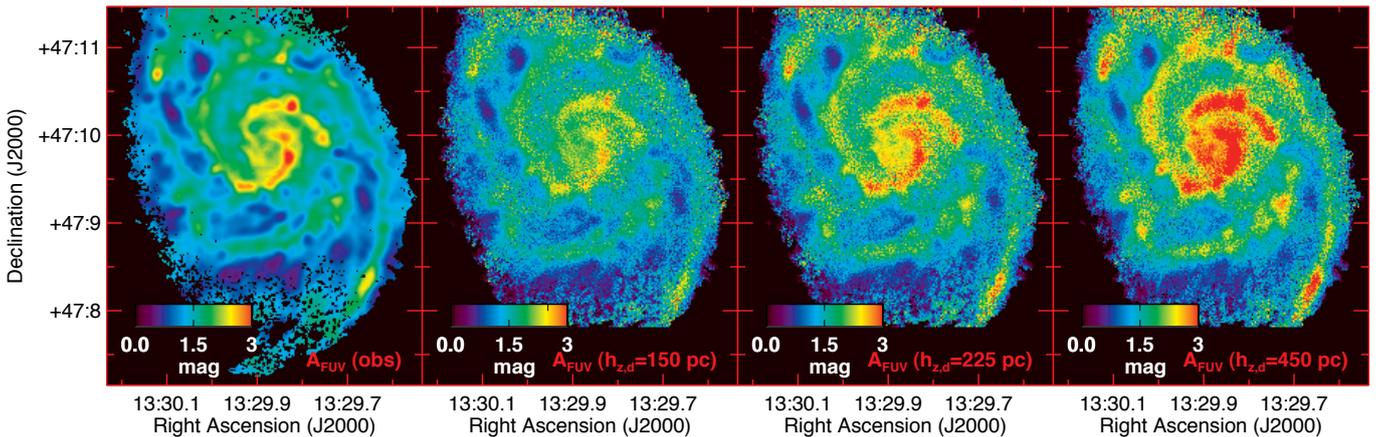
The gas-to-dust ratio is assumed to be around  $\sim 105$ , following the calculations from Mentuch Cooper et al. (2012).

Figure A.3 shows the model SEDs (top), and relative difference between the SEDs (bottom) for radiative transfer simulations with a two-phase dust component consisting of a smooth medium with superimposed a fraction of clumps representing  $0\%$  (equivalent to no clumps),  $20\%$ , and  $50\%$  of the total dust mass, indicated in black solid, green dotted, and red dashed lines, respectively. With increasing clump fractions, the dusty medium becomes more transparent, resulting in increased UV and optical emission of stars in the RT model. The total amount of stellar light that is absorbed by dust decreases from  $10\%$  to  $7\%$  going from a smooth to clumpy medium with  $f_{cl} = 0.5$ , while the amount of stellar photons that is scattered by the dust medium decreases from  $18\%$  to  $15\%$ . The diminished absorption rate is most apparent in the UV wavebands with UV absorption fractions decreasing from  $75\%$  to  $55\%$  going from a smooth to clumpy media with  $f_{cl} = 0.5$ . The UV scattering, on the other hand, increases from about  $10\%$  to  $15\text{--}20\%$  from a smooth to clumpy medium with half of the dust mass distributed in clumps. The clumping fraction of dust has, thus, a direct impact on the UV absorption and scattering rates, which also affects the shape of the  $2175 \text{ \AA}$  feature (see bottom panel of Fig. A.3). In addition to variations in the grain composition and/or size distribution, the clumping of dust might play an important role in explaining the shape of the dust bump at  $2175 \text{ \AA}$ .

The transparency of the ISM in clumpy media reduces the absorbed stellar energy, which is reflected in a decrease of the infrared and (sub)millimeter emission with about  $\sim 10\text{--}15\%$  long



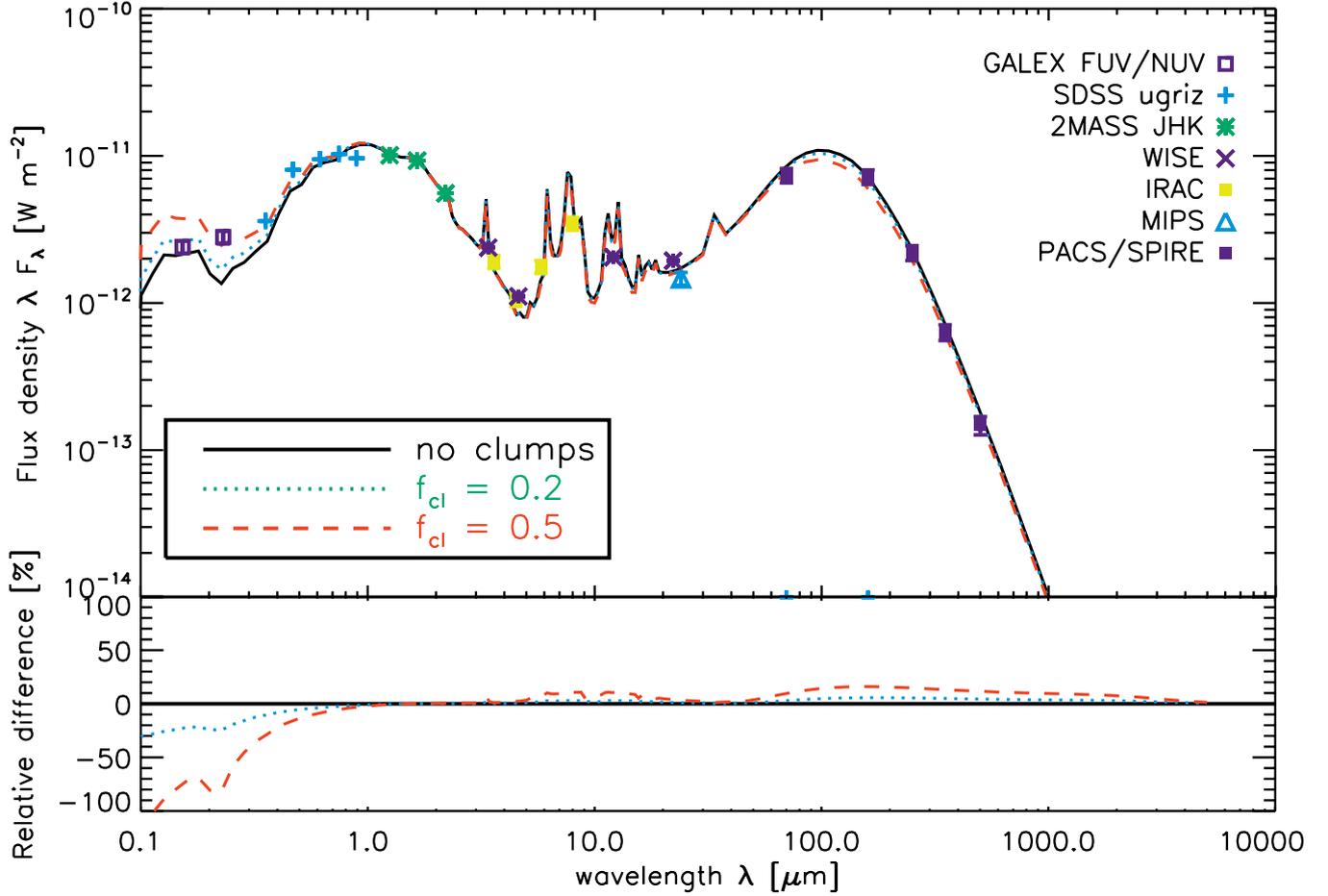
**Fig. A.1.** Effect on the model SED (*top*) derived from the RT simulations for M51 assuming a dust scale height  $h_{z,d} = 150, 225$  and  $450$  pc or, thus, a dust-to-stellar scale height of  $\frac{1}{3}, \frac{1}{2}$  and  $1$ , indicated as a black solid, green dotted and red dashed line, respectively. The bottom panel shows the relative SEDs for RT models with dust scale heights of  $h_{z,d} = 150$  pc (green dotted line) and  $h_{z,d} = 450$  pc (red dashed line) compared to the RT model with  $h_{z,d} = 225$  pc, which was used for the analysis in this paper.



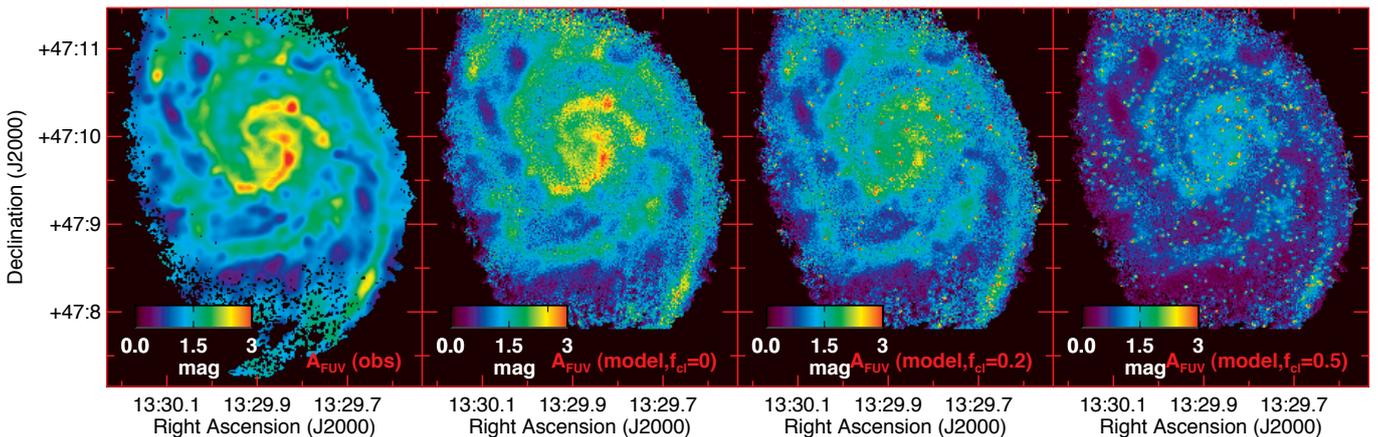
**Fig. A.2.** FUV attenuation map for M51 as derived from observations (*left*), for the RT model exploited in this paper with dust scale height  $h_{z,d} = 150$  pc (*second panel*),  $h_{z,d} = 225$  pc (*third panel*) and  $h_{z,d} = 450$  pc (*last panel*).

wards of  $60\mu\text{m}$ . Similar effects of dust clumping have been shown in RT simulations presented by Bianchi et al. (2000); Pierini et al. (2004); Baes et al. (2011). The decreased dust emission could be compensated for by increasing the dust mass in the RT model, but would still underestimate the FUV attenuation derived for M51 by  $\sim 40\%$  (see Fig. A.4). We, therefore, argue that the 2D structure from the FUV attenuation map, constructed based on FUV and infrared observations, provides an

adequate description of the 3D asymmetry in the distribution of dust in M51. The underestimation of the FUV attenuation upon including a two-phase dust component with clumps suggests that M51 likely does not contain many compact dust clouds without embedded sources. This result for M51 differs from the substantial fractions of compact dust clumps needed to explain the FIR/submm emission observed in NGC 891 ( $f_{cl} = 0.5$ , Bianchi 2008), NGC 4594 ( $f_{cl} = 0.75$ , De Looze et al. 2012a) and



**Fig. A.3.** Effect on the model SED (*top*) derived from the RT simulations for M51 assuming a two-phase dust component with a clump fraction  $f_{cl}$  of 0, 0.2 and 0.5, representative for the relative fraction of dust mass distributed in clumps, indicated as a black solid, green dotted and red dashed line, respectively. The *bottom panel* shows the relative SEDs for RT models with clump fractions  $f_{cl} = 0.2$  (green dotted line) and  $f_{cl} = 0.5$  (red dashed line) compared to the RT model with a smooth dust component, which was used for the analysis in this paper.



**Fig. A.4.** FUV attenuation map for M51 as derived from observations (*left*), for the RT model exploited in this paper with smooth dust component (*second panel*) and the RT model with a two-phase dust component with clump fractions  $f_{cl} = 0.2$  (*third panel*) and  $f_{cl} = 0.5$  (*last panel*).

NGC 4565 ( $f_{cl} = 0.67$ , De Looze et al. 2012b). A first difference between these galaxies and M51 is their edge-on view, which does not allow to constrain the 3D structure of stars and dust as we did for M51, and might, therefore, introduce uncertainties into the model. A second explanation might be the difference in galaxy type between the former galaxies and M51. In late-type galaxies such as M51 (SAbc), we expect to find star-forming

regions with embedding young stars, while the early-type galaxies (Sa-Sb) likely have lower levels of star formation activity, with several compact gaseous clouds where star formation has not (yet) been ignited. The study of a bigger sample of face-on galaxies will be necessary to disentangle between the effects of inclination and galaxy type in explaining the need for compact quiescent dust clouds.